The Magnetosphere: A Focus on Convection and Substorms

Christine Gabrielse

The Aerospace Corporation

Email: christine.gabrielse@aero.org

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Overview

So much cool stuff, so little time…

- How does one cover the magnetosphere in two 50 minute lectures?
  - *Focus: Convection and Substorms*

- Even so…There's a lot more material on convection…and a lot more material on substorms…I tried to look at each through the lens of the other

- Goal:
  - To make convection and substorm concepts accessible to graduate student audience.
  - Use some recent studies as examples to demonstrate current works in the field.

- Introductory material: Definitions and Broad Overview
  - *Dungey Cycle*
  - *Particle motion lite*
- Convection on large scales
- Convection on mesoscales
- Convection and Substorms

This presentation is being recorded. – christine.gabrielse@aero.org

Convection: The Dungey Cycle

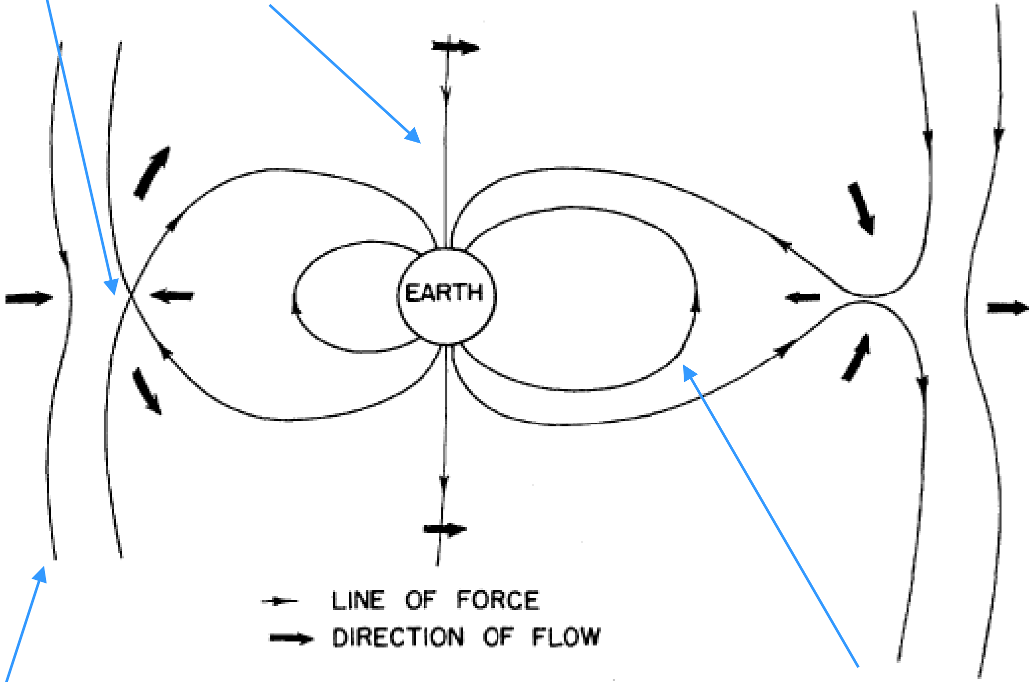
Magnetic Flux Transport: In a steady state, all rates must be equal

Dungey, PRL, (1961): <2 pages, 1914-4032 citations and growing

$$Magnetic\ flux = \Phi = BA\cos(\Theta)$$

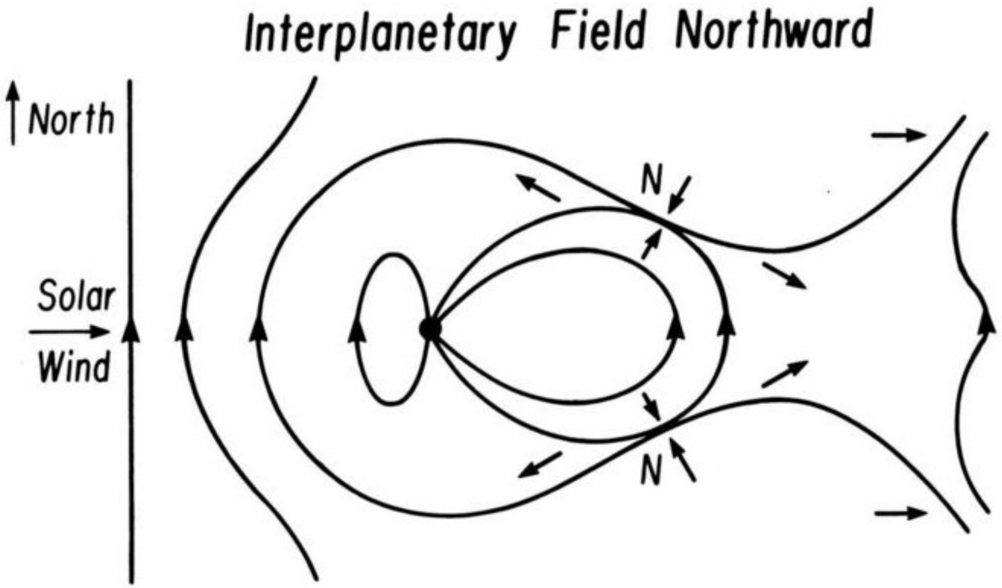
Dayside reconnection

Open magnetic field lines: Earth magnetic field lines connected to IMF.



Interplanetary Magnetic Field (IMF): Sun's magnetic field embedded in the solar wind. (Southward)

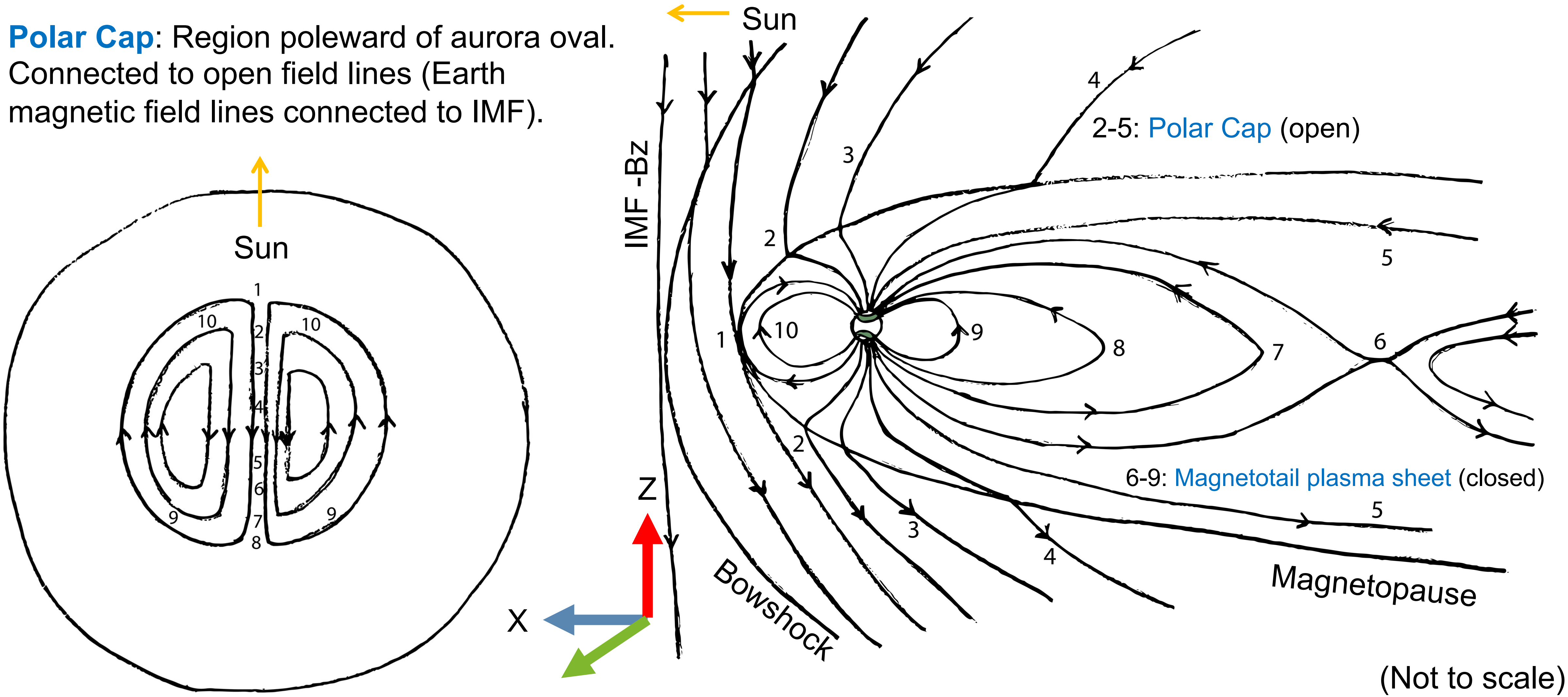
Closed magnetic field lines: Earth magnetic field lines connected to only Earth.



Credit: Dungey (1963?)

Large-Scale Convection: Ionosphere

Polar Cap: Region poleward of aurora oval. Connected to open field lines (Earth magnetic field lines connected to IMF).



2-5: Polar Cap (open)

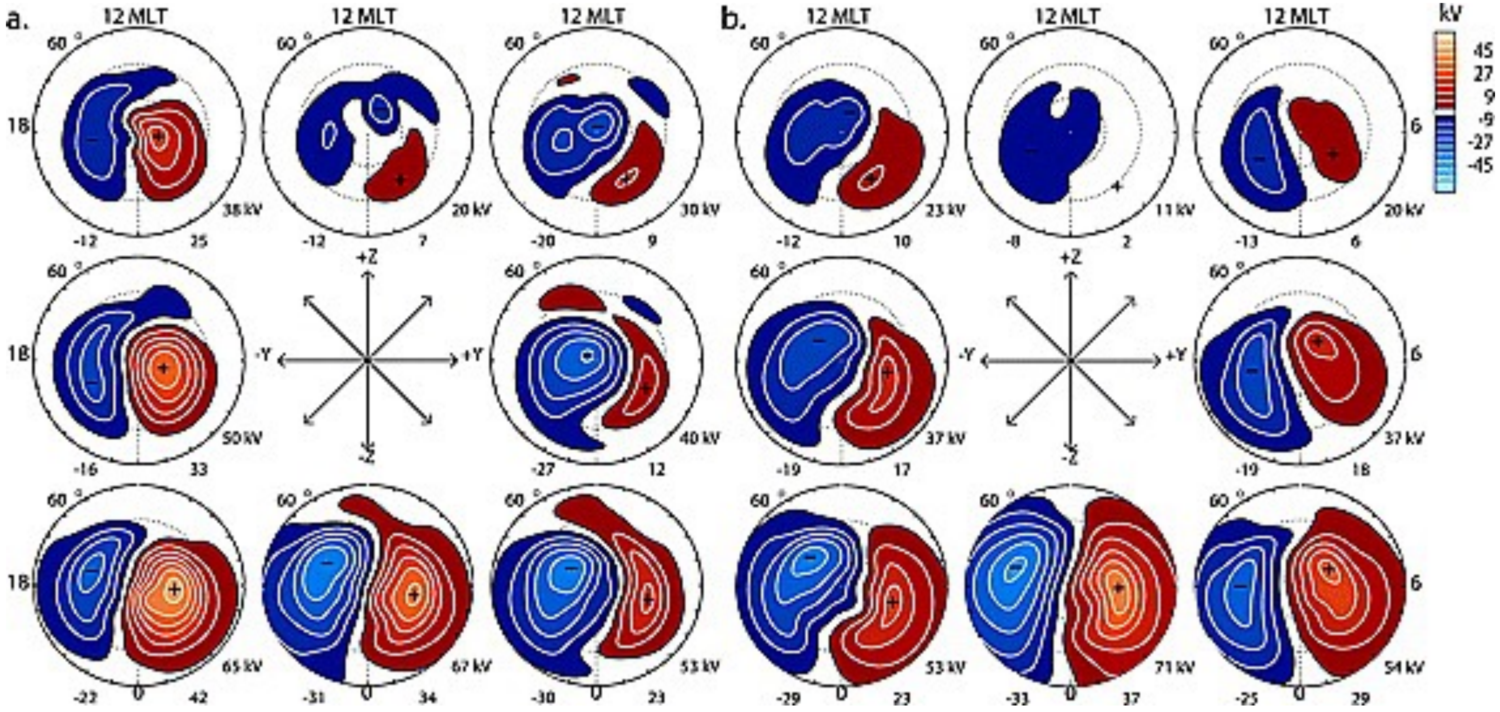
6-9: Magnetotail plasma sheet (closed)

(Not to scale)

See also Figure 9.11 in Kivelson & Russell

Large-Scale Convection: Ionosphere

Contours=Equipotentials=Contours of Motion



Cousins and Shepherd (JGR AGU, 2015)
Data source: SuperDARN HF radar

Convection Pattern
Function of:

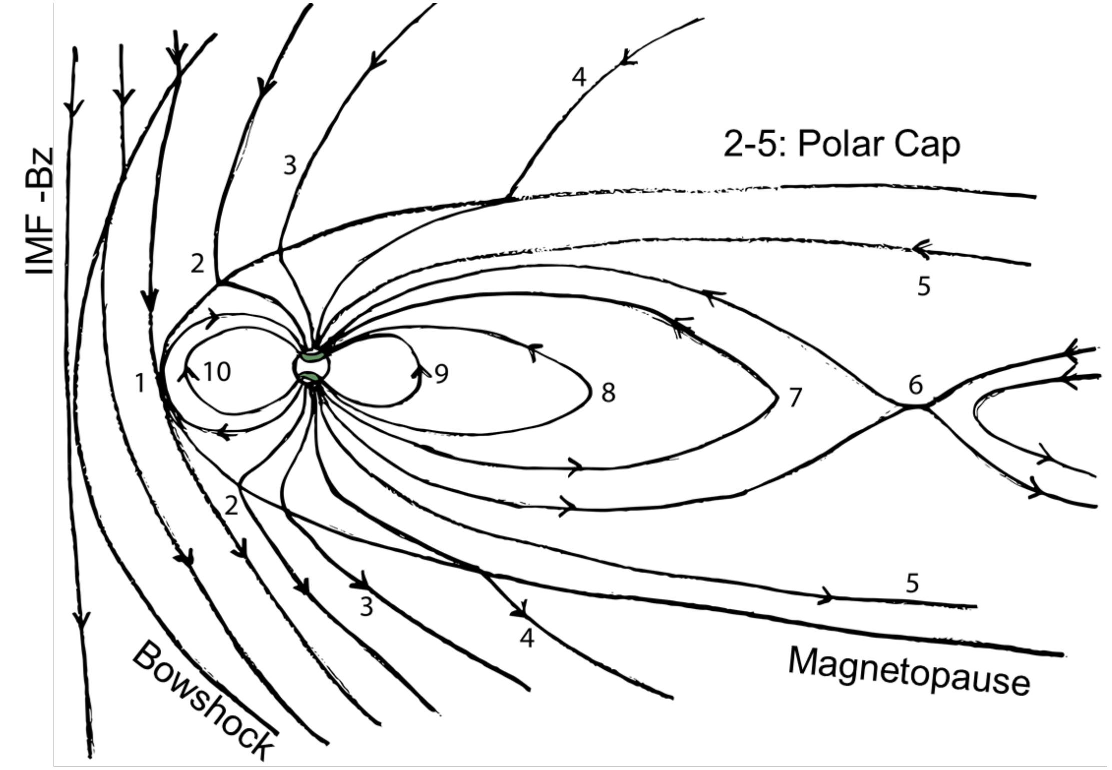
- IMF clock angle
- Hemisphere
- Solar wind velocity
- Most common: 2-cell
- 3-cell and 4-cell possible for +Bz

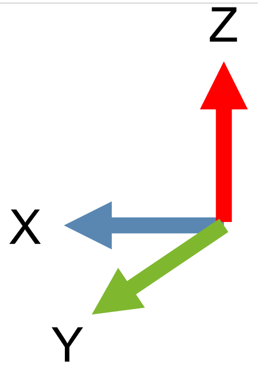
Read More!

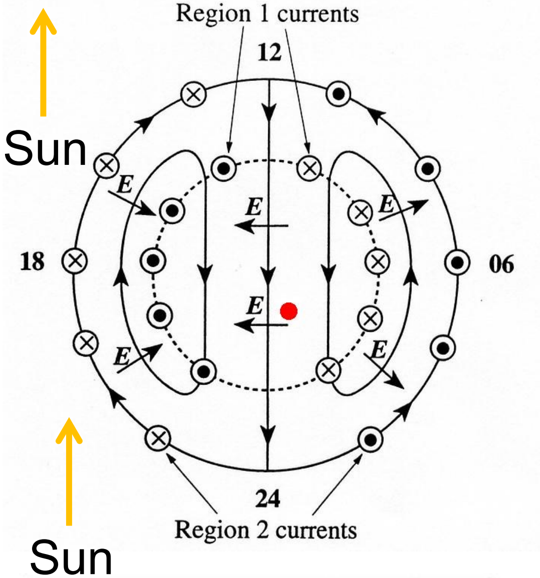
- Heppner (1977) using Ogo 6
- Heppner and Maynard (1987) using Dynamics Explorer 2
- Rich and Hairston (1994) using DMSP
- Weimer (1995) using DE satellite data
- Ruohoniemi and Greenwald (1996) using ground-based HF radar data
- Papitashvili and Rich (2002) using DMSP
- Ruohoniemi and Greenwald (2005) and Cousins and Shepherd (2015) using SuperDARN HF radar data
- Haaland et al. (2007) using Cluster data

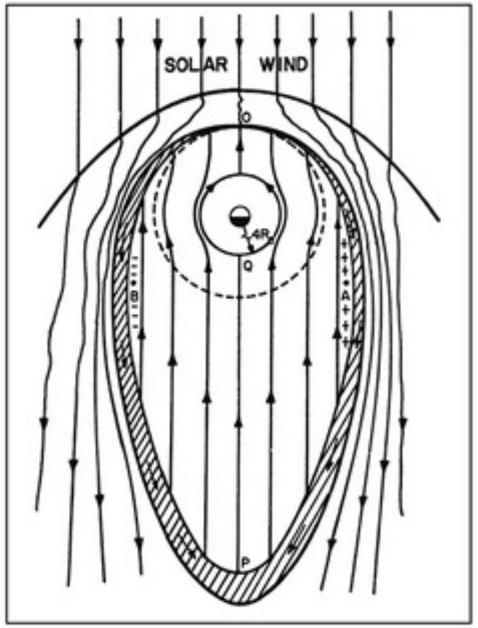
Large-Scale Convection: Magnetosphere-Ionosphere Coupling

Cowley (AGU Geophys. Monogr. Series, 2000)









$$E = \frac{\phi}{D}$$

$$\boldsymbol{E} = -\boldsymbol{V}\times\boldsymbol{B}$$

$$E_{PC} = \frac{\phi}{2R_{PC}}$$

$R_{PC} \approx 0.2\,R_E$

$V_{PC} \approx 330\,m/s$

$B_{PC} \approx 62000nT$

$\phi \approx 52$ kV

Note: 330 m/s is high, background flow closer to ~100 m/s

$$E_{Tail} = ?$$

$$E_{Tail} = \frac{\phi}{2R_{tail}}$$

$R_{tail} \approx 20\,R_E$

$$E_{Tail} \approx 0.2\,mV/m$$

9.4.2 in Kivelson & Russell

Convection: Particle Motion

Particle trajectories, magnetic and electric fields

<u>Guiding-center motion</u>

$$\mathbf{V}_{\mathrm{GC}} = \frac{\mathbf{dr}}{\mathbf{dt}} = \mathbf{V}_{\mathbf{E}\times\mathbf{B}} + \mathbf{V}_{\nabla\mathbf{B}}$$

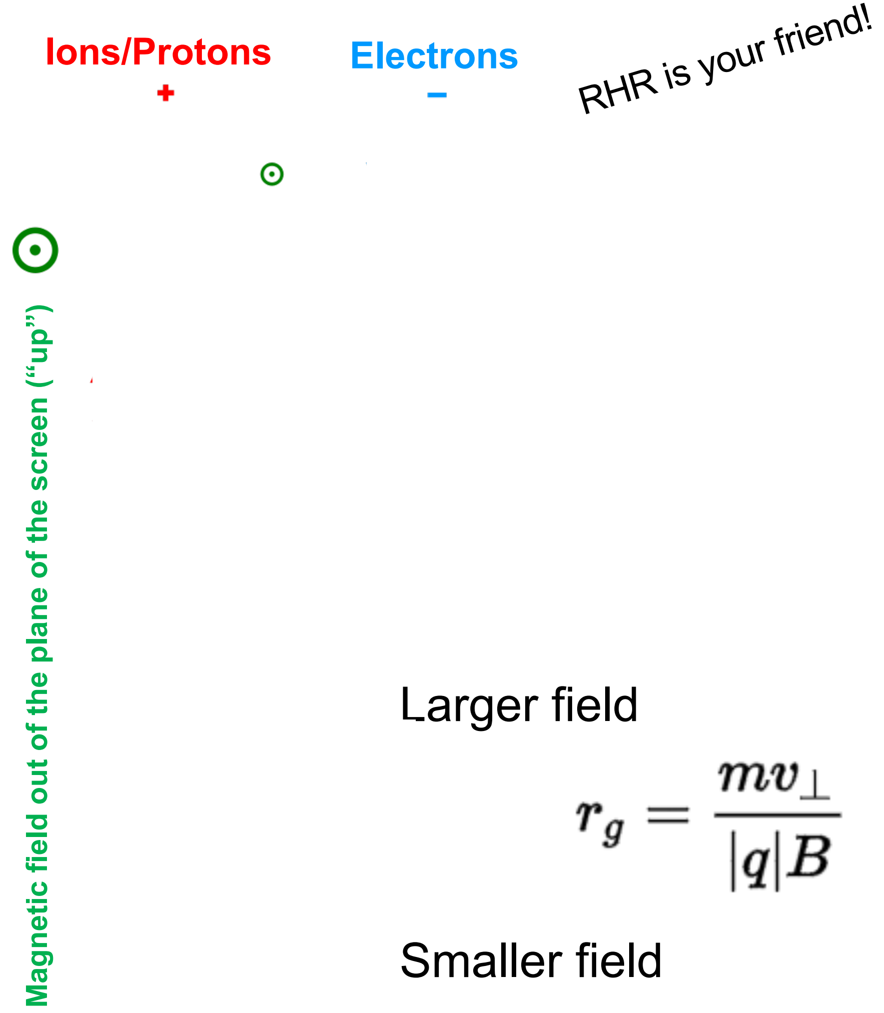
--*Northrop*, 1963
(90° pitch angles)

$$\vec{v}_f = \frac{1}{q}\frac{\vec{F}\times\vec{B}}{B^2}$$

$$\vec{v}_E = \frac{\vec{E}\times\vec{B}}{B^2}$$

$$\mathbf{V}_{drift} = \boxed{\frac{mv_\perp^2}{2qB}}\frac{(-\nabla_\perp B)\times\mathbf{B}}{B^2}$$

More energetic particles drift faster



$$r_g = \frac{mv_\perp}{|q|B}$$

Image based on figure from: https://en.wikipedia.org/wiki/Guiding_center

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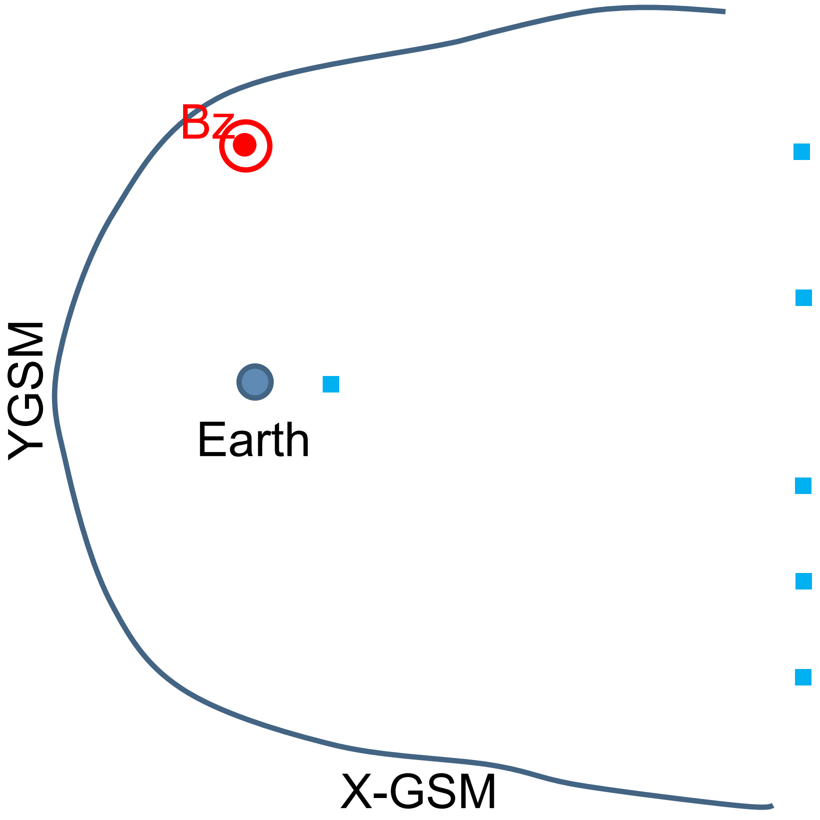

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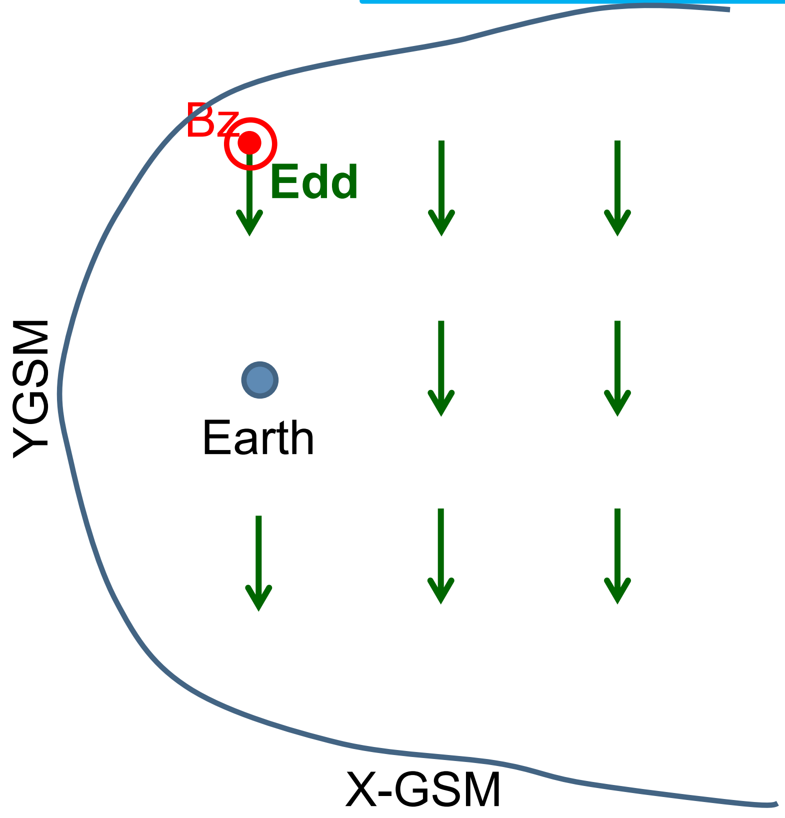
--*Northrop*, 1963

Edd: *dawn-dusk electric field*

← Sun

$$\vec{v}_E = \frac{\vec{E} \times \vec{B}}{B^2}$$

Draw convection E vectors



Convection: Particle Motion

Particle trajectories, the magnetic and electric fields

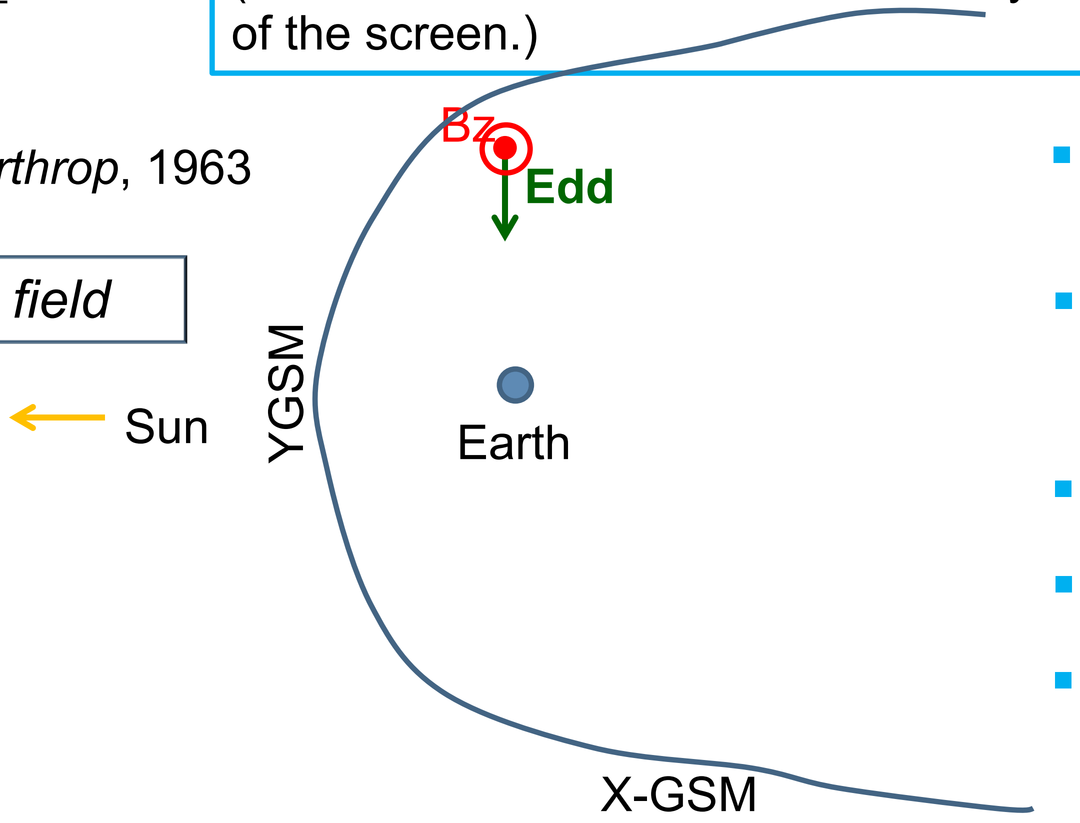
Launch particles from several points downtail. How will they move? (Electrons? Ions?)
(For now, assume B is constant everywhere pointing out of the screen.)

Guiding-center motion

$$\mathbf{V}_{\mathrm{GC}} = \frac{\mathbf{dr}}{\mathbf{dt}} = \mathbf{V}_{\mathbf{E}\times\mathbf{B}}$$

--*Northrop*, 1963

Edd: *dawn-dusk electric field*



$$\vec{v}_E = \frac{\vec{E} \times \vec{B}}{B^2}$$

Convection: Particle Motion

Particle trajectories, the magnetic and electric fields

Contours of motion can be drawn from these relationships.

Guiding-center motion

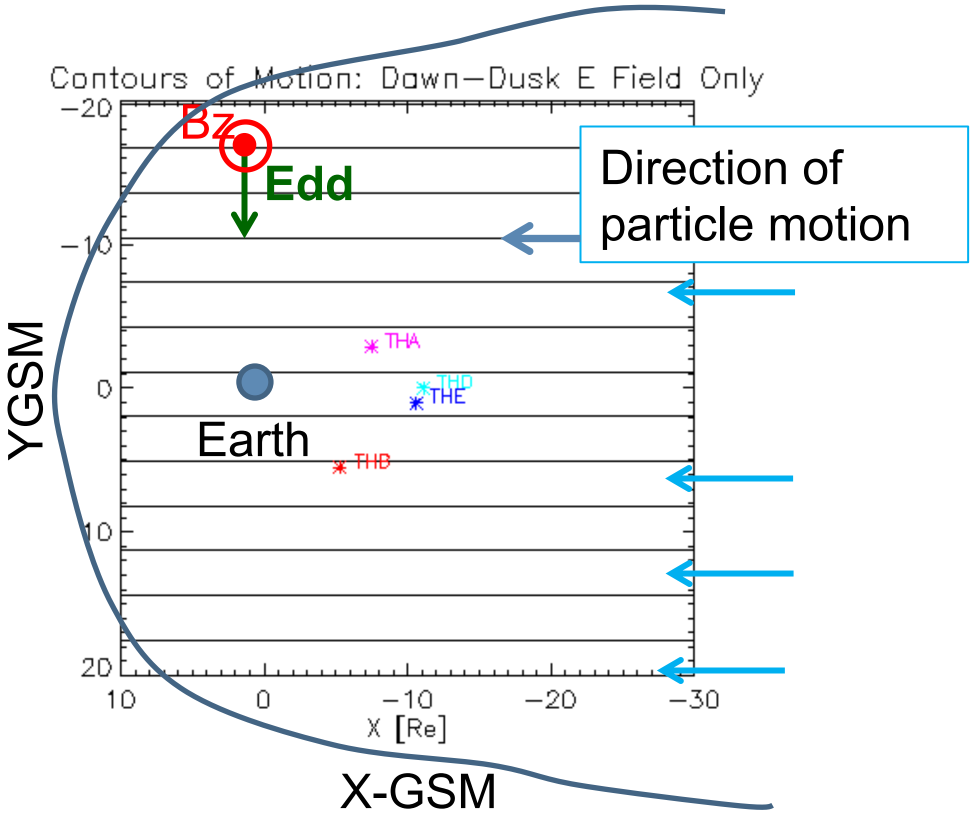
$$\mathbf{V}_{\mathrm{GC}} = \frac{\mathbf{dr}}{\mathbf{dt}} = \mathbf{V}_{\mathbf{E}\times\mathbf{B}}$$

--*Northrop*, 1963

Edd: *dawn-dusk electric field aka "convection Efield"*

$$\vec{v}_E = \frac{\vec{E}\times\vec{B}}{B^2}$$

← Sun



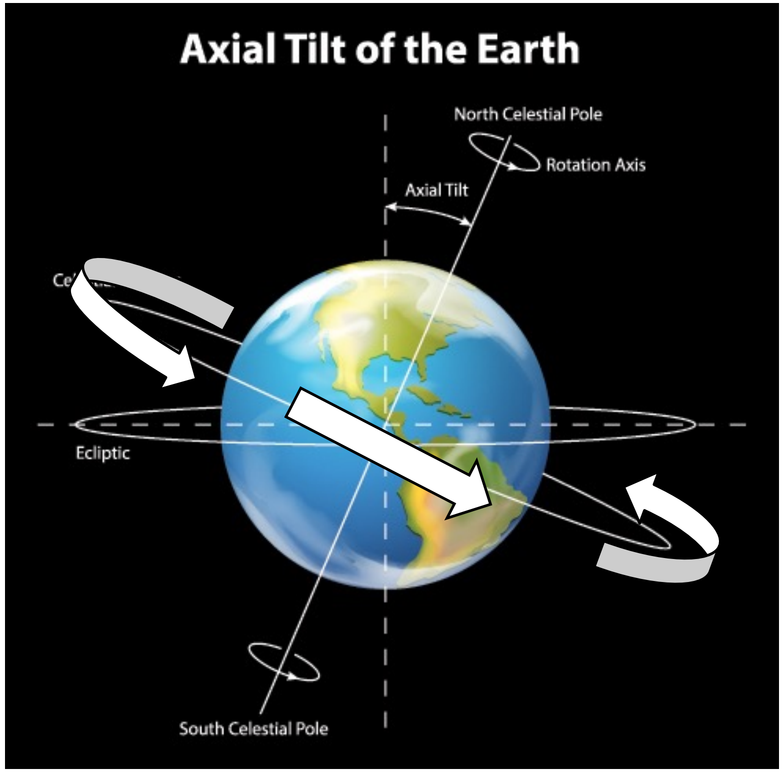
Direction of particle motion

Convection: Particle Motion

Particle trajectories, magnetic and electric fields

What affect might the rotation of the Earth have on our particle trajectories in space?

1. Consider the stationary reference frame outside of the Earth.
2. Consider ionospheric plasma is only partially ionized, and neutral collision frequency is high.



Credit: BlueRingMedia/Shutterstock.com

How do we calculate the co-rotation electric field?

$$\boldsymbol{E} = -\boldsymbol{V}\times\boldsymbol{B}$$

What is **V**?

$$\boldsymbol{V} = \varpi_{E} \times \mathbf{r}$$

$$\boldsymbol{V} = \frac{2\pi}{24\,h}\,\mathrm{r}\,\boldsymbol{e}_{\boldsymbol{\phi}}$$

What does the co-rotation electric field profile look like on the equatorial plane?

See ch. 5 in Baumjohann &Treumann, 10.5.7 in Kivelson & Russell

Convection: Particle Motion

Particle trajectories, magnetic and electric fields

Contours of motion can be drawn from these relationships.

Guiding-center motion

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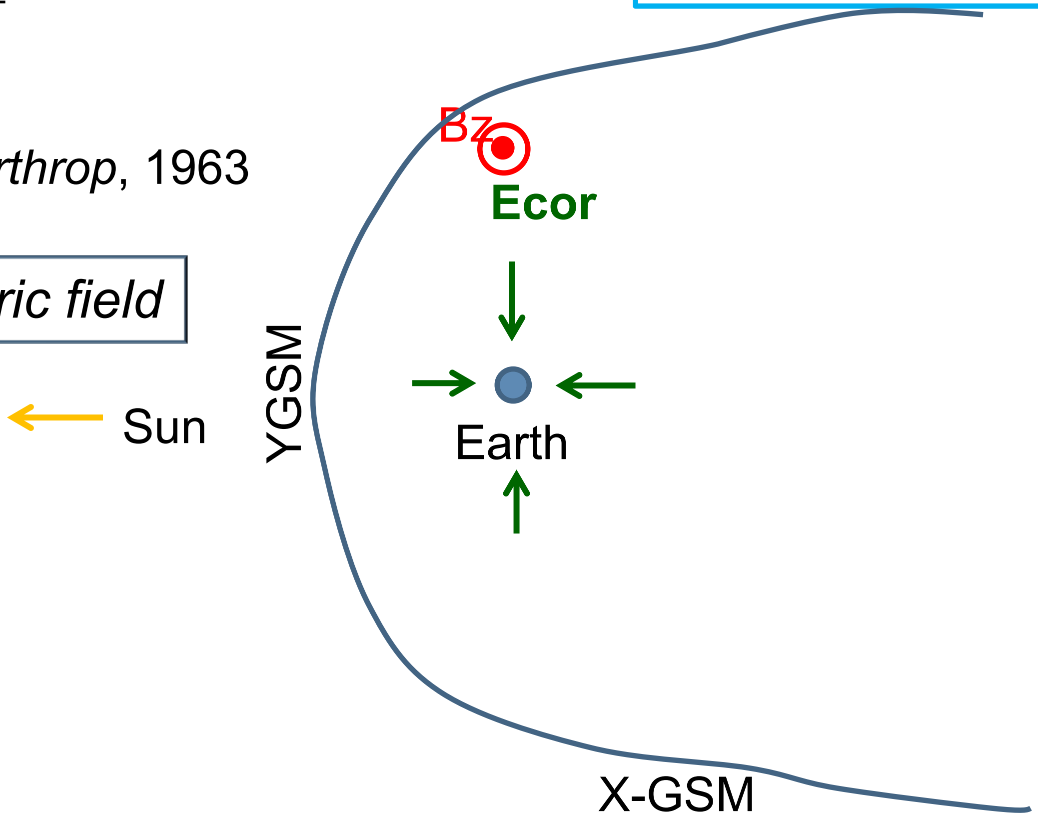
--*Northrop*, 1963

Ecor: *Corotation electric field*

$$\vec{v}_E = \frac{\vec{E}\times\vec{B}}{B^2}$$

Draw corotation E vectors



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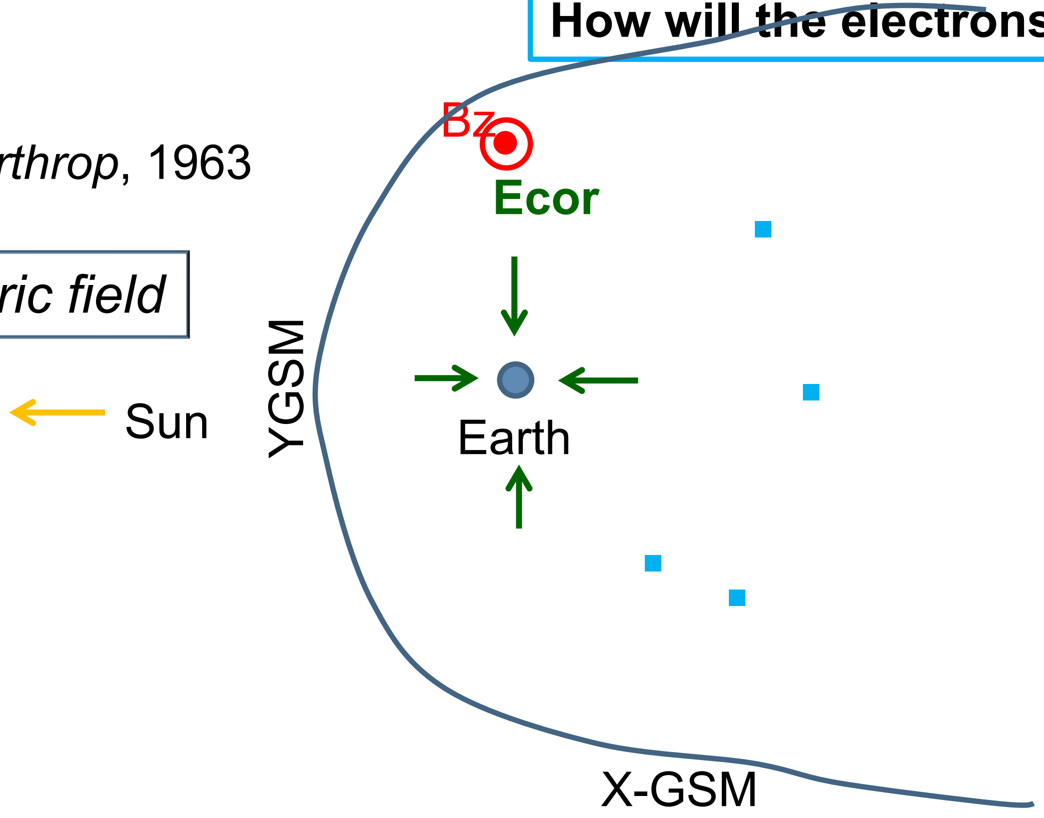
Ecor: *Corotation electric field*

$$\vec{v}_E = \frac{\vec{E}\times\vec{B}}{B^2}$$

**Launch particles from several points.
How will the electrons move? The ions?**



Convection: Particle Motion

Particle trajectories, magnetic and electric fields

Contours of motion can be drawn from these relationships.

Guiding-center motion

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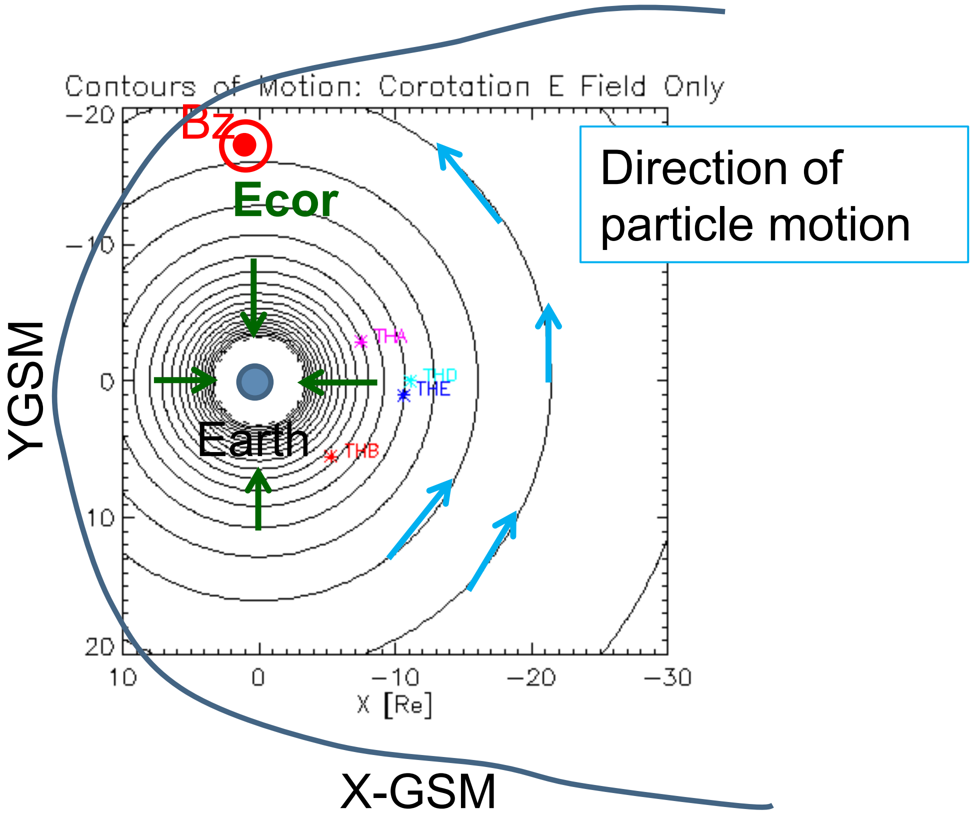
--*Northrop*, 1963

Ecor: *Corotation electric field*

← Sun

$$\vec{v}_E = \frac{\vec{E}\times\vec{B}}{B^2}$$



Convection: Particle Motion

Particle trajectories, magnetic and electric fields

What direction is the gradient?
(Now treat B as a dipole at Earth.)

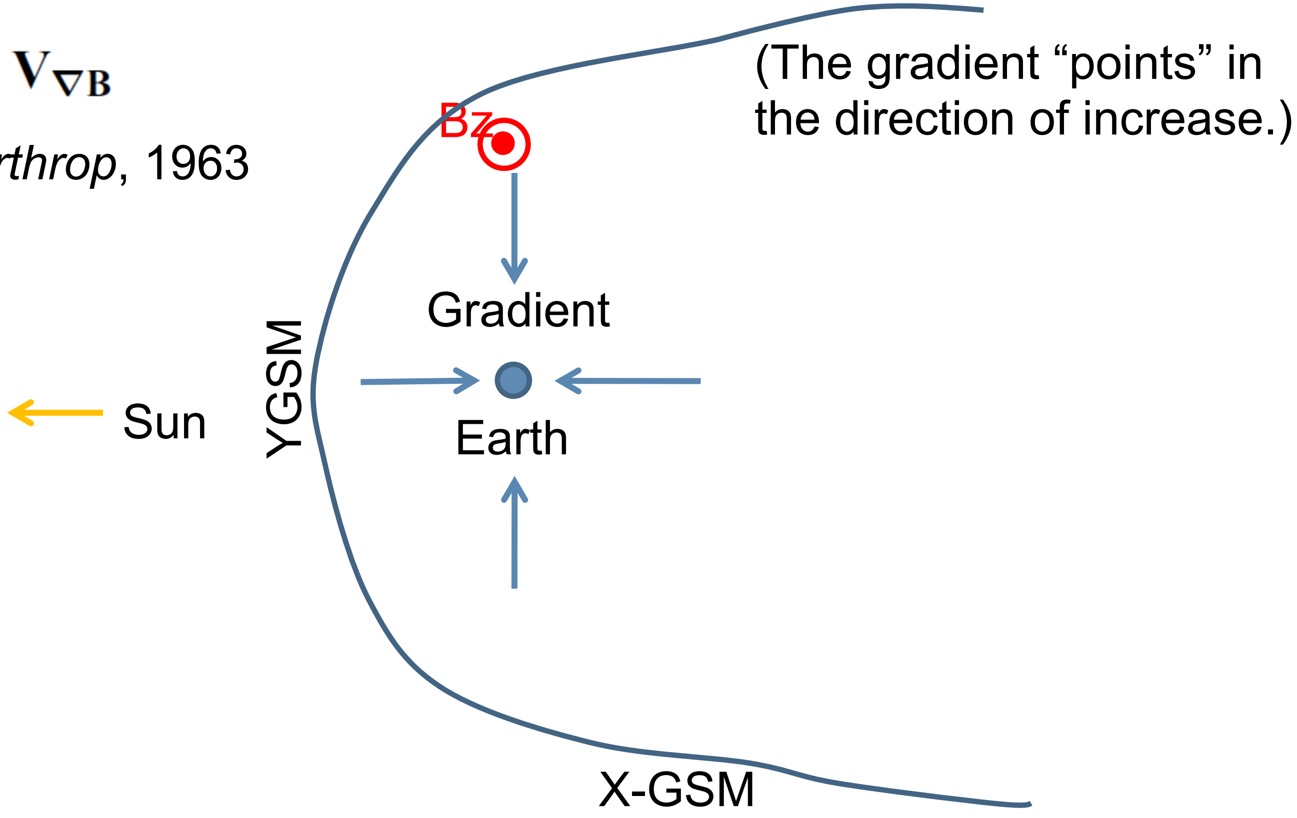
Guiding-center motion

$$\mathbf{V}_{\mathrm{GC}} = \frac{\mathbf{dr}}{\mathbf{dt}} = \mathbf{V}_{\mathbf{E}\times\mathbf{B}} + \mathbf{V}_{\nabla\mathbf{B}}$$

--*Northrop*, 1963

Grad-B Drift

$$\mathbf{v}_{drift} = \frac{mv_\perp^2}{2qB}\frac{(-\nabla_\perp B)\times\mathbf{B}}{B^2}$$



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Convection: Particle Motion

Particle trajectories, magnetic and electric fields

<u>Guiding-center motion</u>

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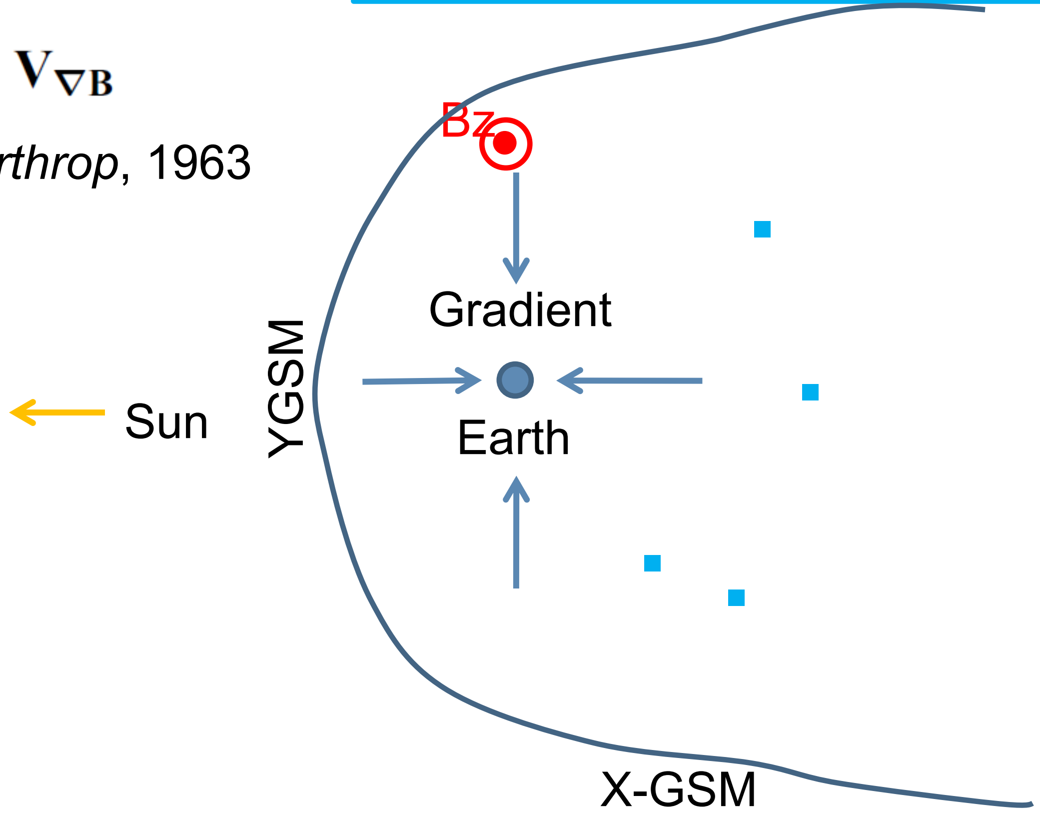
Grad-B Drift

$$\mathbf{v}_{drift} = \frac{mv_\perp^2}{2qB}\frac{(-\nabla_\perp B)\times\mathbf{B}}{B^2}$$

Launch energetic particles.
How will electrons move?



Convection: Particle Motion

Particle trajectories, magnetic and electric fields

Contours of motion can be drawn from these relationships.

Guiding-center motion

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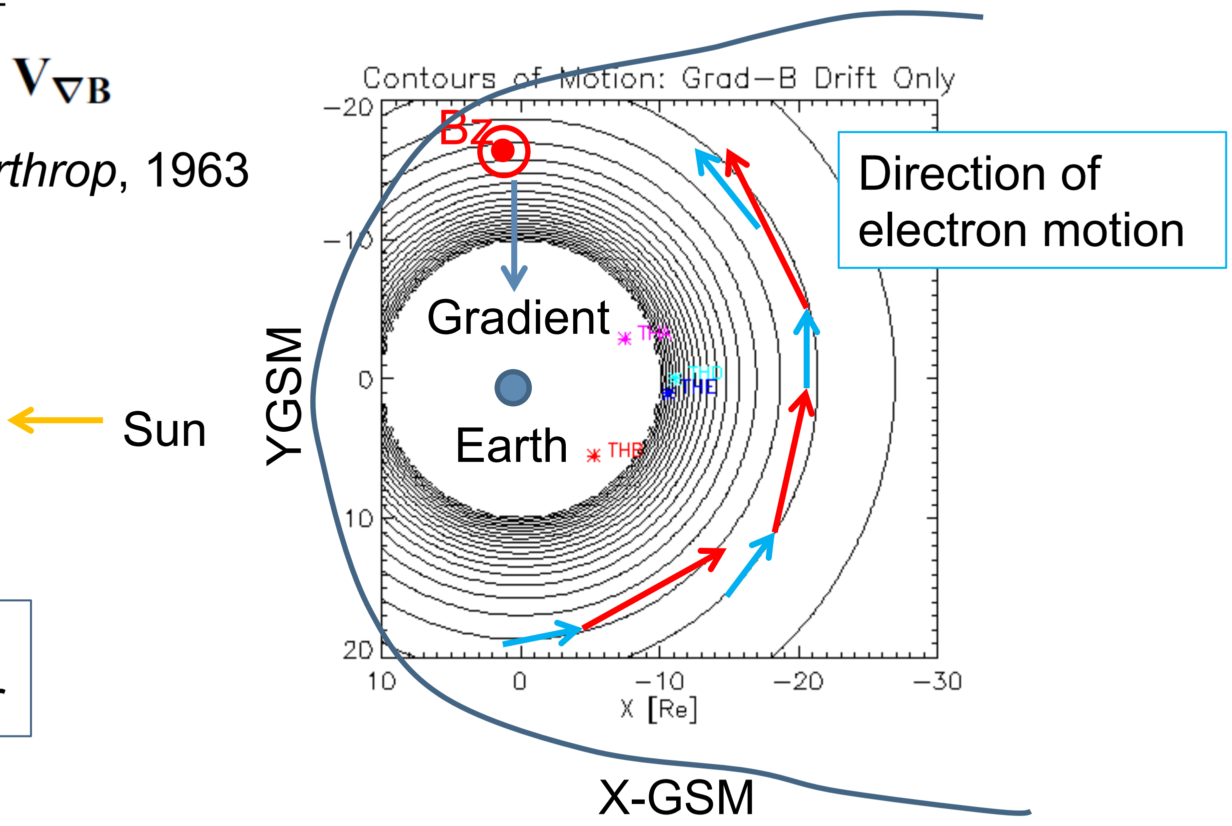
--*Northrop*, 1963

Grad-B Drift

$$\mathbf{v}_{drift} = \frac{mv_{\perp}^{2}}{2qB}\frac{(-\nabla_{\perp}B)\times\mathbf{B}}{B^{2}}$$

Energy-dependent:
Energetic e- drift faster



Direction of electron motion

Convection: Particle Motion

Particle trajectories, magnetic and electric fields

Guiding-center motion

$$\mathbf{V}_{\mathrm{GC}} = \frac{\mathbf{dr}}{\mathbf{dt}} = \mathbf{V}_{\mathbf{E}\times\mathbf{B}} + \mathbf{V}_{\nabla\mathbf{B}}$$

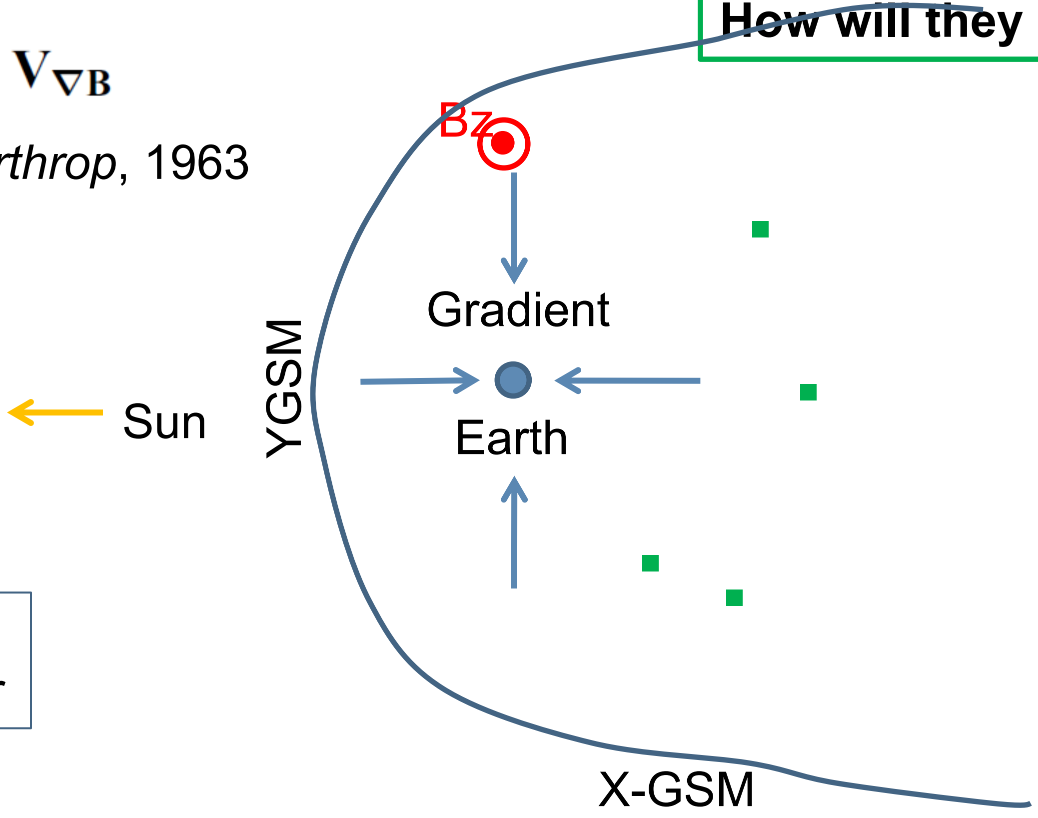
--*Northrop*, 1963

Grad-B Drift

$$\mathbf{v}_{drift} = \frac{mv_{\perp}^{2}}{2qB}\frac{(-\nabla_{\perp}B)\times\mathbf{B}}{B^{2}}$$

Energy-dependent:
Energetic e- drift faster

Launch energetic i+.
How will they move?



Convection: Particle Motion

Particle trajectories, magnetic and electric fields

Contours of motion can be drawn from these relationships.

Guiding-center motion

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--*Northrop*, 1963

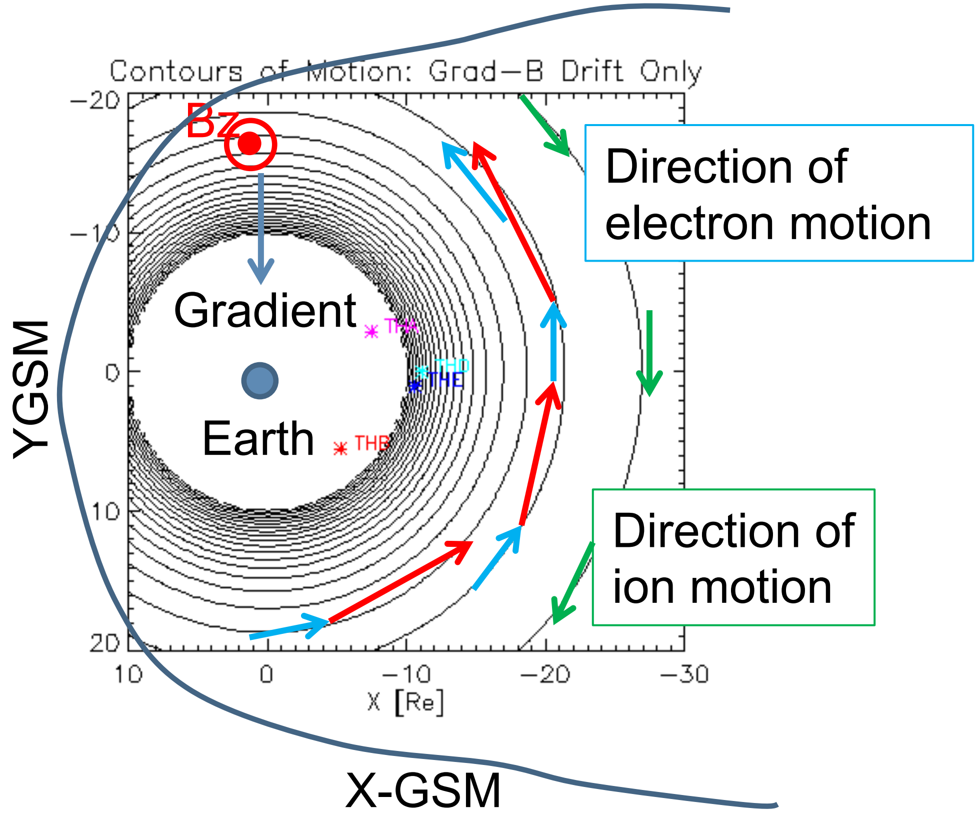
Grad-B Drift

$$\mathbf{v}_{drift} = \frac{mv_\perp^2}{2qB}\frac{(-\nabla_\perp B)\times\mathbf{B}}{B^2}$$

← Sun

Energy-dependent:
Energetic e- drift faster

Energetic i+ drift
opposite direction
(clockwise)



Direction of electron motion

Direction of ion motion

Convection: Particle Motion

Particle trajectories, magnetic and electric fields

Guiding-center motion

$$V_{GC} = \frac{dr}{dt} = V_{E\times B} + V_{\nabla B} \boxed{+ V_R}$$

--*Northrop*, 1963

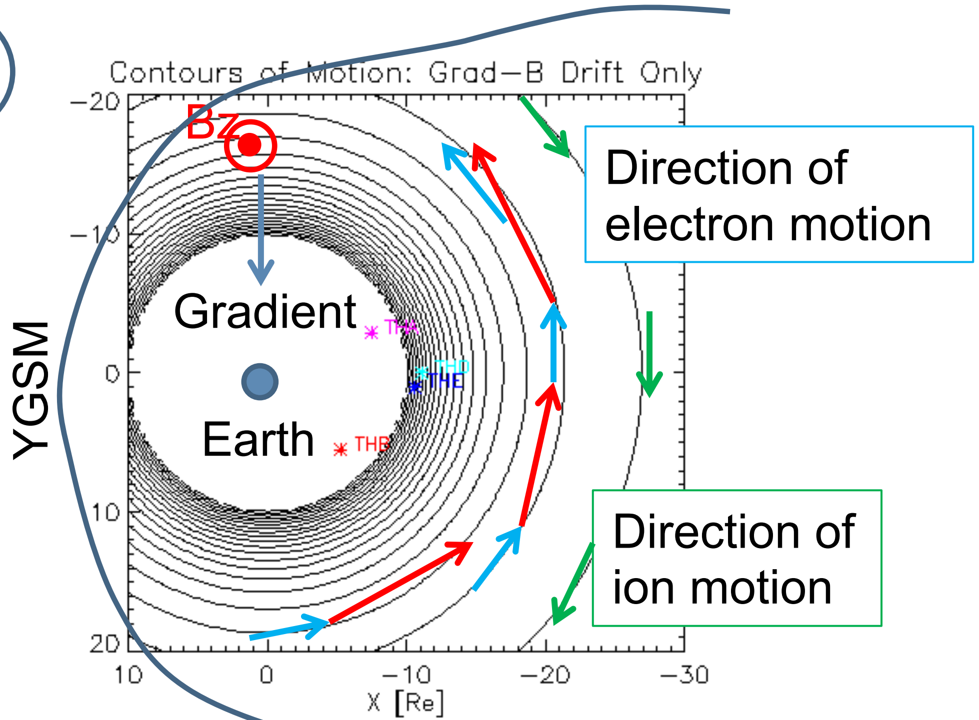
Curvature Drift

$$\vec{v}_R = \frac{2K_\parallel}{qB}\frac{\vec{R}_c\times\vec{B}}{R_c^2 B}$$

← Sun

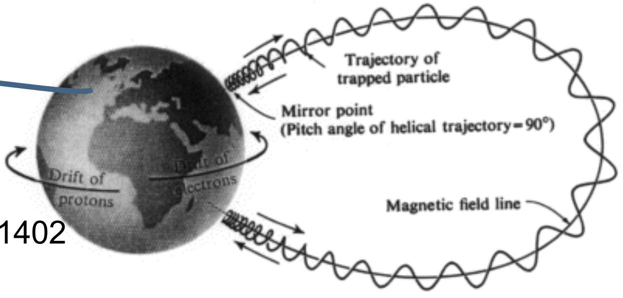
Energy-dependent:
Energetic e- drift faster

Energetic i+ drift
opposite direction
(clockwise)



Direction of
electron motion

Direction of
ion motion



R.E. Mars, 2002, Lawrence Livermore
National Laboratory Report UCRL-ID-151402

Convection: Particle Motion

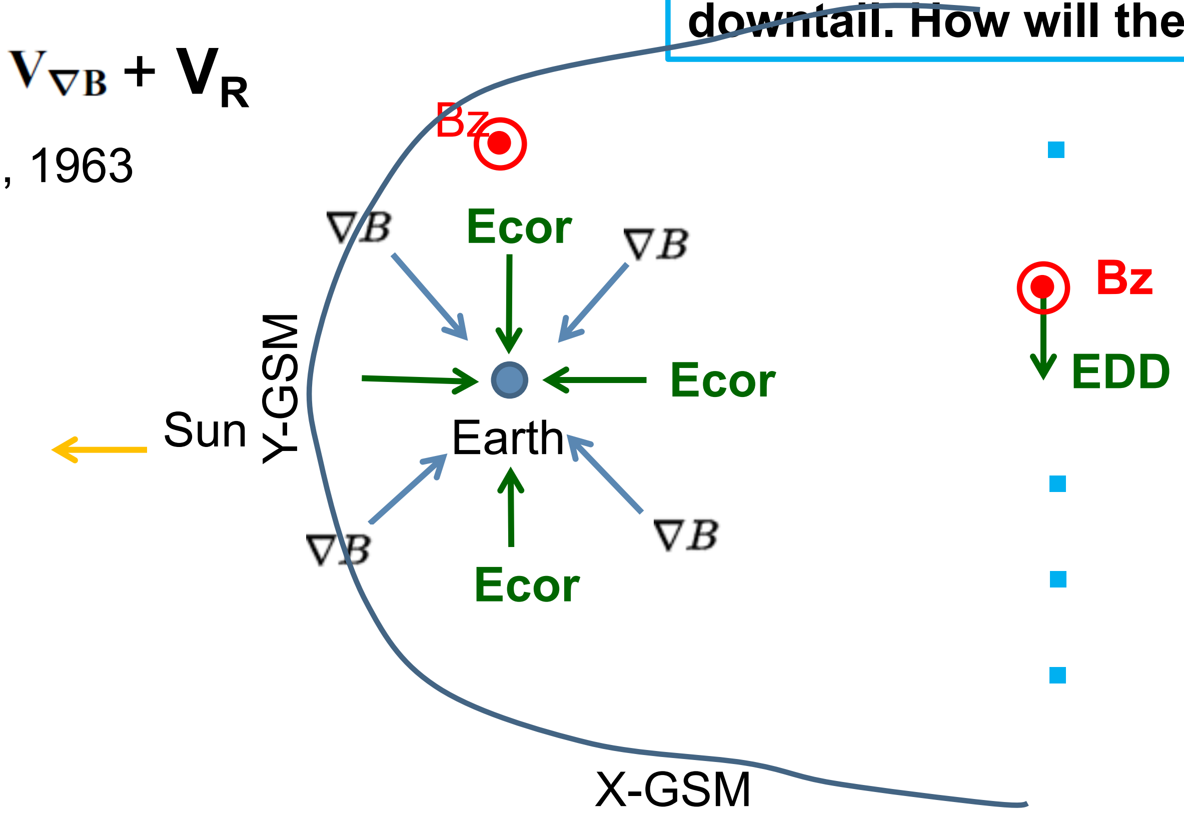
Particle trajectories, magnetic and electric fields

Guiding-center motion

$$\mathbf{V}_{\mathrm{GC}} = \frac{\mathbf{dr}}{\mathbf{dt}} = \mathbf{V}_{\mathbf{E}\times\mathbf{B}} + \mathbf{V}_{\nabla\mathbf{B}} + \mathbf{V}_{\mathbf{R}}$$

--*Northrop*, 1963

Launch e- from several points downtail. How will the e- move?



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Convection: Particle Motion

Particle trajectories, magnetic and electric fields

Contours of motion can be drawn from these relationships.

Guiding-center motion

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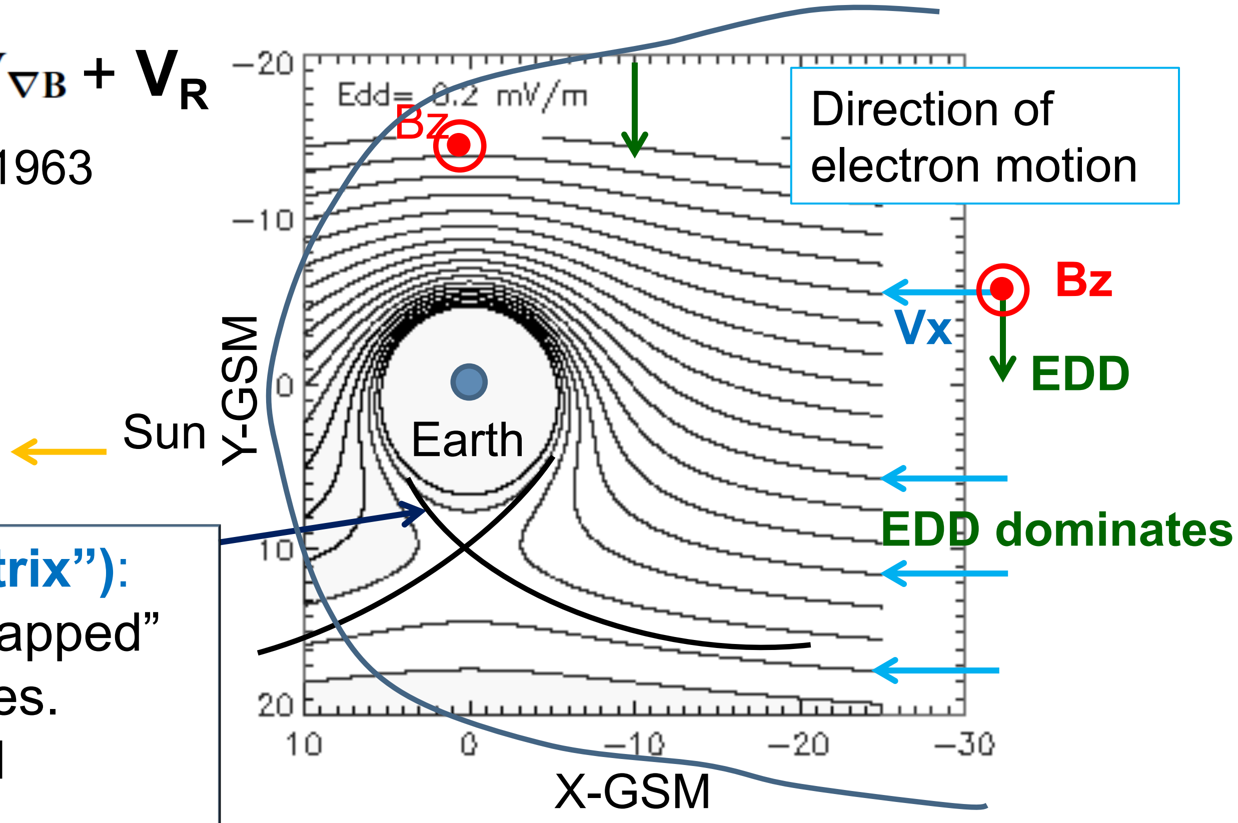
--*Northrop*, 1963



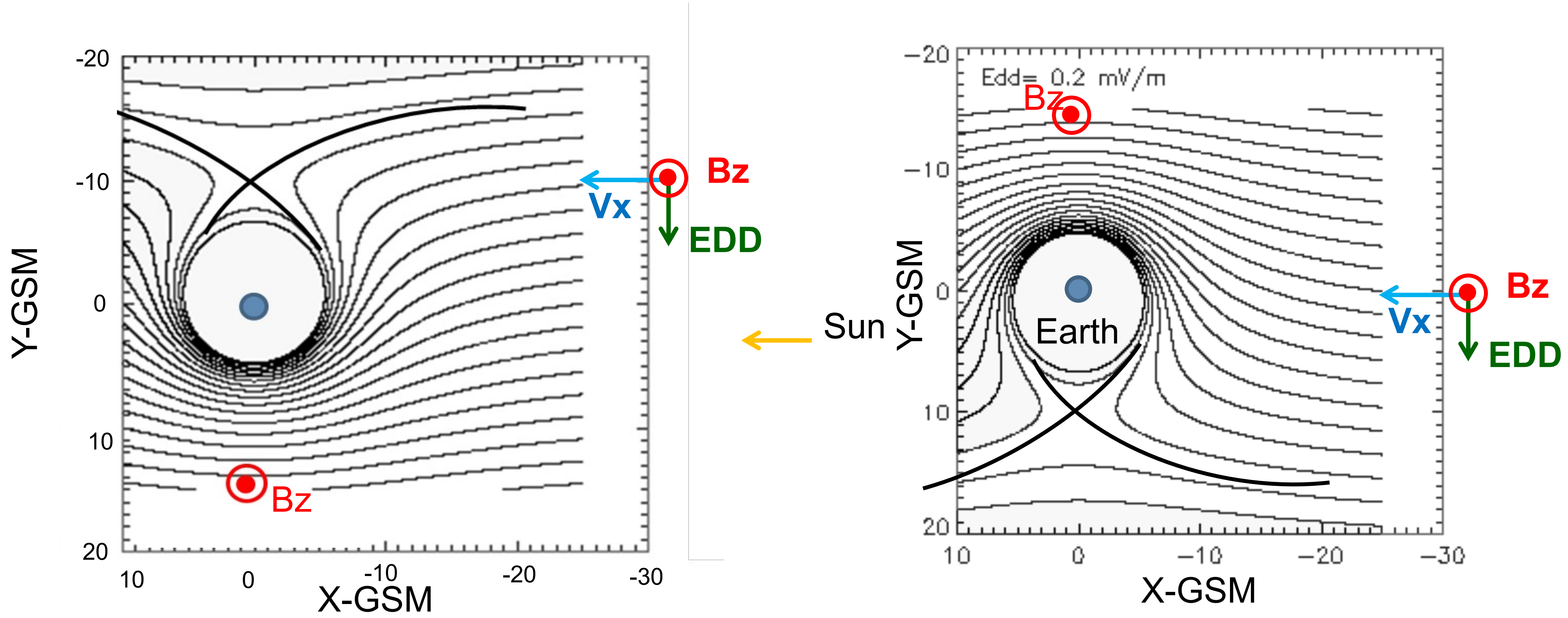
Alfvén Layer ("separatrix"): Boundary separating "trapped" and "convecting" particles. Inside, Ecor, gradB, and curvature drift dominate.

Convection: Particle Motion

Particle trajectories, magnetic and electric fields

$$\mathbf{v}_{drift} = \frac{mv_\perp^2}{2qB}\frac{(-\nabla_\perp B)\times\mathbf{B}}{B^2}$$

Which plot below represents ion motion and the ion Alfven Layer?



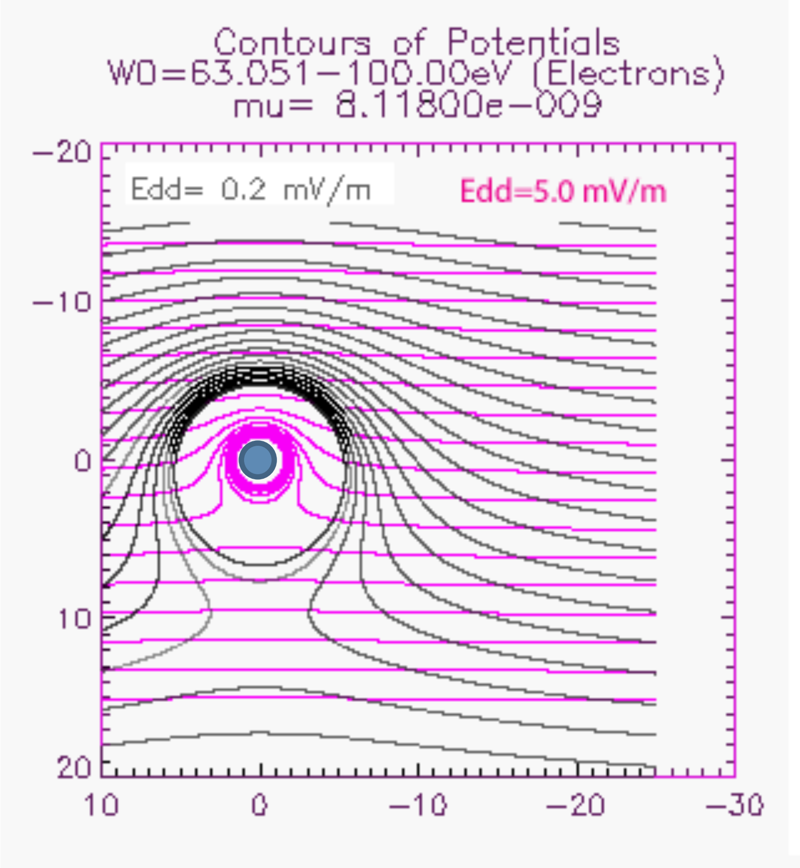
Both are correct! The left is for energetic (hot) ions, while the right is for cold ions that do not gradB drift.

But how do particles get past the Alfvén Layer?

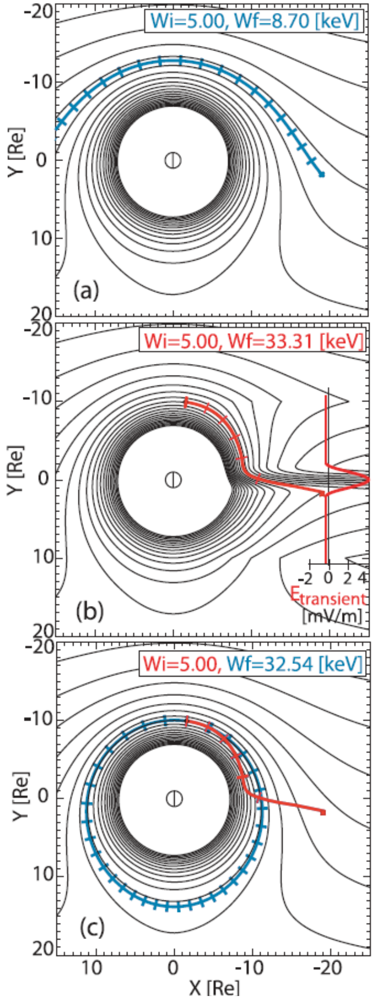
In terms of convection

Enhanced solar wind driving: Faster **V**→Larger **Edd**
Walker and Kivelson, 1975

← Sun



Localized fast flows:
Faster **V**→Larger **E**
Gabrielse et al. (AGU JGR, 2012)

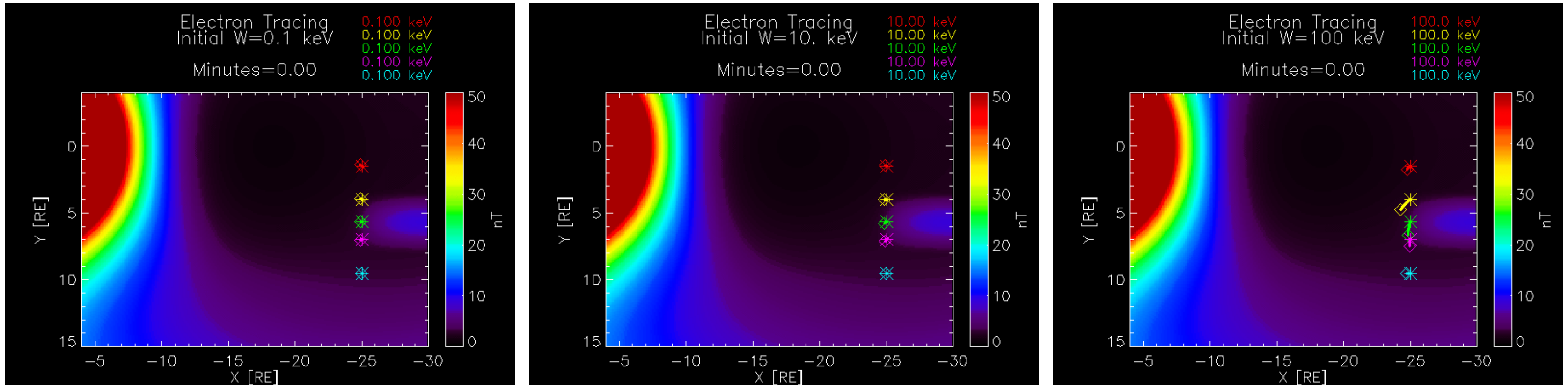


e.g., Pfitzer and Winckler, 1969; Baker et al., 1979; Moore et al., 1981; Aggson et al., 1983; Mauk and Meng, 1987; Reeves et al., 1990

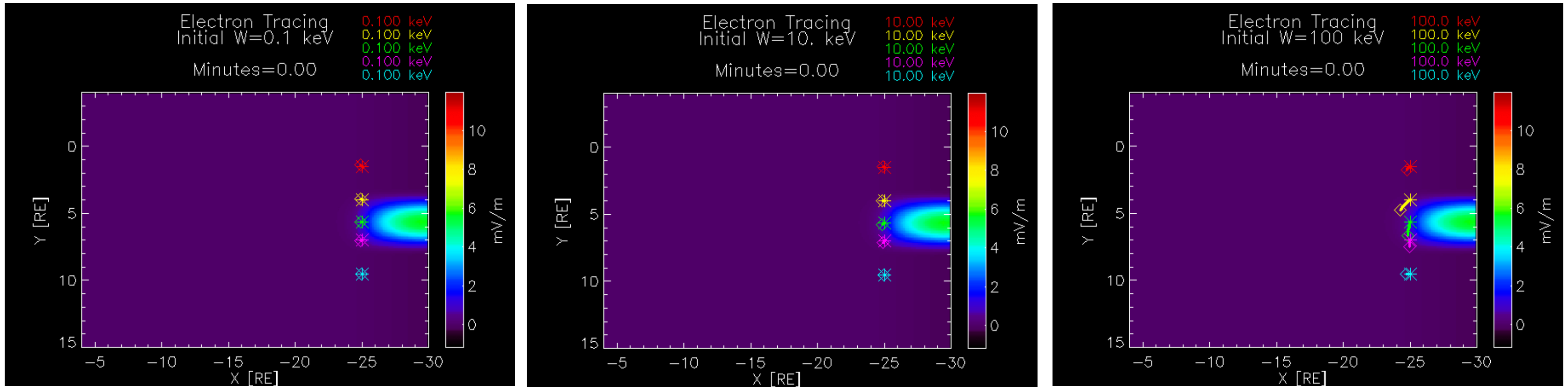
Particle Transport

Initial energy: 0.1 keV 10 keV 100 keV

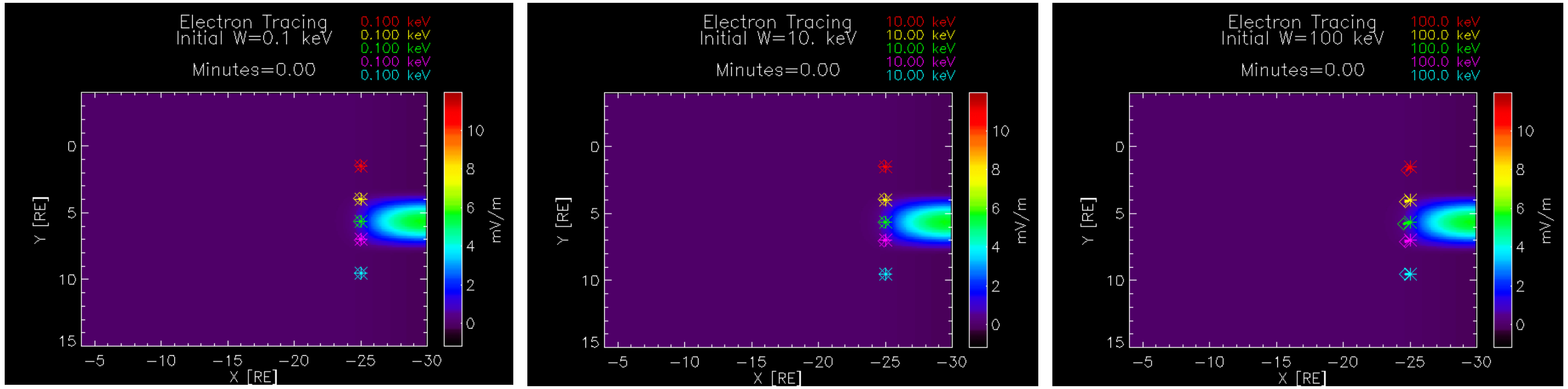
E and B
B plotted



E and B
E plotted



No B
E plotted



Gabrielse et al.
(AGU JGR, 2016)

Convection on Mesoscales

Embedded in the background convection pattern, mesoscale phenomena carry the bulk of the load.

Angelopoulos et al. (1992; 1994) showed bursty bulk flows transport >60% of the magnetic flux earthward on the nightside: >10 min of enhanced flow with bursts exceeding 400 m/s.

And mesoscale phenomena start at the beginning with dayside reconnection…

Mesoscale in the magnetosphere: ~1 to several R_E wide
Mesoscale in the ionosphere: ~50-500 km wide

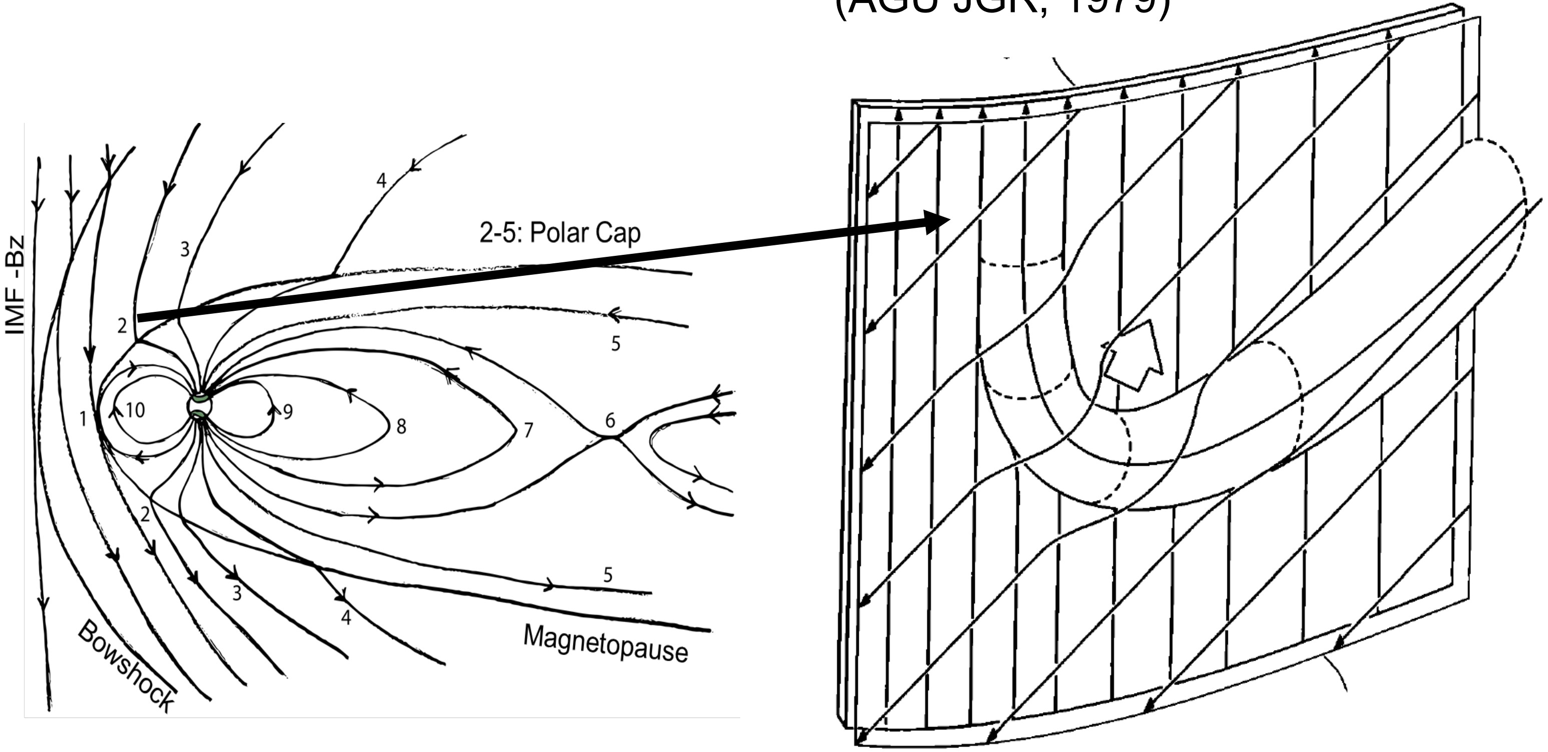
Stop for Break

Convection on Mesoscales

The point of the 2nd half of lecture is to introduce you to different datasets by showing you actual data relevant to mesoscale convection in the magnetosphere.

Convection on Mesoscales

Dayside and Polar Cap Convection



Modified from Russell and Elphic (AGU JGR, 1979)

SuperDARN observed fast flow

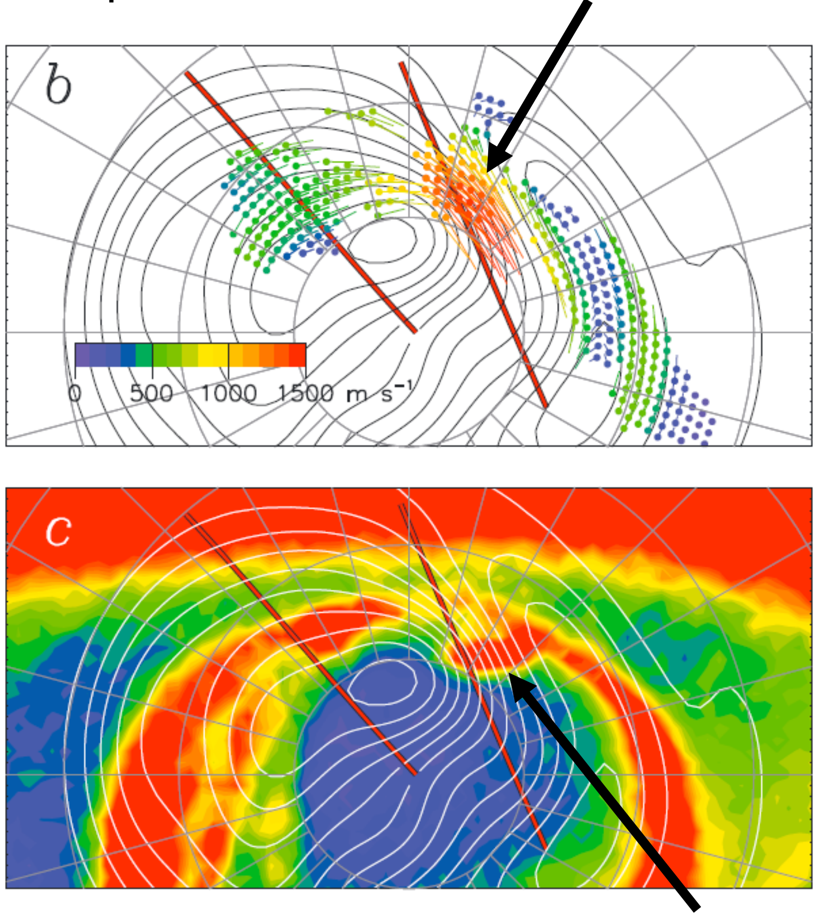


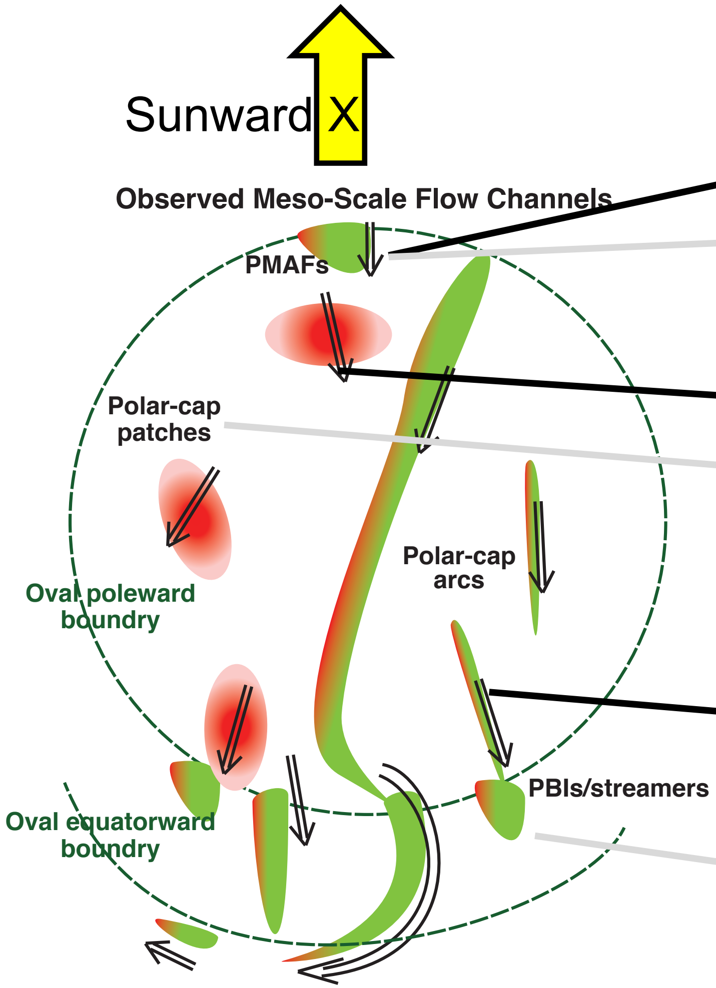
IMAGE FUV/WIC observed PMAF

Milan et al. (AGU JGR, 2016)

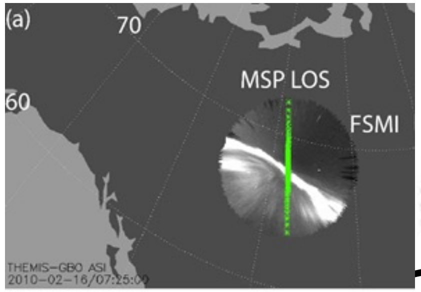
Transient magnetopause reconnection has been related to polar cap patches (Carson et al., 2006; Lockwood and Carlson, 1992), which have been related to poleward moving auroral forms (PMAFs) (Wang et al., 2016)

Convection on Mesoscales

Dayside and Polar Cap Convection



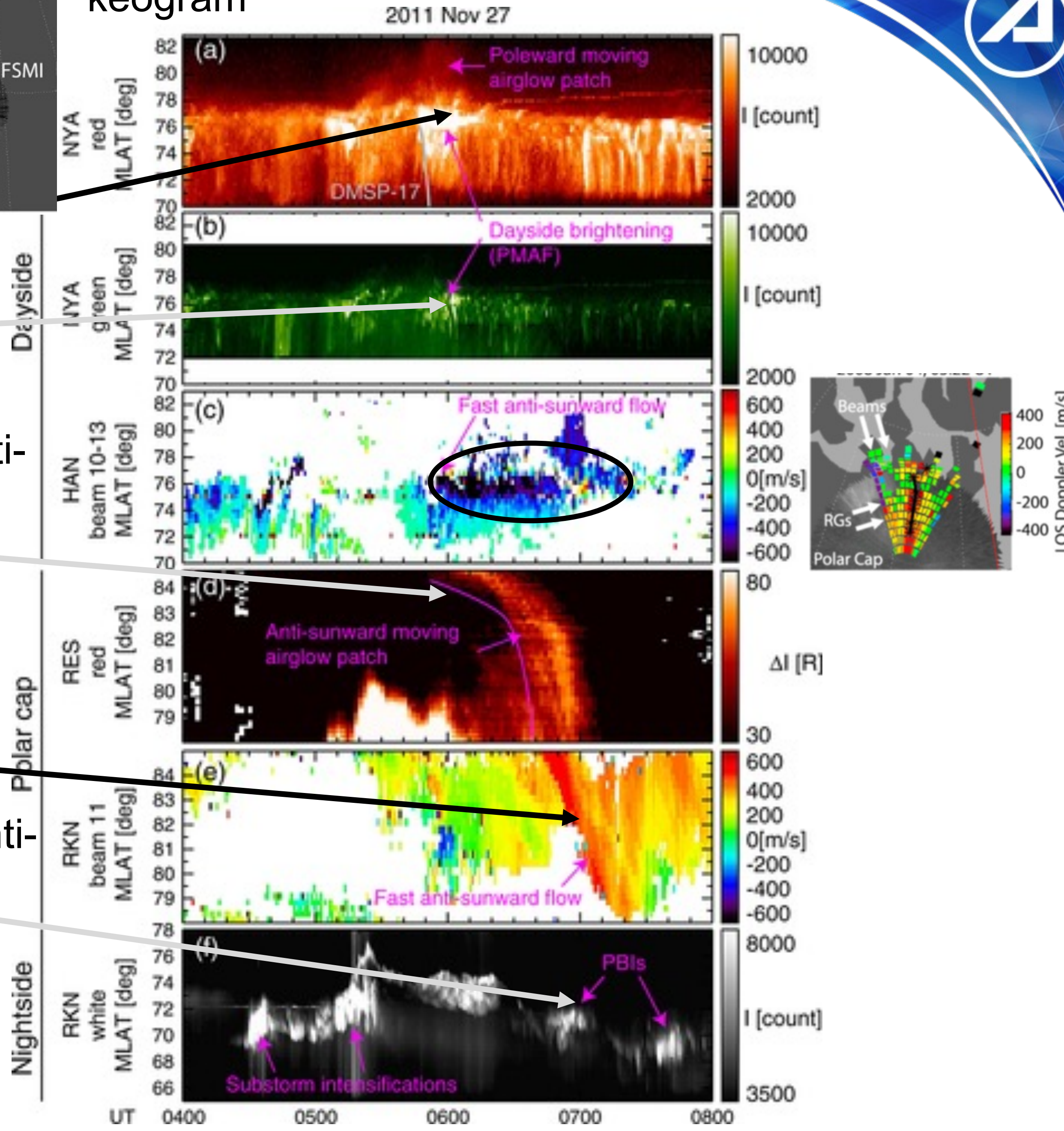
Modified from Lyons al. [AGU JGR 2016]



SuperDARN anti-sunward flow

SuperDARN anti-sunward flow

"keogram"



Nishimura et al. (AGU JGR, 2014)

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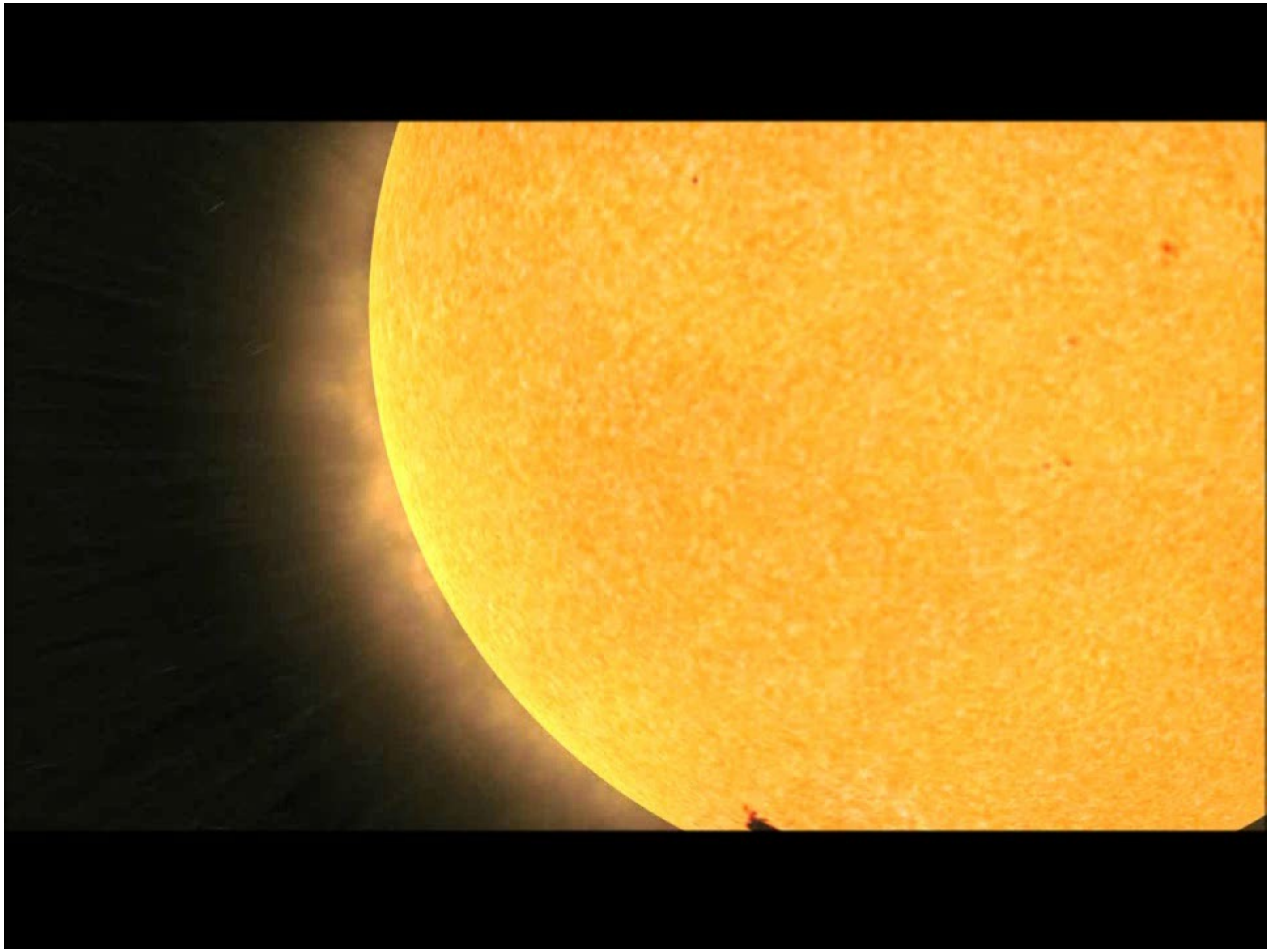
Convection: Substorm

Not steady-state: Increase load in the tail requires explosive response

1. **Growth Phase**
    a) Tail stretching
    b) Flux loading in the tail
    c) Solar wind energy stored in magnetosphere
2. **Expansion Phase**
    a) Initiated by substorm onset
    b) Release of "pent-up energy"
3. **Recovery Phase**
    a. Magnetosphere returns to "ground state"

Duration: ~90 minutes



Video: Courtesy of NASA

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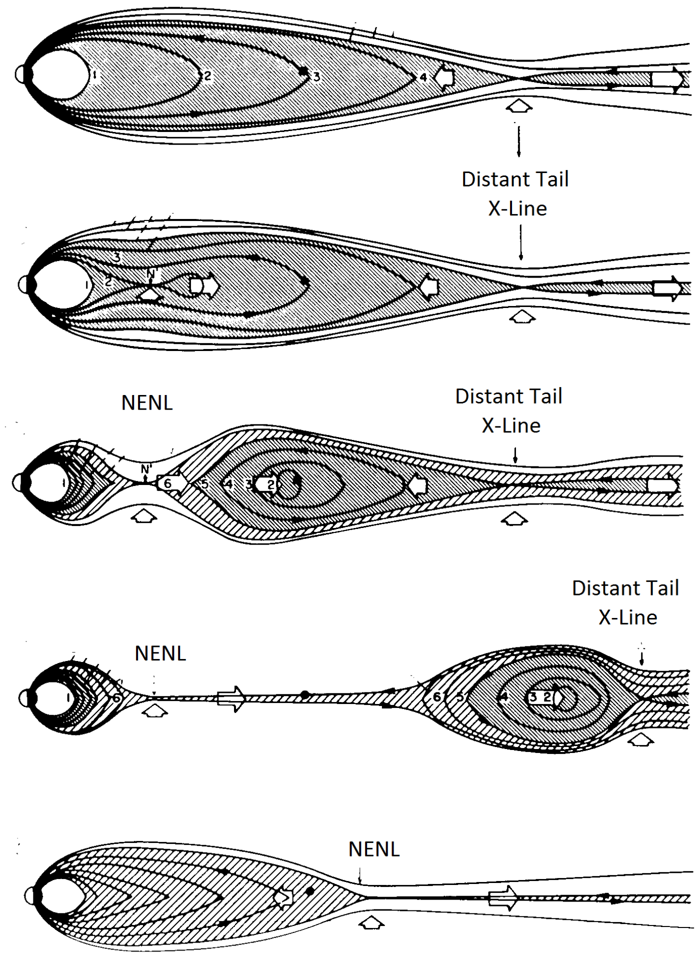
Substorm Overview

Phenomenology



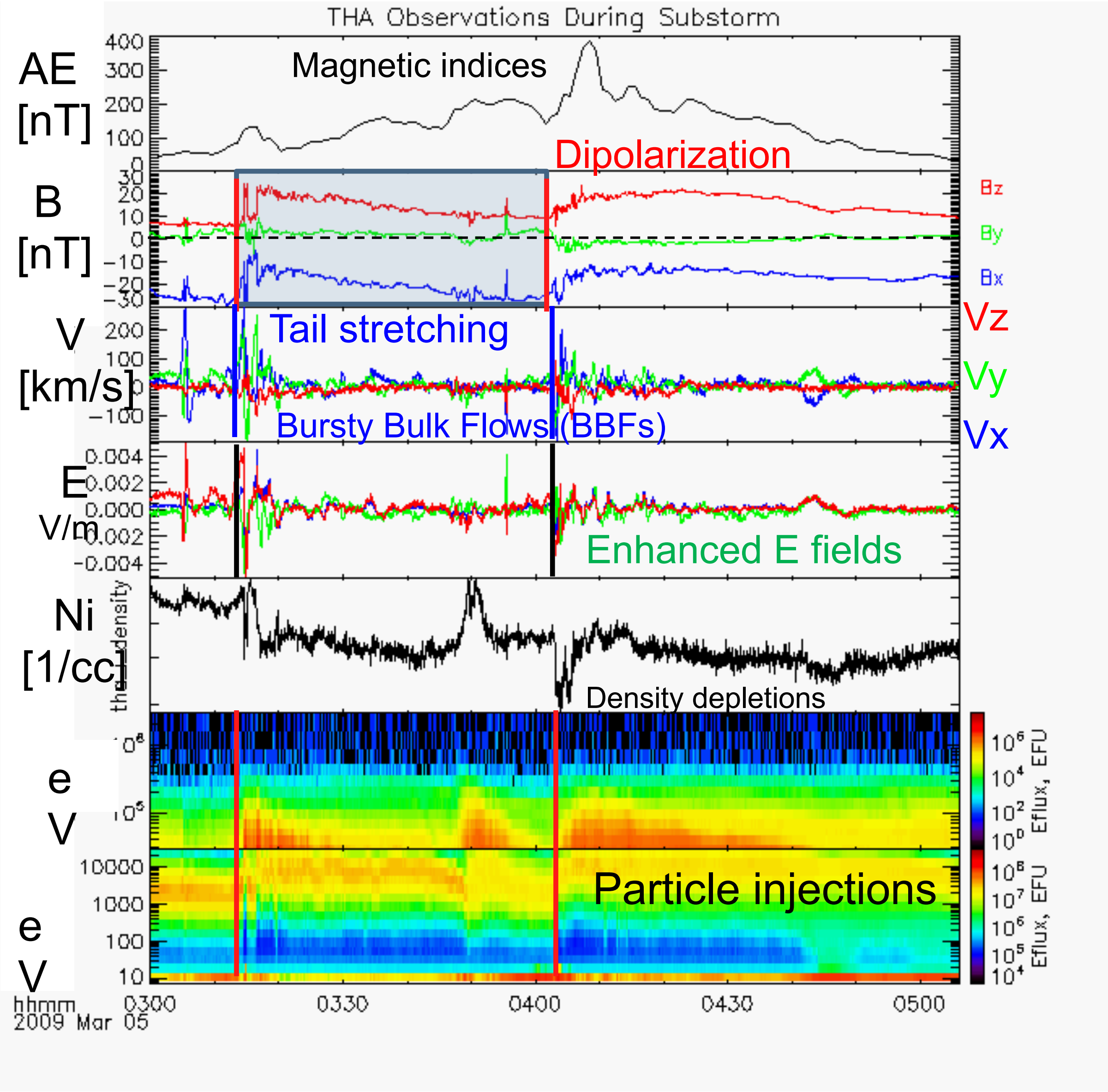
Modified from Hones (AGU JGR, 1977)

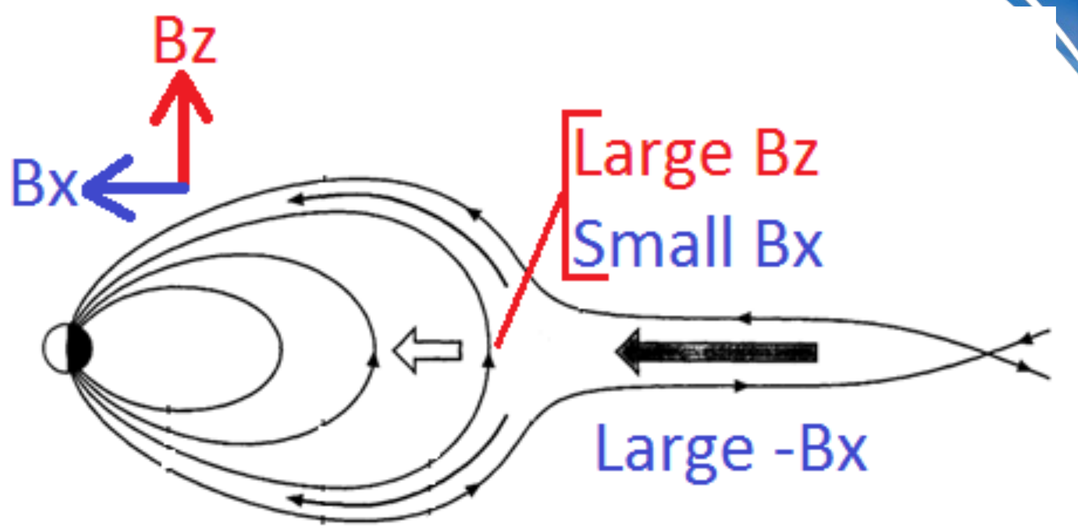
Three Timing Sequences Proposed:

1. Near-Earth Neutral Line (NENL)
    1. Tail stretching & additional load→reconnection ~20-30RE
    2. Earthward flows transport magnetic flux
    3. Large-scale dipolarization and substorm current wedge form ~6.6-12 RE
    4. Hones et al., 1973; 1977; Nice review: McPherron et al. 2020
2. Current Disruption (CD)
    1. Tail stretching thins current sheet to such extent that instability forms ~9-12 RE
    2. Current is diverted along field lines into ionosphere (substorm current wedge), dipolarization forms
    3. Triggers reconnection and/or earthward flows
    4. Lopez et al., 1990; Lui et al., 1991; Lui et al., 2011
3. Streamer-triggered substorm
    1. Flow forming at distant X-line travels earthward (observed optically as a streamer)
    2. Reaches thin, unstable current sheet near 9-12 RE
    3. Causes instability resulting in substorm
    4. Nishimura et al., 2011; 2013; Nice review: McPherron et al. 2020

Substorm Overview

Magnetic field reconfiguration: In situ observations





Often corresponds to:

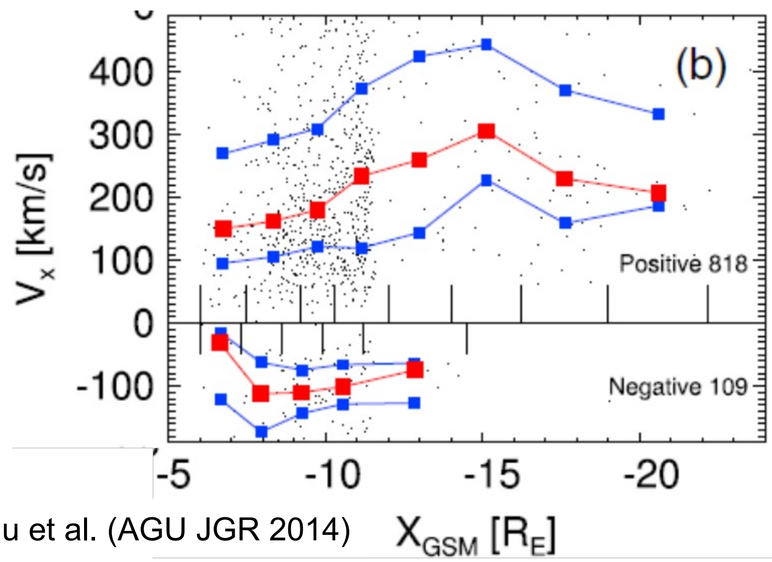
- Substorm (AE index)
- Fast flows (>400 km/s)
- Strong electric fields (2-10s mV/m)
- Plasma density depletions
- Particle injection (flux increase)

https://gem.epss.ucla.edu/mediawiki/index.php/GEM_Tutorials#2017_Summer_Workshop

Substorm: Convection on Mesoscales

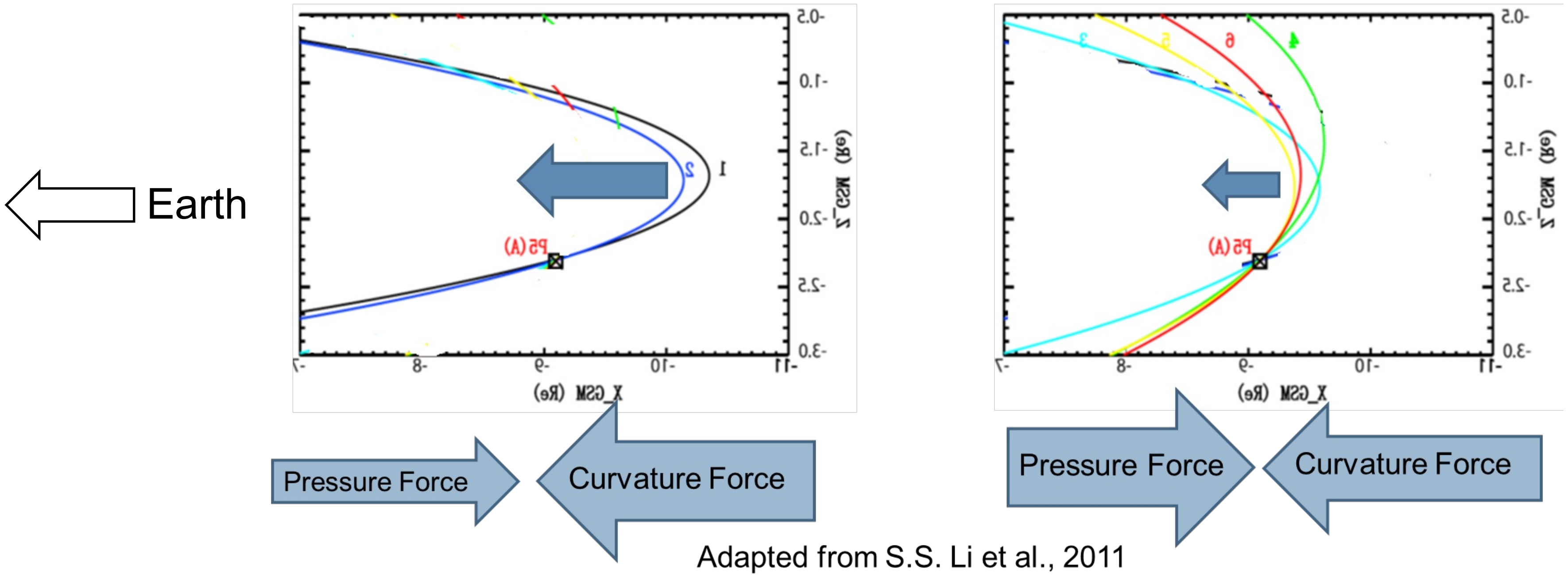
Magnetotail Transients: In terms of local forces

1. Reconnection
2. Curvature force of field line > Pressure force from dipole
3. Flux tube accelerates earthward
4. Pressure force increases, flux tube brakes



Liu et al. (AGU JGR 2014)

Curvature gradient force vs. Pressure gradient force



Adapted from S.S. Li et al., 2011

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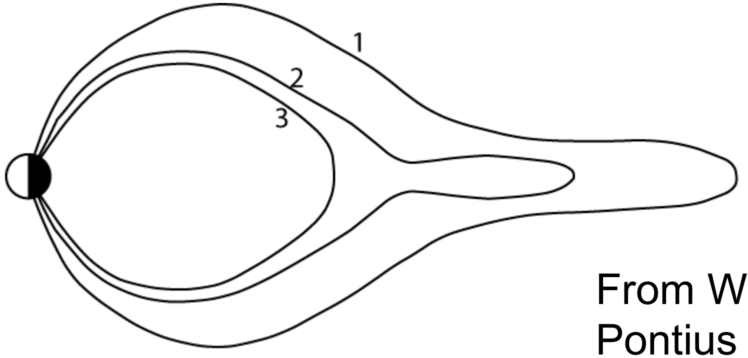
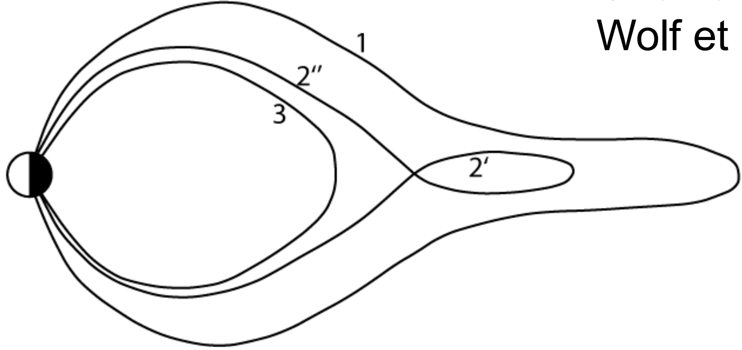
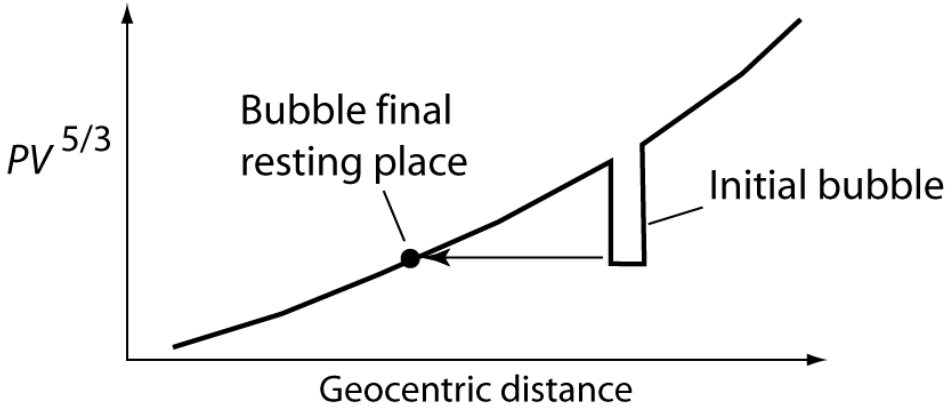
Substorm: Convection on Mesoscales

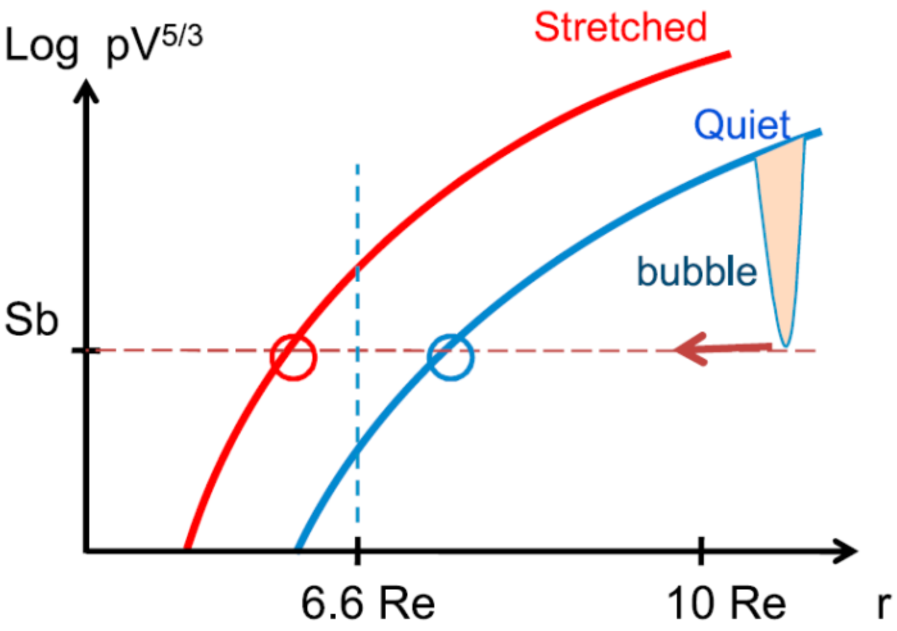
Magnetotail Transients: In terms of entropy





From Wolf GEM Tutorial 2012
Pontius and Wolf, 1990
Chen and Wolf, 1993; 1999
Wolf et al., 2009





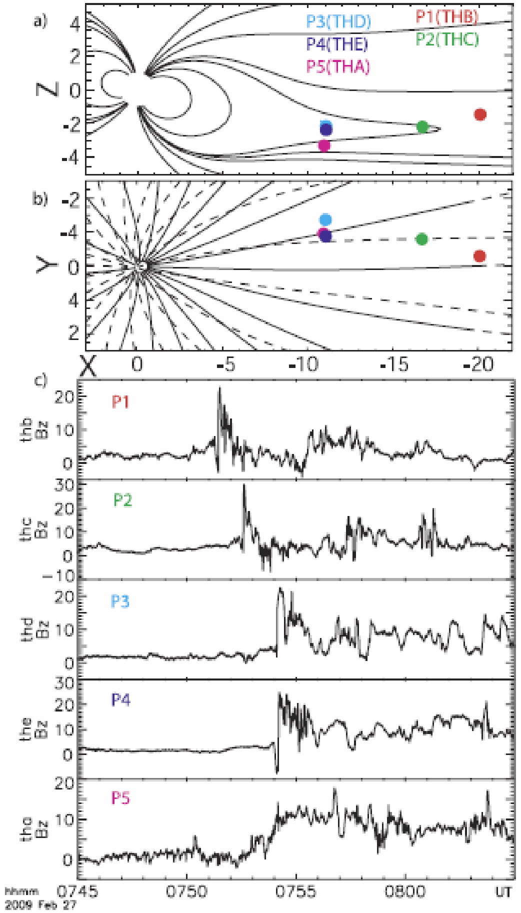
Sergeev et al. (AGU JGR 2012)

Low entropy plasma bubbles

- Global description
- Formulated in terms of Rayleigh-Taylor Instability
- GEO penetration depends on tail preconditioning (Sergeev et al., 2012)

Magnetic Flux Transport

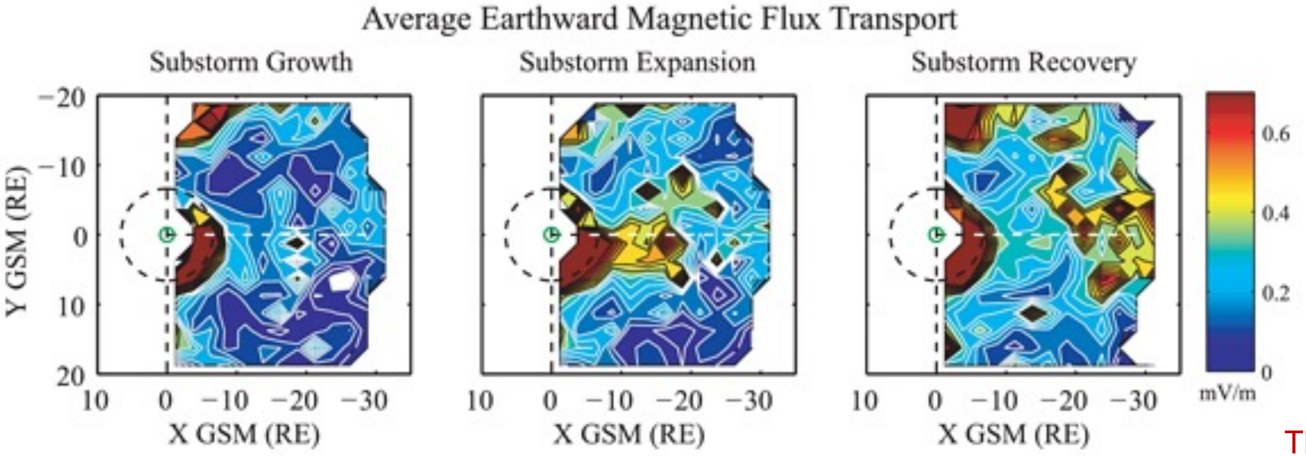
Most during substorm expansion phase



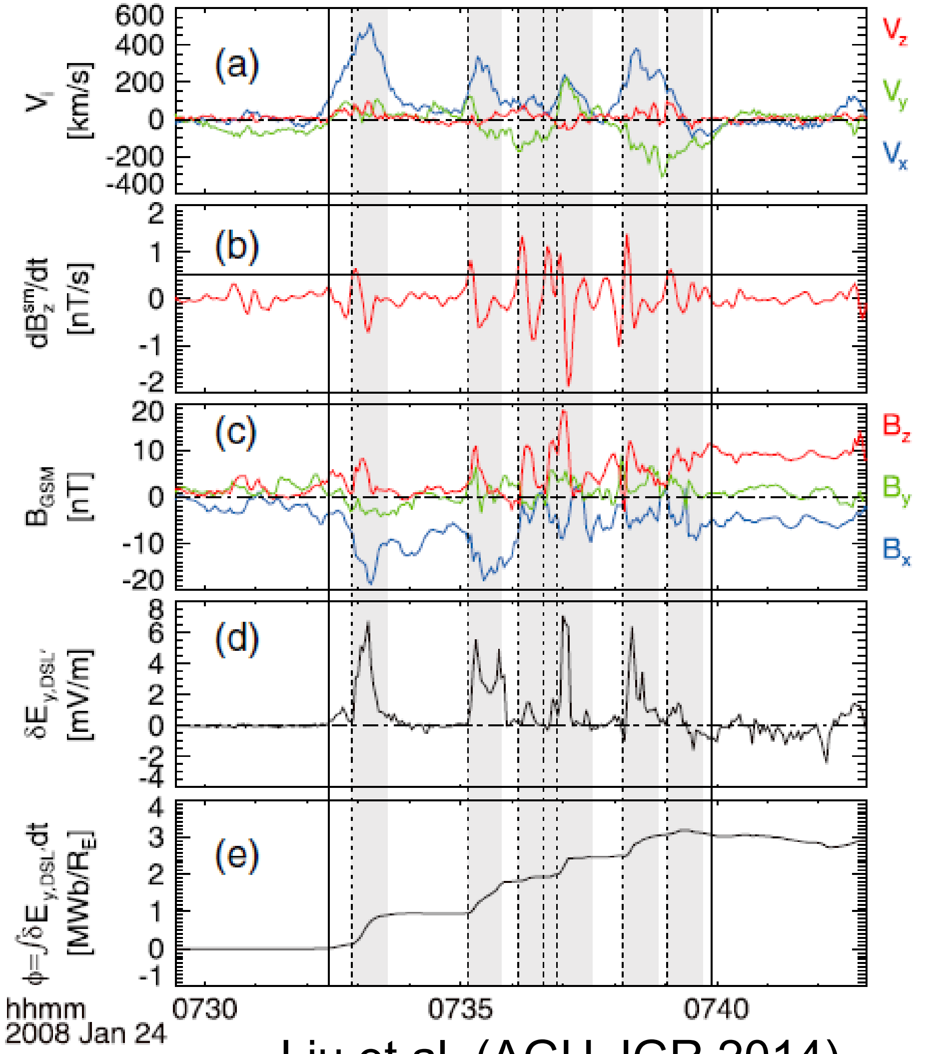
Runov et al. (AGU GRL 2009)

- **>60% magnetic flux carried in BBFs** (Angelopoulos et al., 1992; 1994)
- **~70% of BBF flux transport within dipolarizing flux bundles (DFBs)** (Liu et al., 2014)
- **Most magnetic flux transport occurs during substorm expansion phase** (e.g., Kissinger et al., 2012; Lyons et al., 2012; Merkin et al., 2020)

Kissinger et al. (AGU JGR 2012)



Multiple DFBs in a Bursty Bulk Flow (BBF)

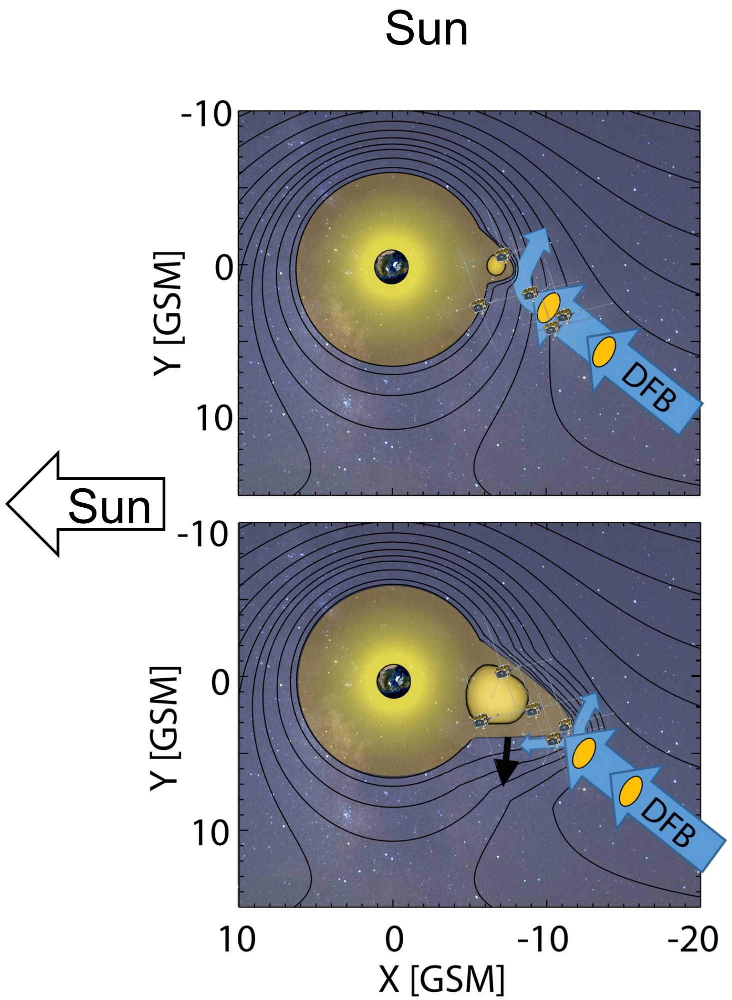


Liu et al. (AGU JGR 2014)

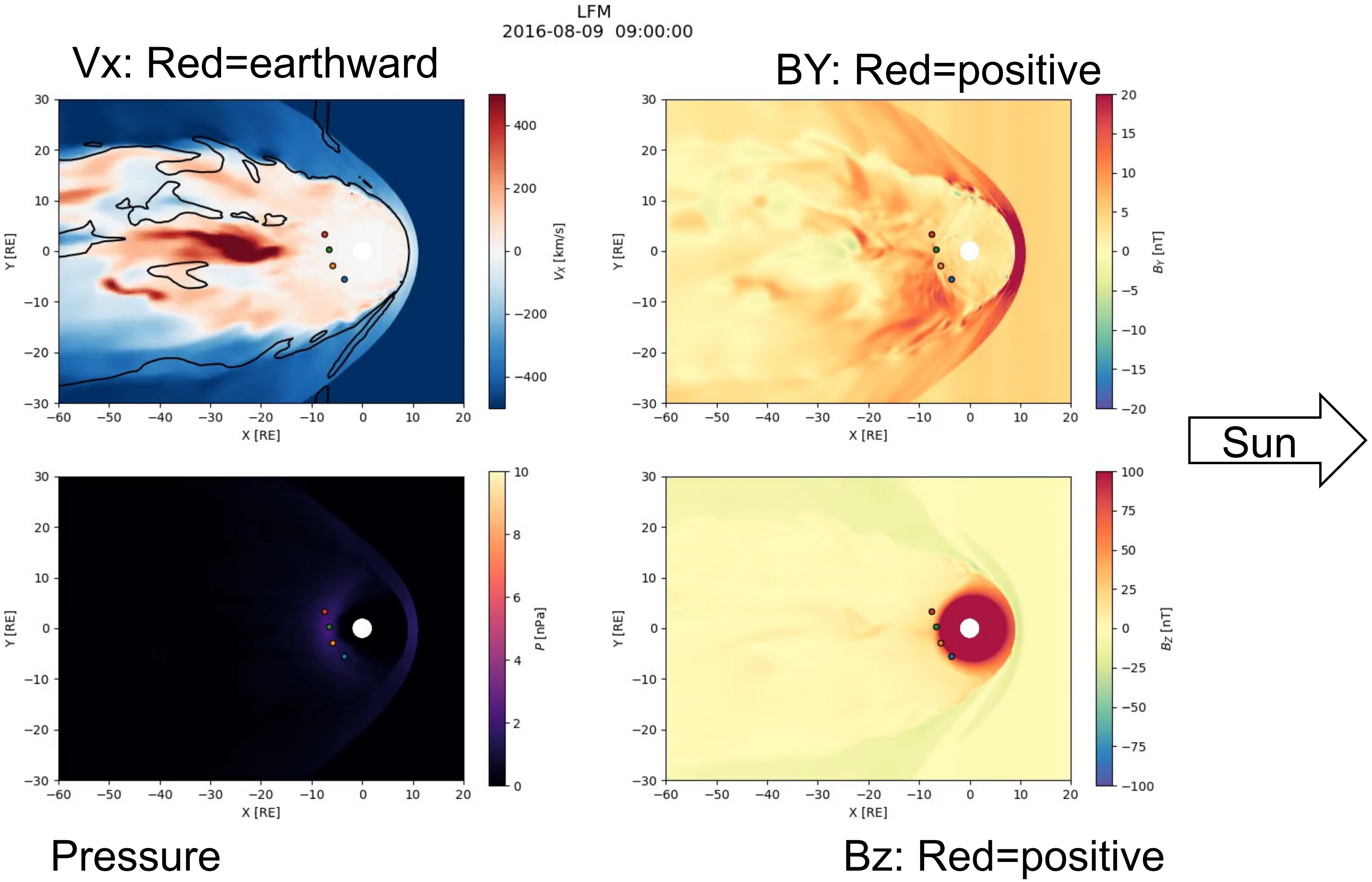
Substorm Dipolarization

Magnetic Flux Pileup

Shiokawa et al., 1997; Baumjohann et al., 1999; Baumjohann, 2002; Nakamura et al., 2009; 2013



Modified from Gabrielse et al. (AGU JGR 2019)
Observations: Bz, Vx, and Injections



Merkin et al. (AGU JGR 2019)
Modeling: Bz and Vx

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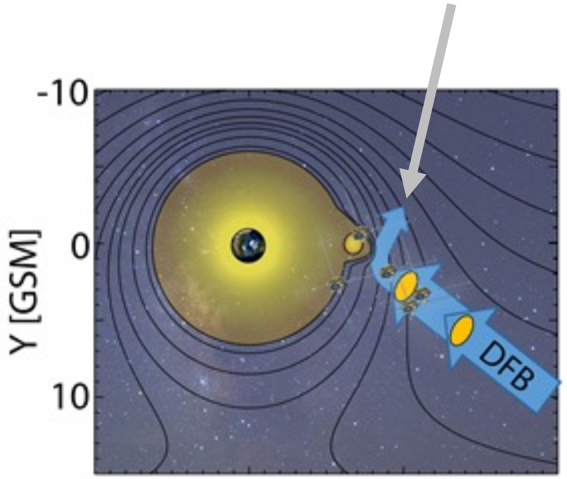
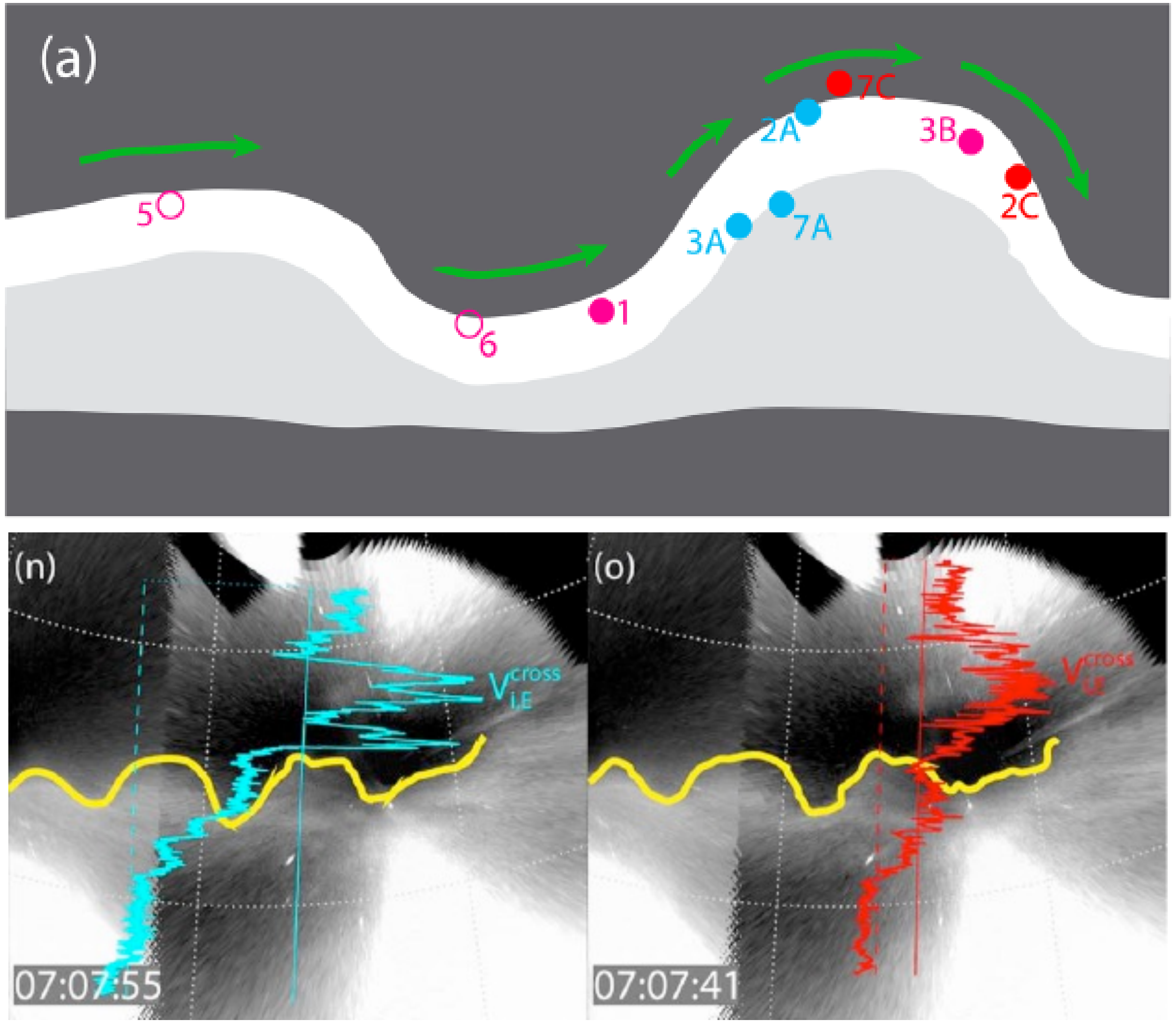
Diverted Flows



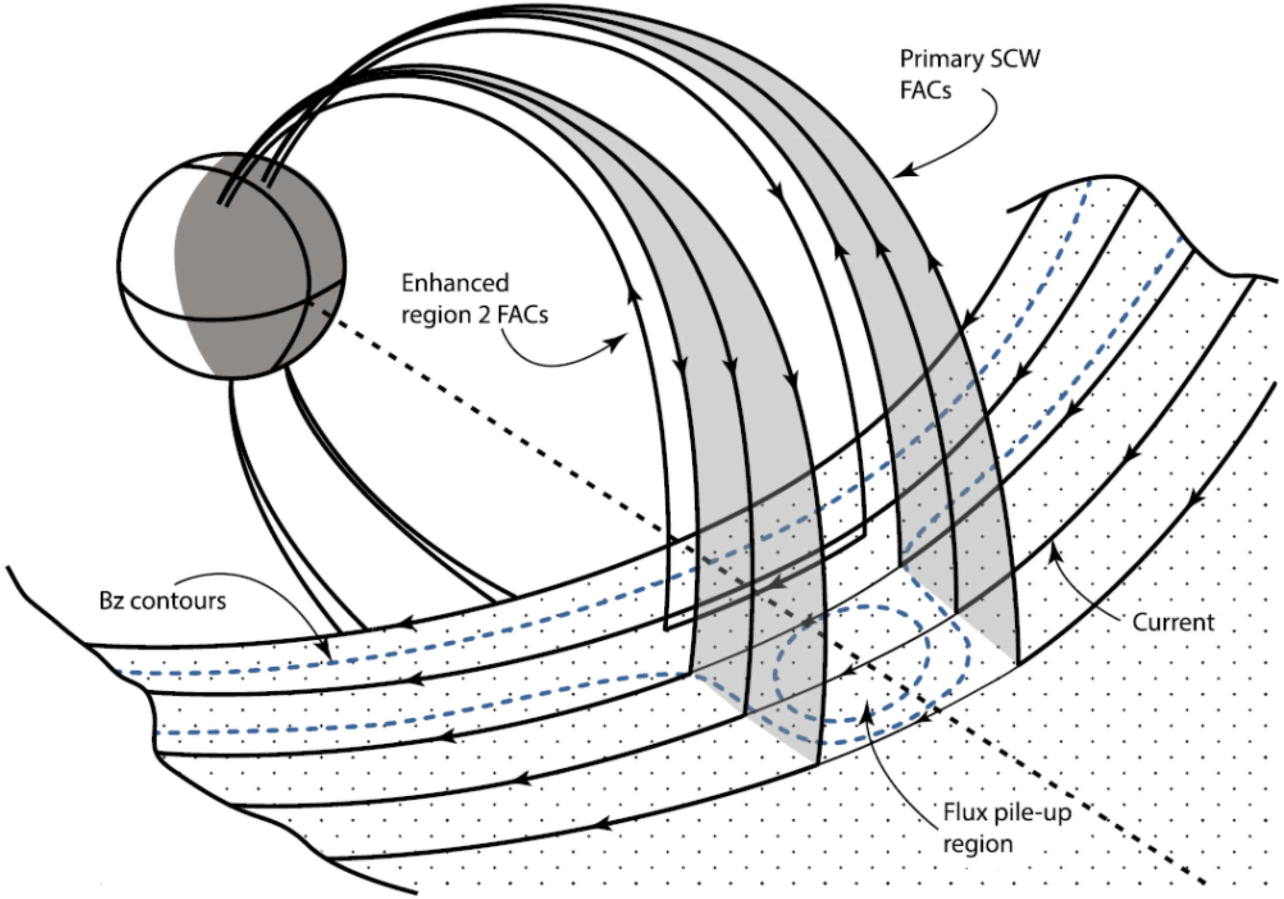
SWARM and THEMIS ASIs used to correlate diverted, fast flows to omega bands post-midnight: Kelvin-Helmholtz



Modified from Liu et al. (AGU GRL 2018)

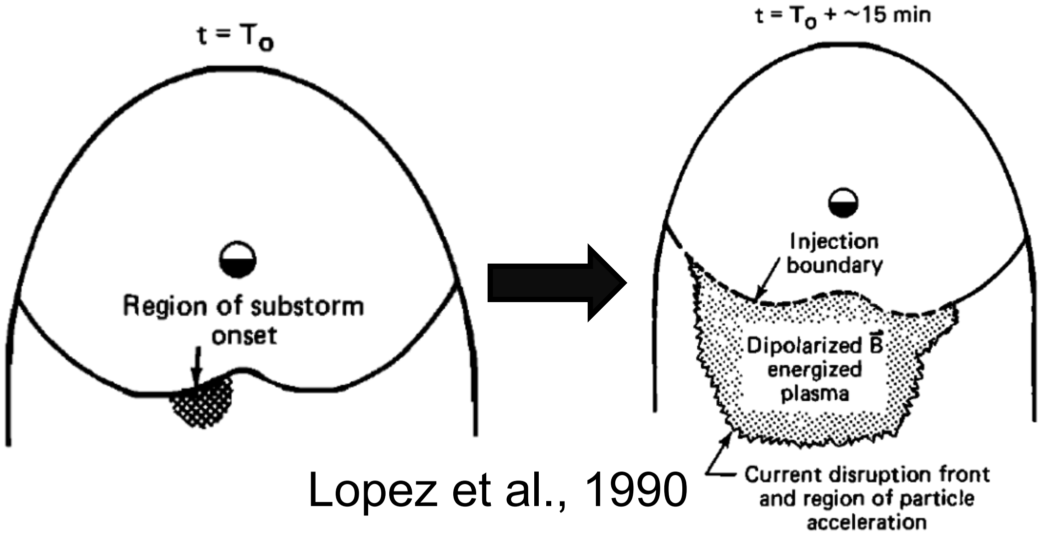
Substorm Current Wedge

A result –or cause— of large-scale dipolarization



Dipolarization diverts current: Kepko et al. (Space Science Reviews, 2014)

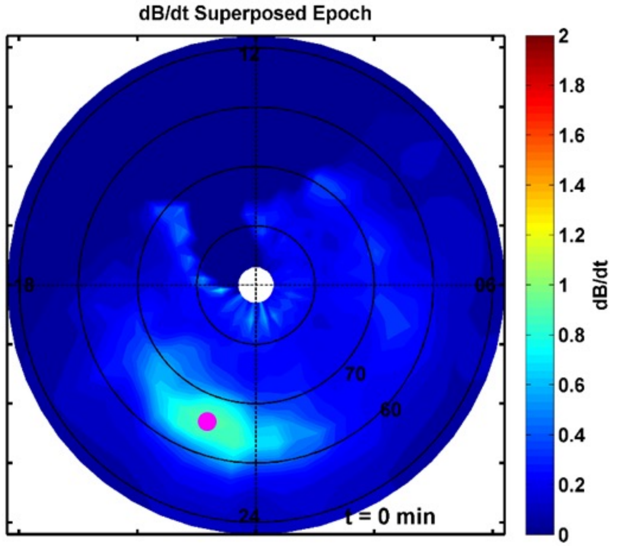
Current disruption: Instabilities form when current sheet becomes too thin.



Lopez et al., 1990

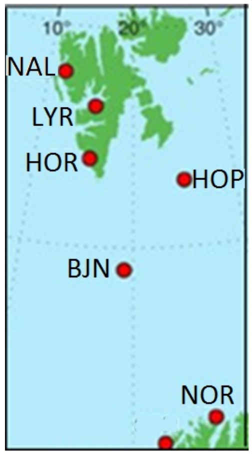
Substorm Current Wedge

Magnetic Indices

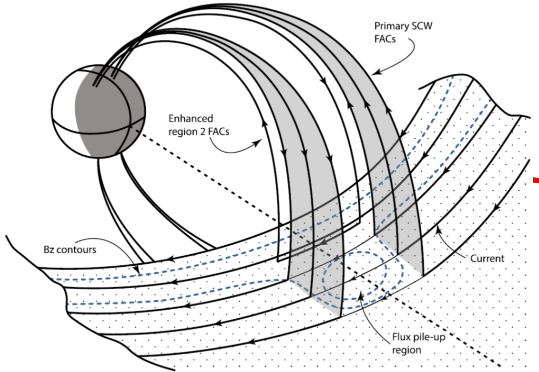


From James Weygand

Where on this map is the SCW closing in the ionosphere?

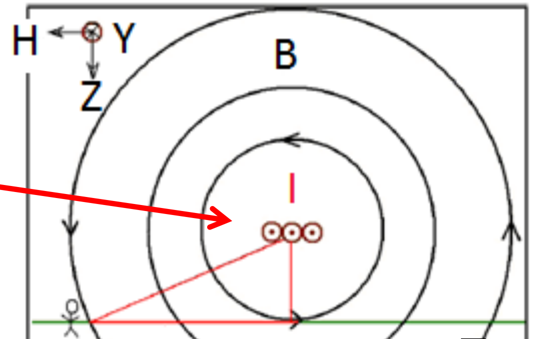


Kepko et al. (Space Science Reviews, 2014)



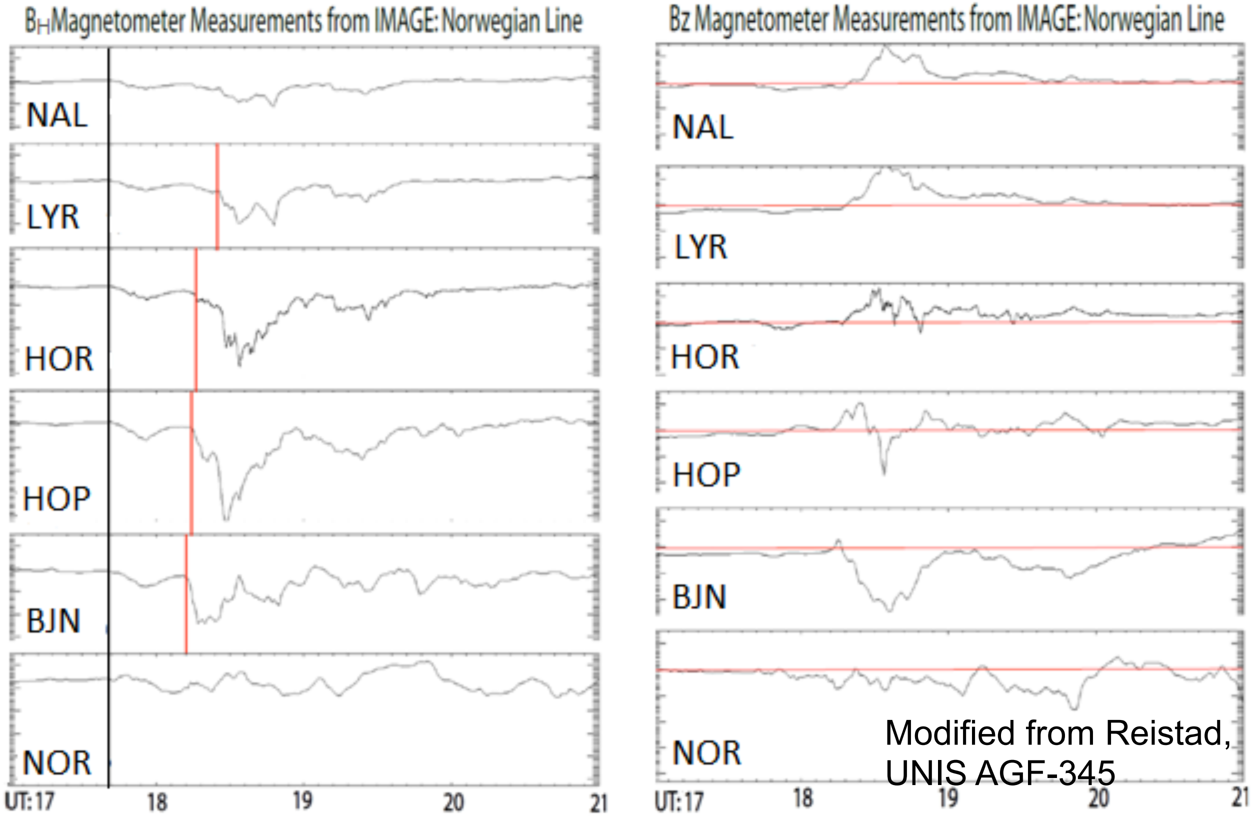
Substorm Current Wedge

Ampere's Law



<-Pole

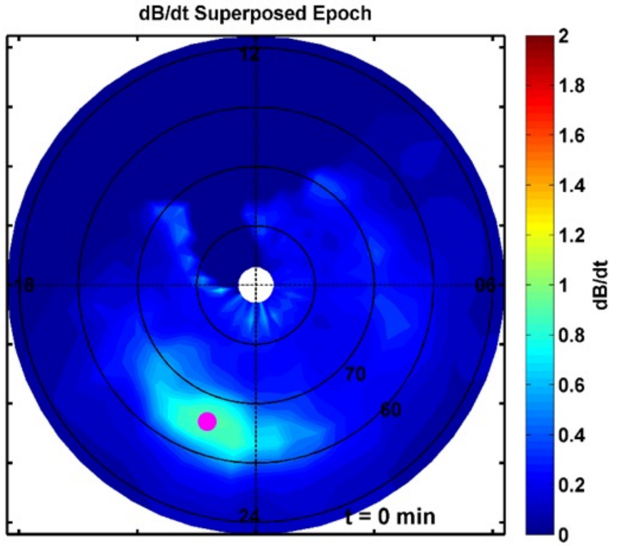
Equator->



Modified from Reistad, UNIS AGF-345

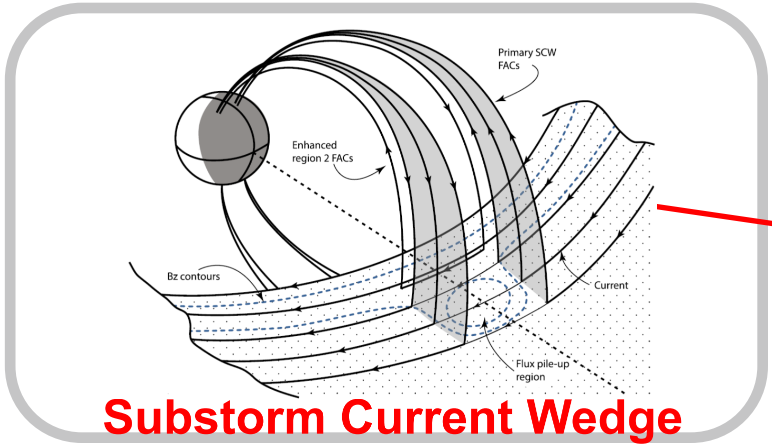
Substorm Current Wedge

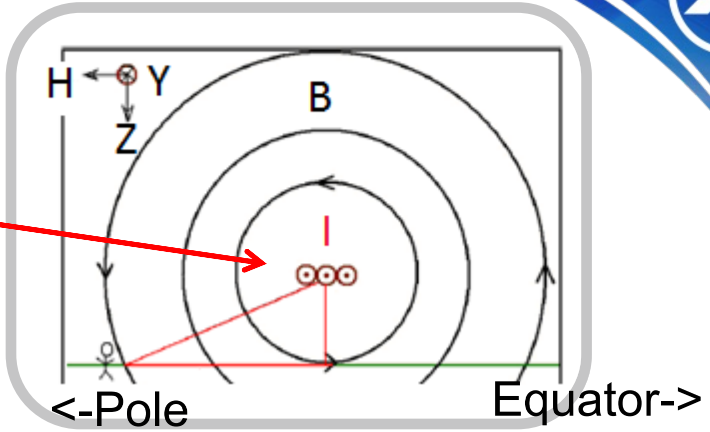
Magnetic Indices

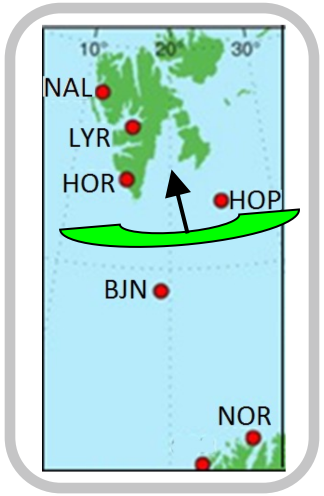


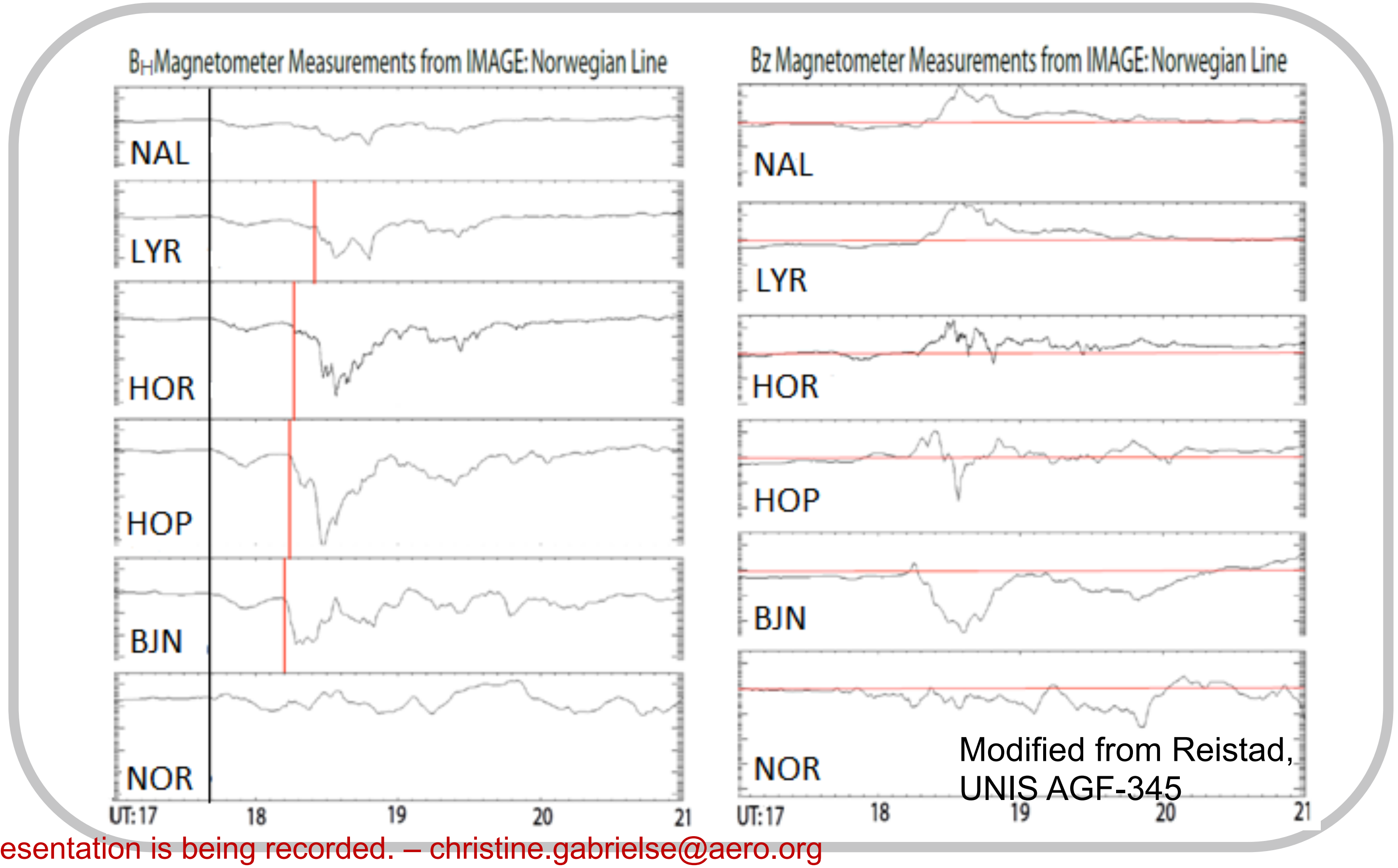
From James Weygand

Kepko et al. (Space Science Reviews, 2014)







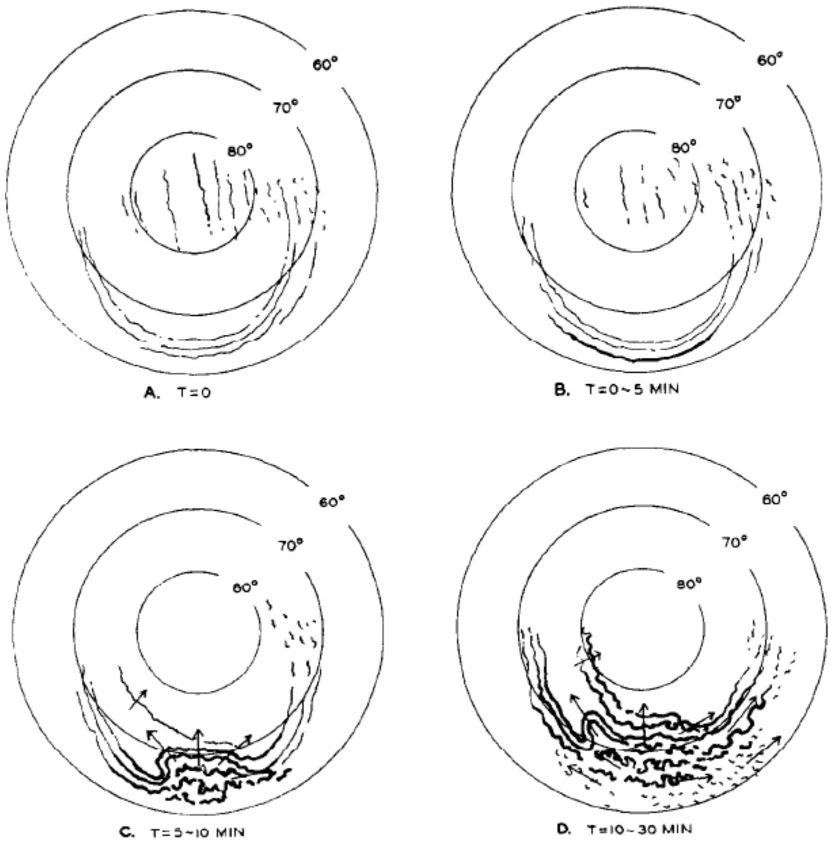


Modified from Reistad, UNIS AGF-345

Substorm Overview

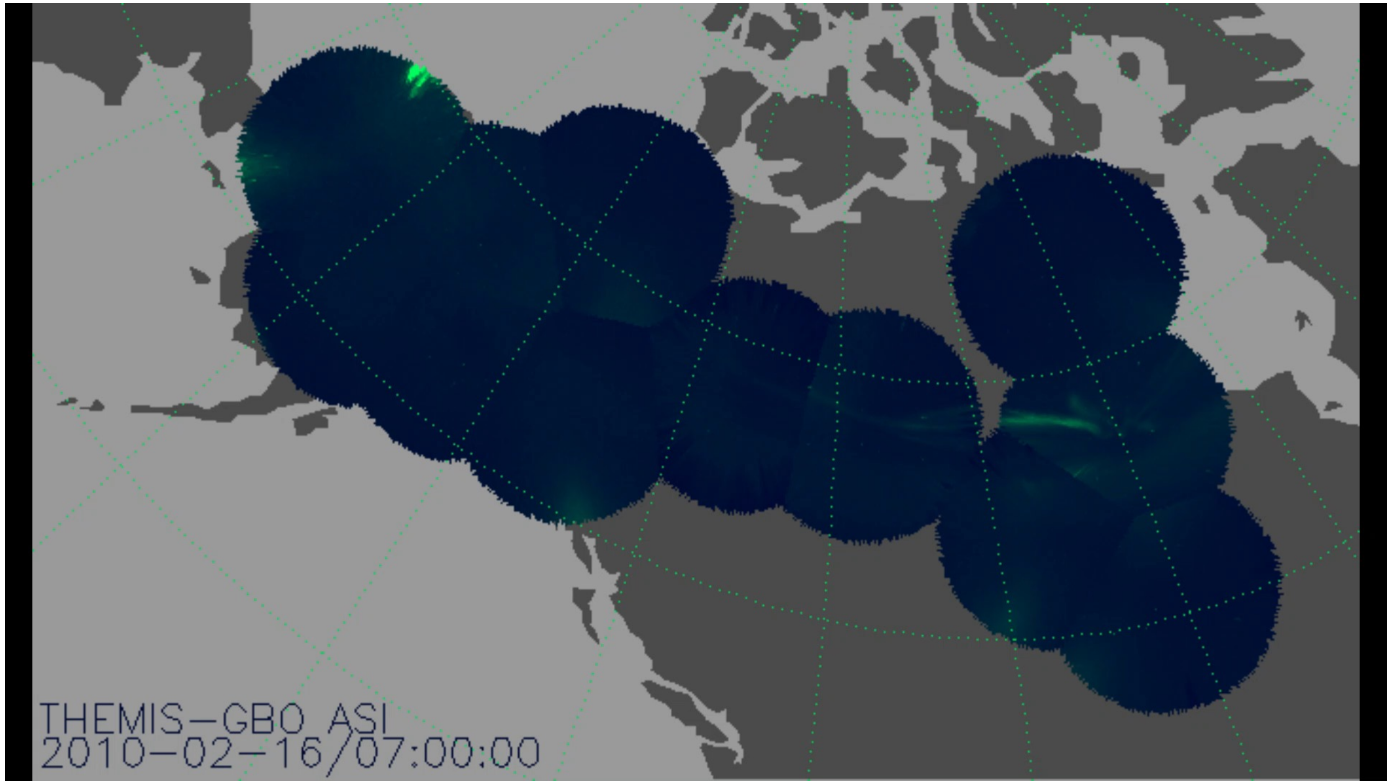
Auroral Onset: The OG Substorm Onset Definition

Auroral Onset



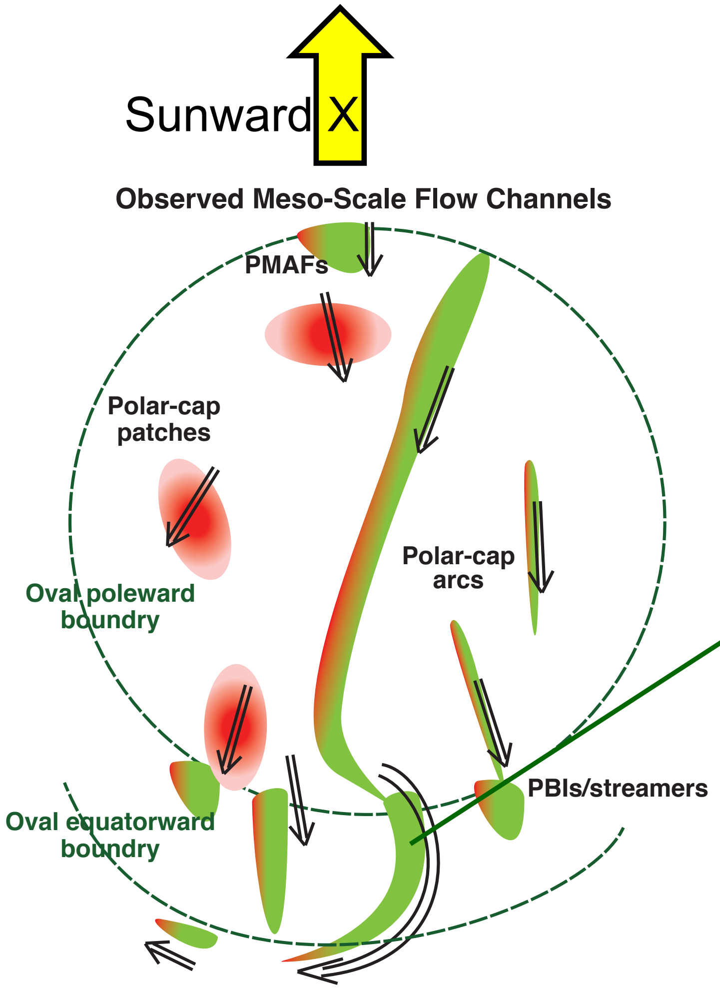
Adapted from Akasofu
(Planetary and Space Sciences, 1964)
886—1578 citations

THEMIS white light All-Sky-Imagers (false color)



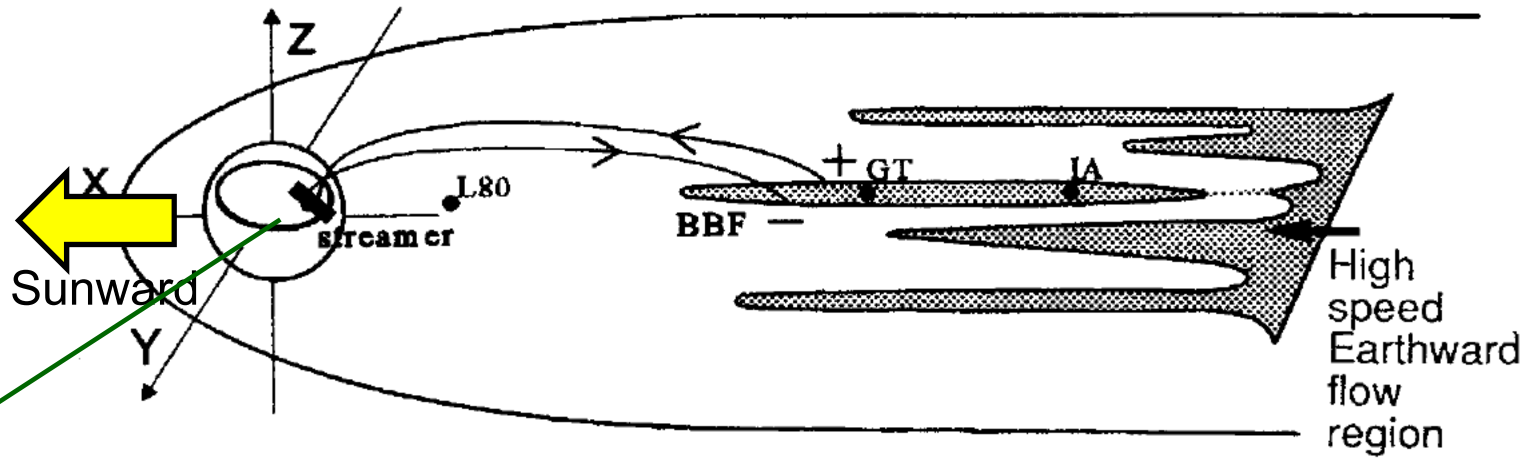
Convection on Mesoscales

Magnetotail Flows and the Ionosphere



Modified from Lyons al. [AGU JGR 2016]

Sergeev et al. [AGU GRL 2000]



Convection on mesoscales can be studied in the magnetosphere by satellites or in the ionosphere by low-earth orbiting satellites or ground-based radar.

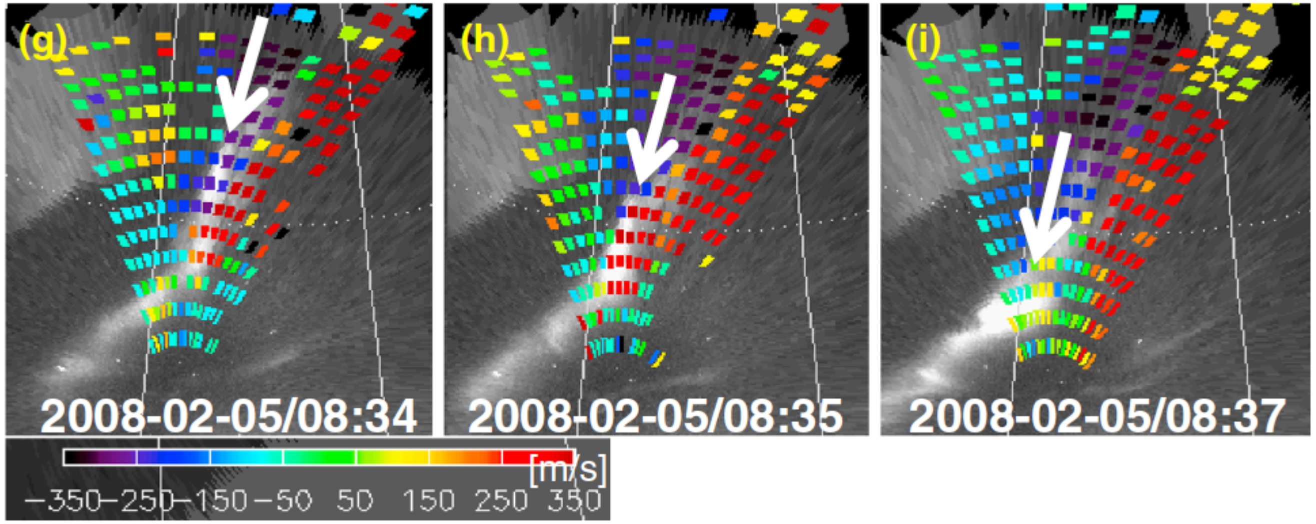
Although, McPherron et al. (2020) found not every flow burst had an associated streamer.

Convection on Mesoscales

Magnetotail Flows and the Ionosphere

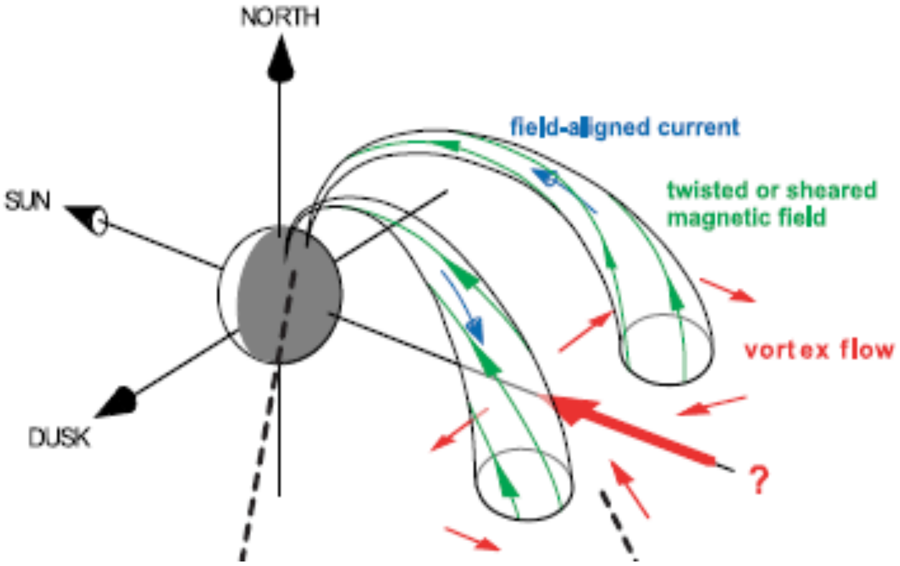
Flow shear from vortex creates field-aligned current, accelerates electrons and creates equatorward-traveling auroral streamers.

Blue=poleward/tailward flow, red=equatorward/earthward flow



Gallardo-Lacourt et al. (AGU JGR, 2014)

Note: The optical signature (streamer) lies west of the equatorward/earthward flow and east of the poleward/tailward flow, right at the flow shear region.



Birn et al. (AGU JGR, 2004)

Read more!
Henderson et al., 1998; Sergeev et al., 1999, 2000; Lyons et al., 1999, 2002; Kauristie et al., 2000; Zesta et al., 2000; Zou et al., 2010, 2013.

Plasma Transport

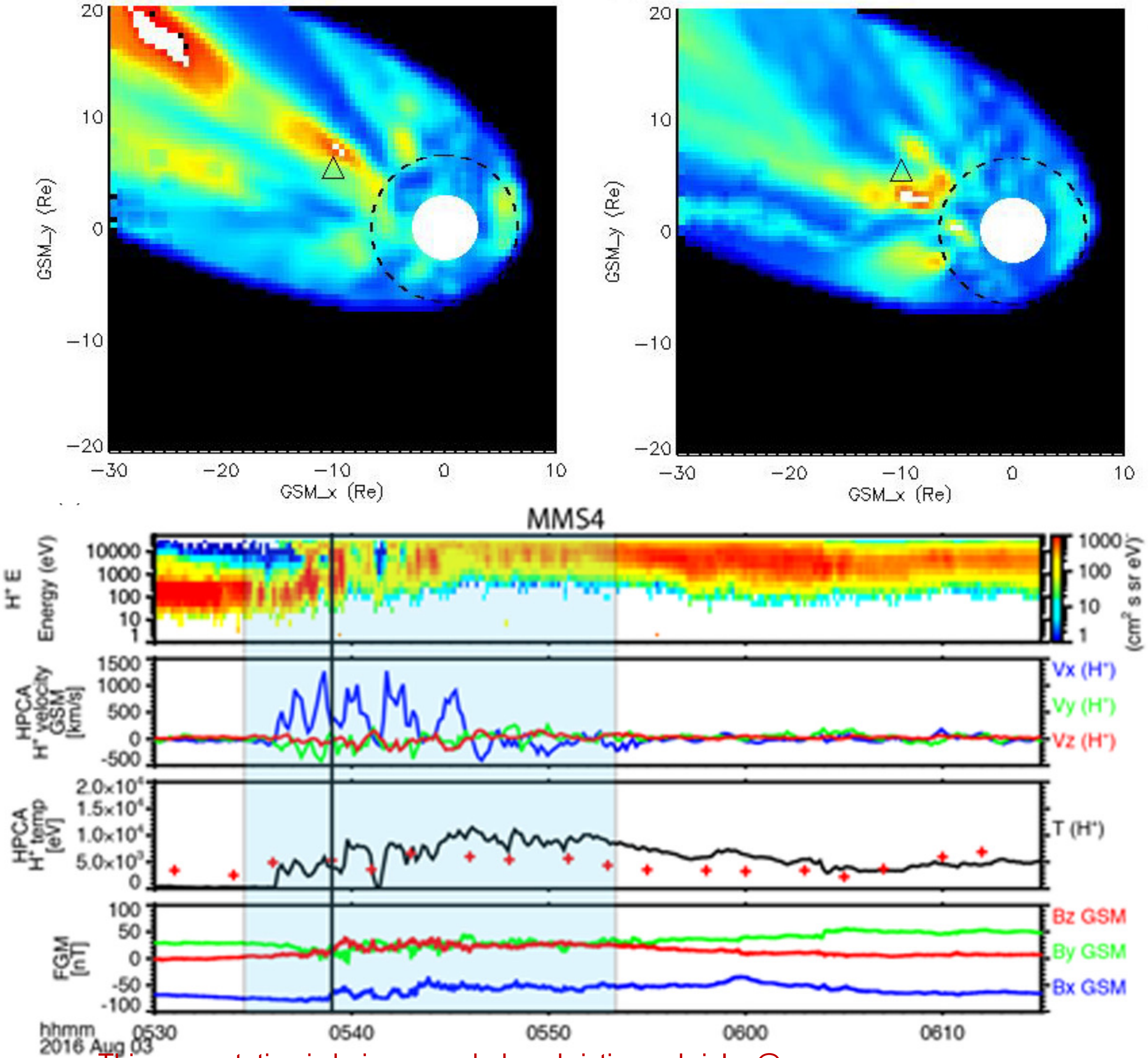
TWINS: 2D map of ion temperature

Compare with in situ data (MMS)

Are particles transported all the way earthward from X-line?

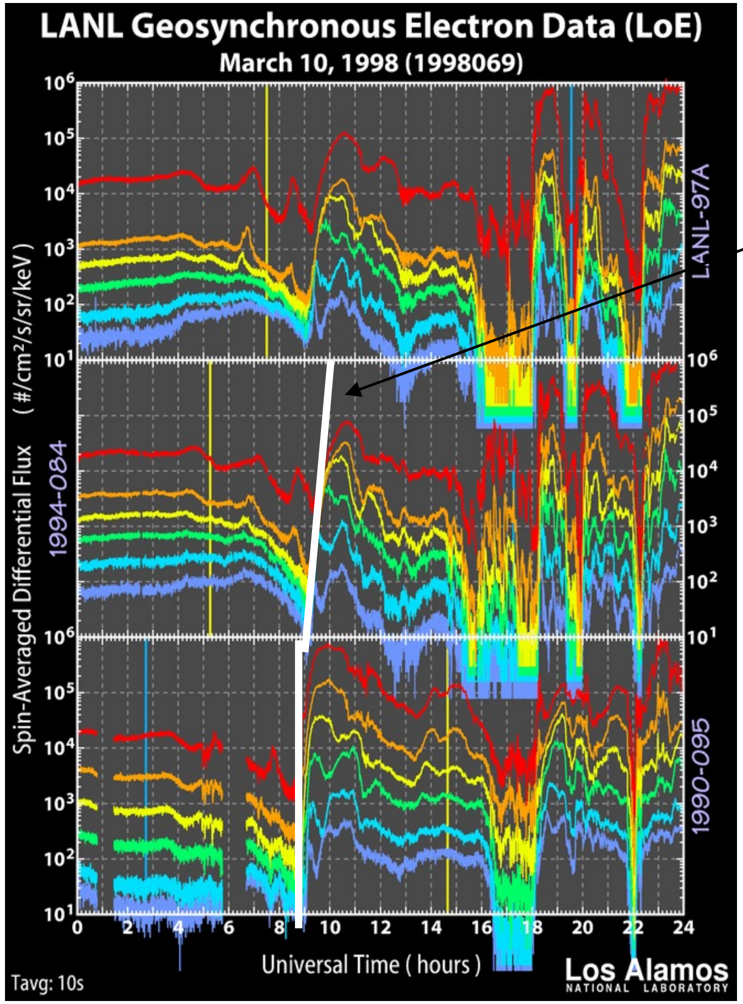
How much energy transferred from X-line to inner magnetosphere?

Keesee, GEM 2020 Presentation (see also Keesee et al., 2014)



Particle Transport

Particle Injections



http://hpde.gsfc.nasa.gov/LWS_Space_Weather/LANL_descrip.html

Reeves et al. (Proceedings of the 3rd International Conference for Substorms,1996)

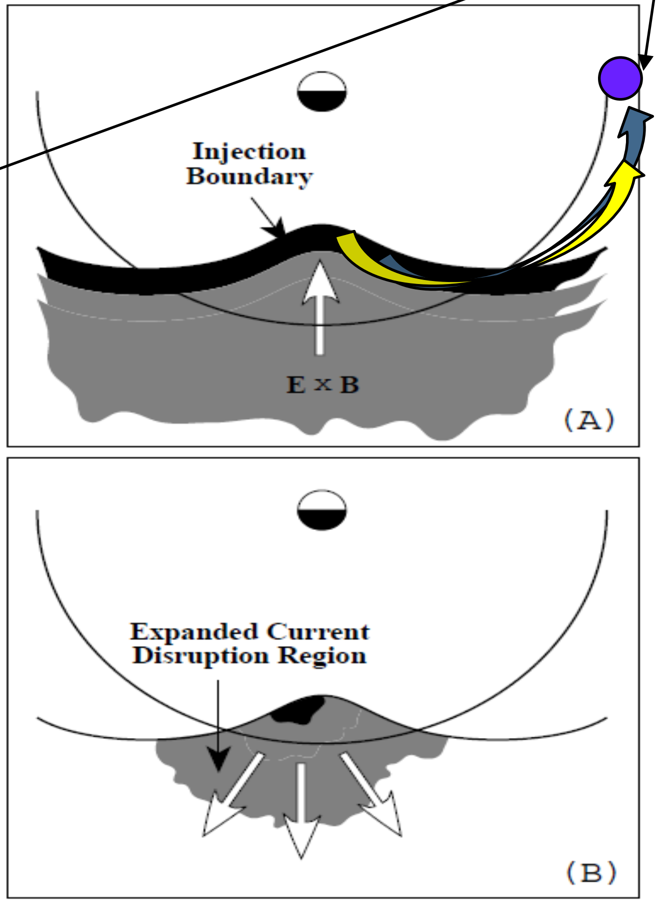


Figure 1: Propagation of the substorm injection region predicted by (A) the Convection Surge model and (B) the Current Disruption model.

$$\mathbf{v}_{drift} = \frac{mv_\perp^2}{2qB}\frac{(-\nabla_\perp B)\times\mathbf{B}}{B^2}$$

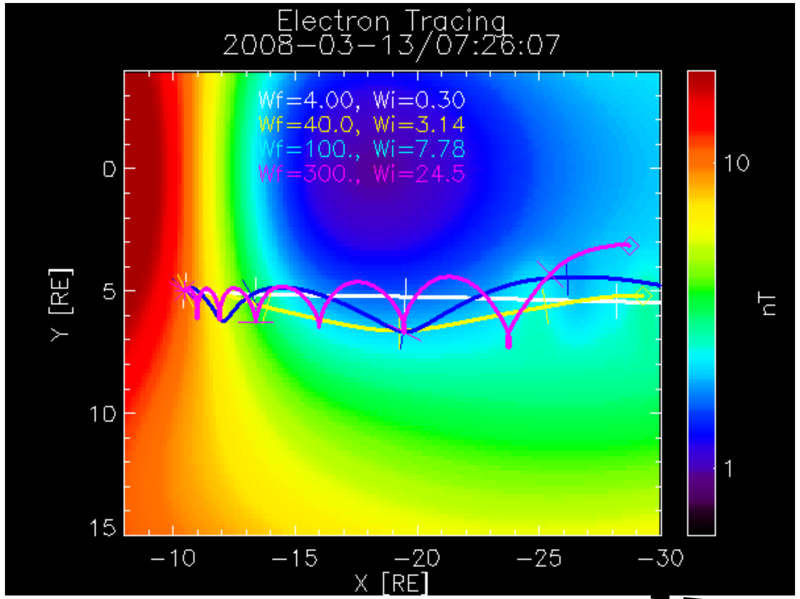
Earthward motion of a boundary between hot and cold plasma

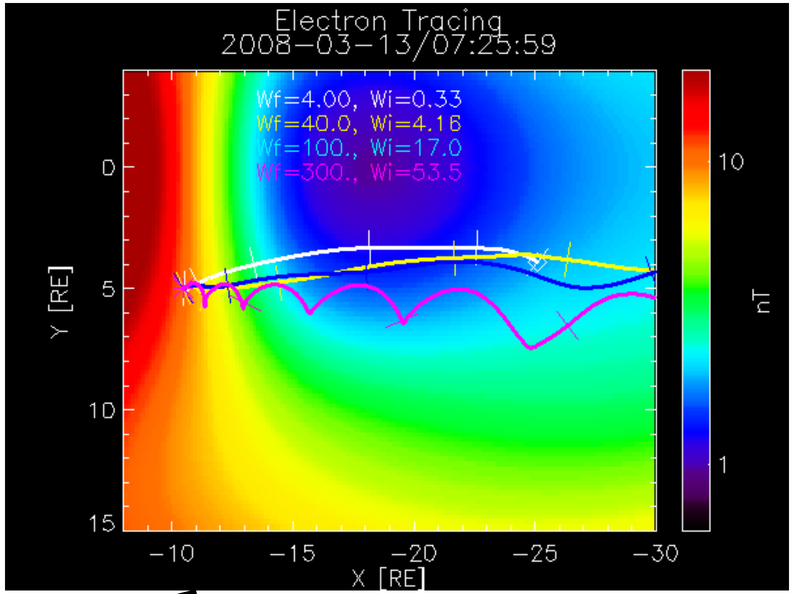
(e.g., Mauk and McIlwain, 1974; Konradi et al., 1975; Mauk and Meng, 1983; Reeves et al., 1990; Birn et al., 1997)

Tailward propagation also observed

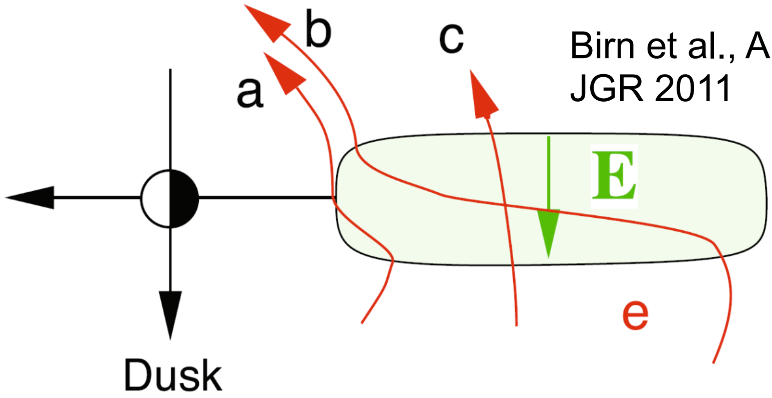
(e.g., Lopez et al., 1990; Spanswick et al., 2010; Gabrielse et al., 2019)

Particle Transport



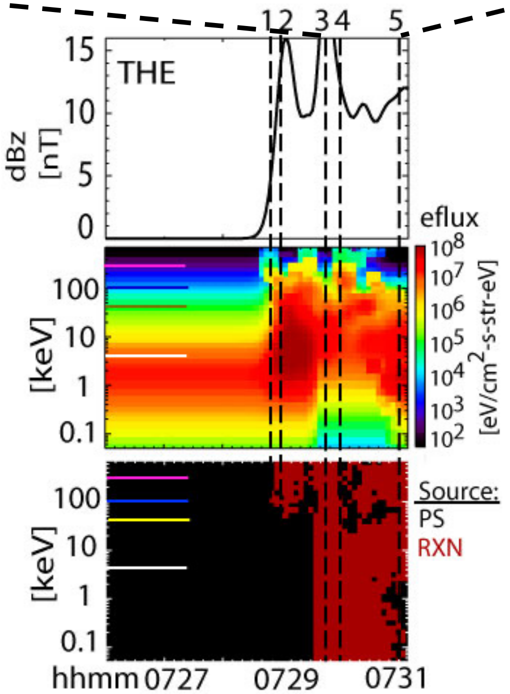


Explains how very localized DFB can energize particles more than what can be gained by drifting across a narrow potential drop.



Birn et al., AGU JGR 2011

Typical picture without localization in X

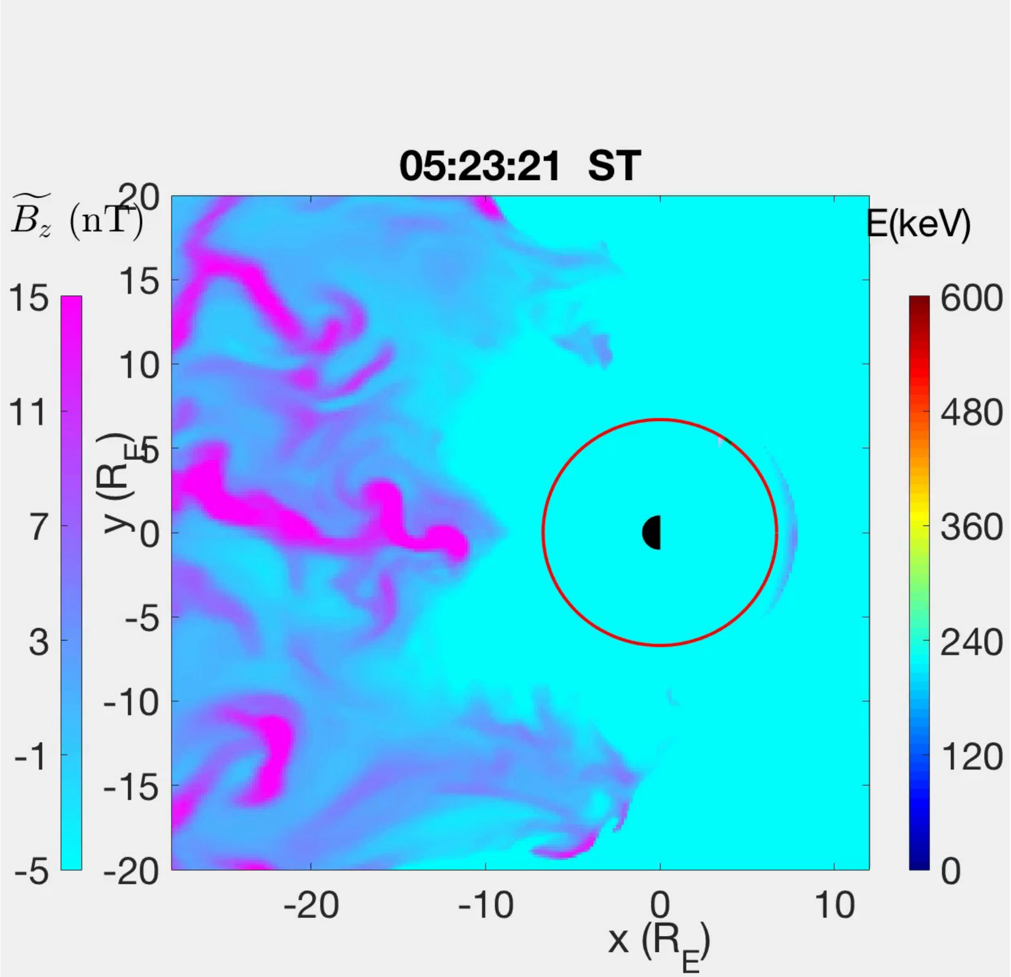


Gabrielse et al. (AGU JGR 2017) for electrons

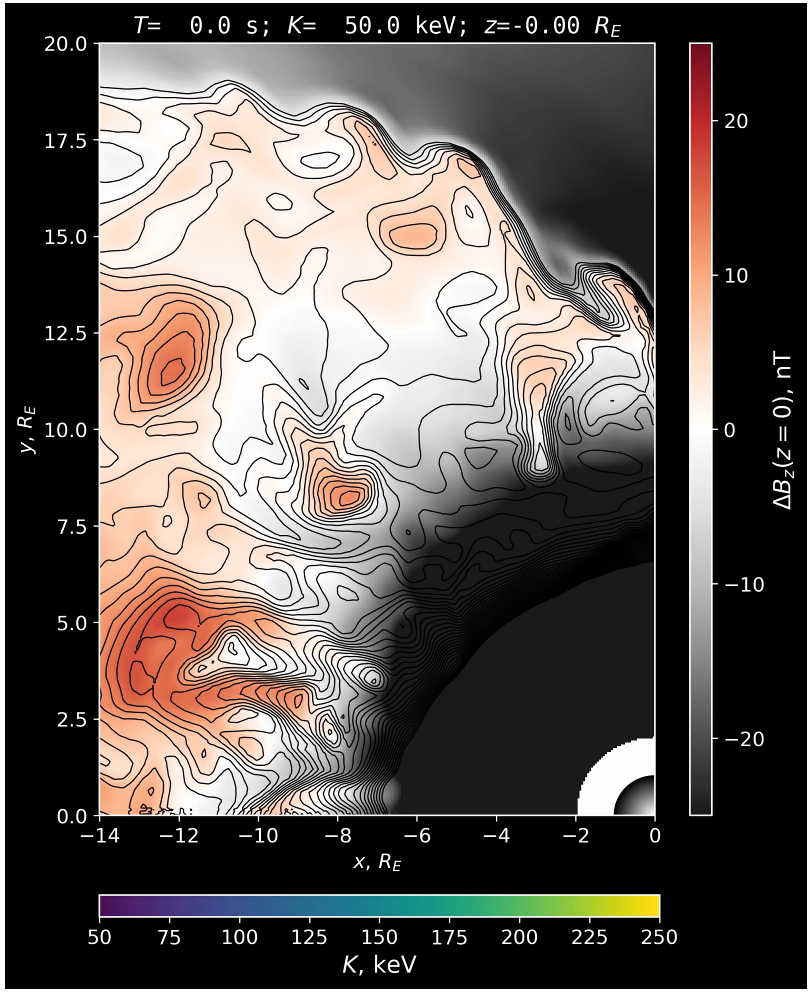
Only considering adiabatic energization, largest source is from dB/dt of the dipolarization front a trapped electron finds itself gradB drifting about.

Particle Transport

Both studies below suggest energization can be non-adiabatic



Eshetu et al. (AGU JGR 2019): Electrons in LFM MHD
See also: Kim et al. (2000) and Sorathia et al. (2018) for electron trapping in MHD fields



Ukhorskiy et al. (AGU JGR 2019): Ions in MHD

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Recent Work and Looking Forward

Convection and Substorms

- **This lecture focused more on IMF -Bz, but IMF +Bz and IMF BY→interhemispheric asymmetries many open questions**
- **Mesoscale convection is important! Especially when discussing substorms.**
- **Working on quantifying mesoscale phenomena contribution**
  - *Newest papers hot off the press:*
    - GEM Focus Group on Magnetotail Dipolarization and its Effects in the Inner Magnetosphere: bit.ly/DIPFG
- **Quantify contributions to radiation belts and ring current**
- **Vast implications on the ionosphere-thermosphere-mesosphere**
  - *Global models (e.g., GITM) now computationally capable of including mesoscales (flows, Efields, precipitation, etc.)*
  - *Models need data inputs to inform how to run!*
  - *Feedback has implications on magnetosphere*
- **NASA is taking White Papers for its Decadal Survey. I repeat other speakers' call for multi-point observations**
  - *As you saw, hard to visualize with single data points in space. Decadal is our chance to influence NASA on what science is important to study.*
  - *Missions that spread in azimuth (MLT) and radial distance*
  - *Utilize imaging and 2D datasets*
  - *Observe aurora/precipitation simultaneously to in situ particle data*
  - *Coordinate with improving modeling efforts that include smaller scales*

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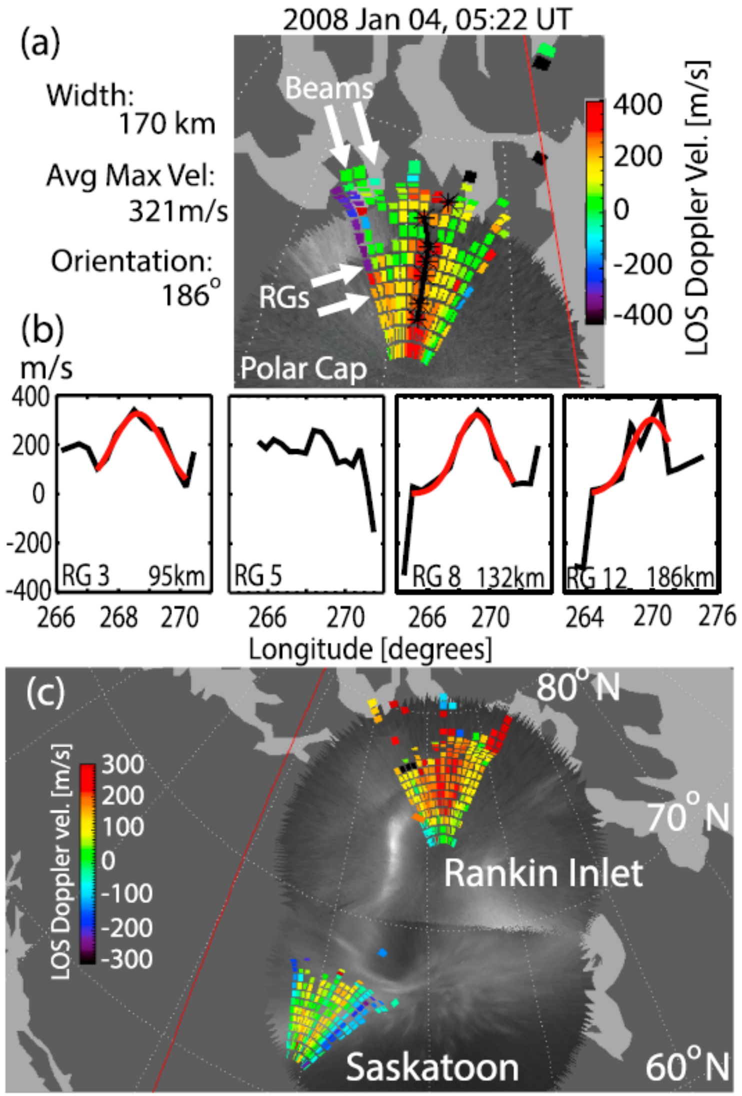
Thank you!

(Back Up)

Convection on Mesoscales

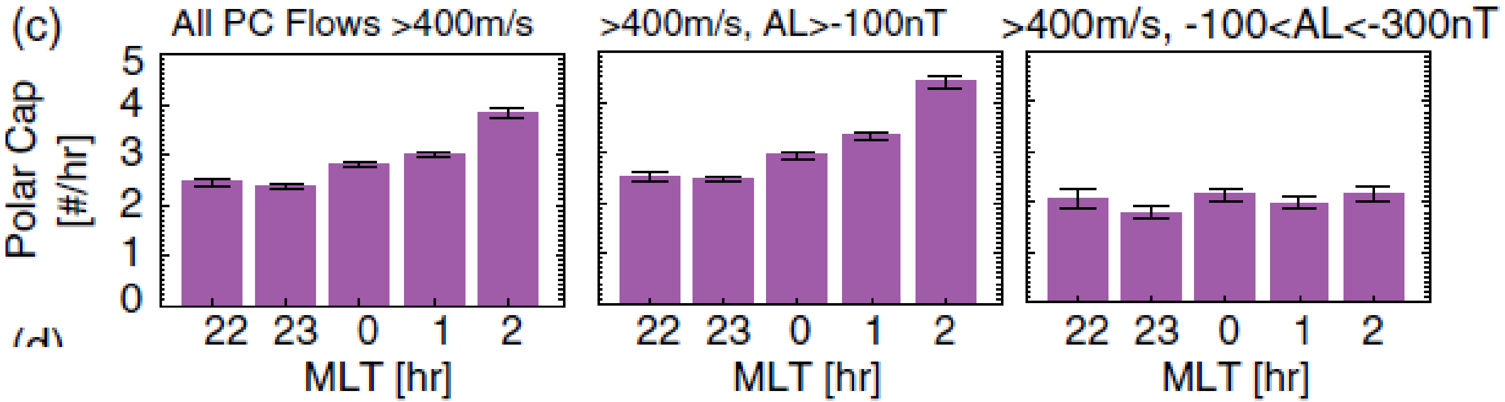
Polar Cap Convection

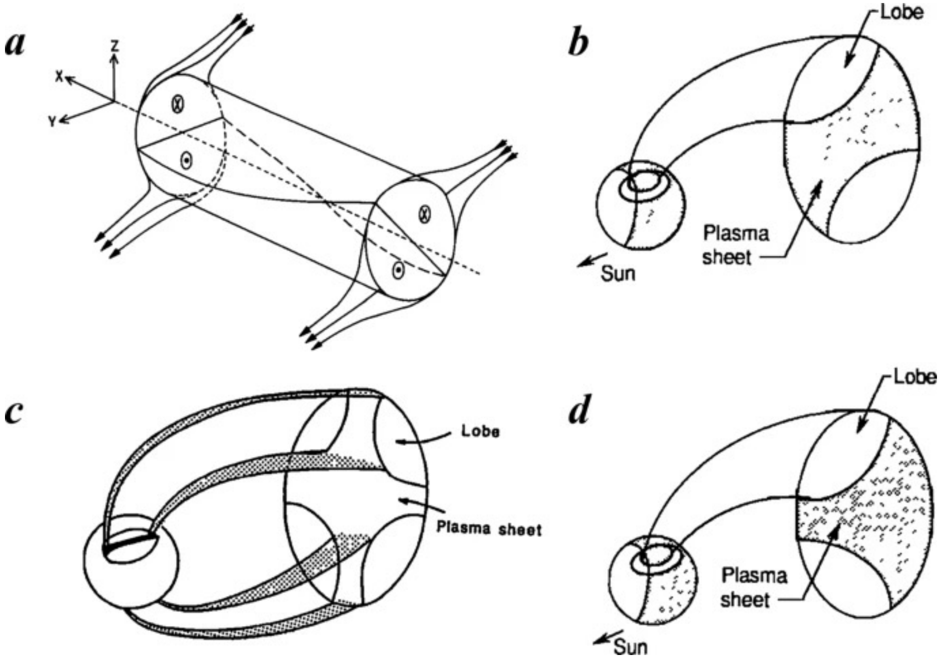


Gabrielse et al. (AGU JGR, 2018)

Fewer polar cap mesoscale flows during substorm.

Similar trend as polar cap arcs: sun-aligned auroral arcs in polar cap (Berkey et al., 1976; Hosokawa et al., 2011; Valladares et al., 1994; Hosokawa et al., 2020)





Polar Cap Arcs

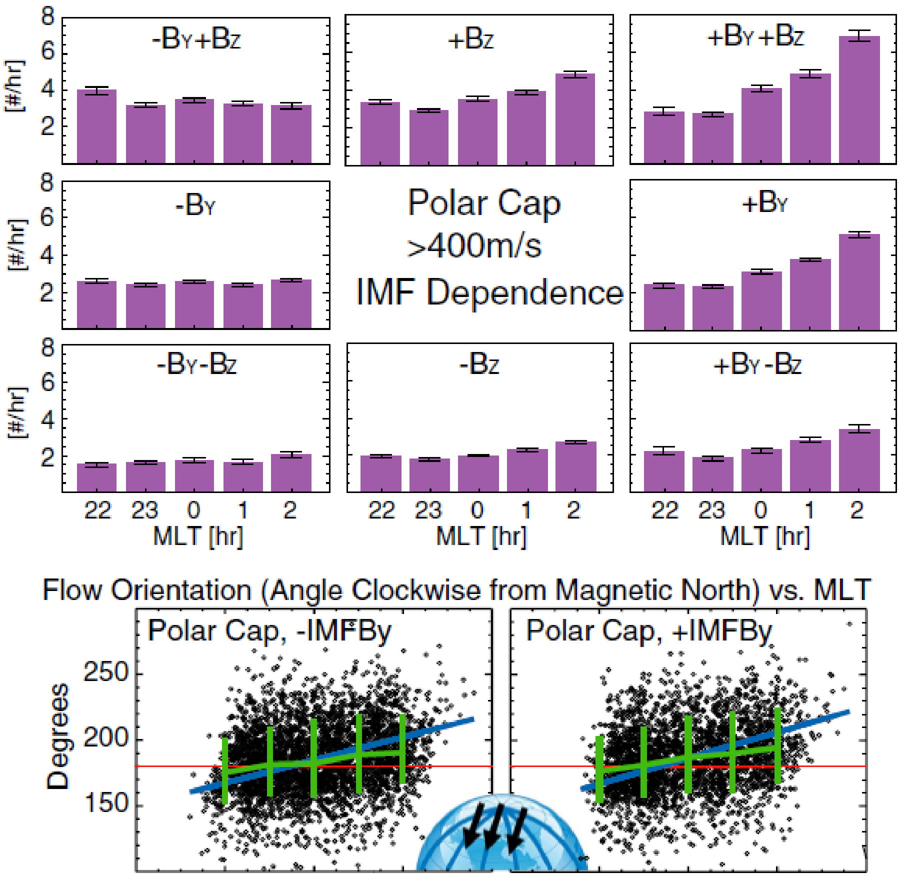
From Hosokawa et al., 2020: (**a**) Tail twist for duskward IMF (Cowley 1981) (**b**) oval-aligned TPA at dusk due to twist of tail plasma sheet for duskward IMF (Makita et al. 1991), (**c**) bifurcated tail plasma sheet mapping to a TPA in the center of the polar cap, and a conjugate TPA in the other hemisphere (Obara et al. 1988), (**d**) possible relations of polar cap arc formation with plasma sheet configuration (taken from Makita et al. 1991)

Convection on Mesoscales

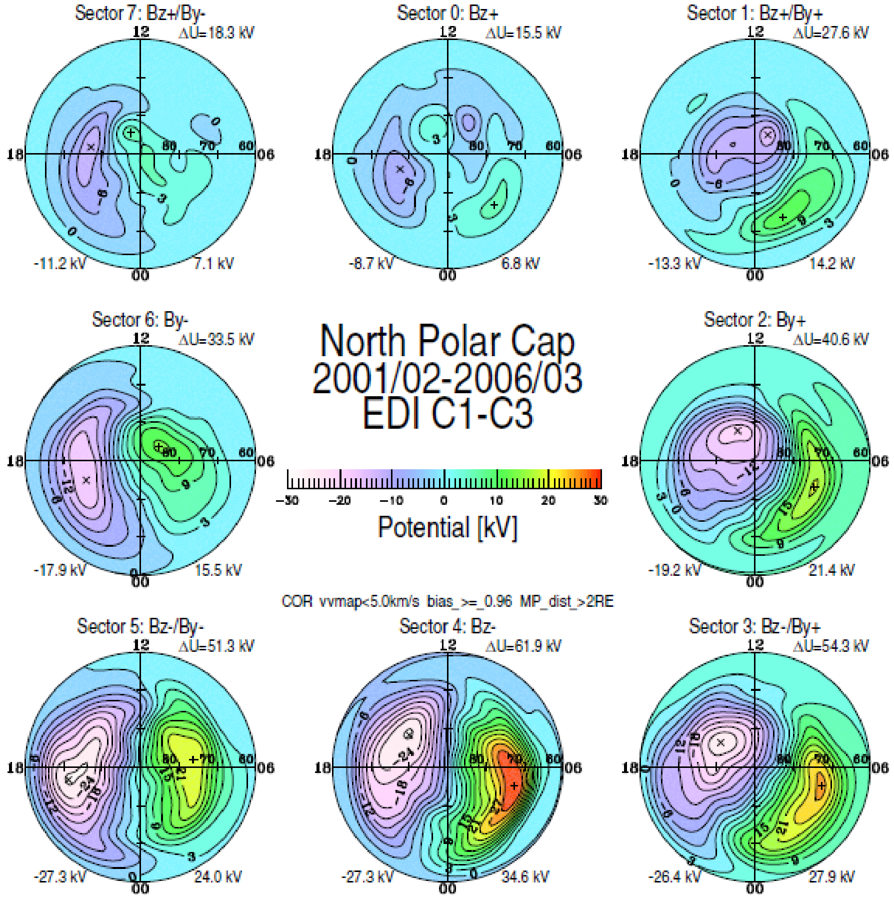
Polar Cap Convection

Mesoscale equatorward flows tend to be embedded in the background flow

- More post-midnight during +BY like polar cap arcs (Hosokawa et al., 2011) and majority of background convection.
- Tilted east-to-west



Gabrielse et al. (AGU JGR, 2018)

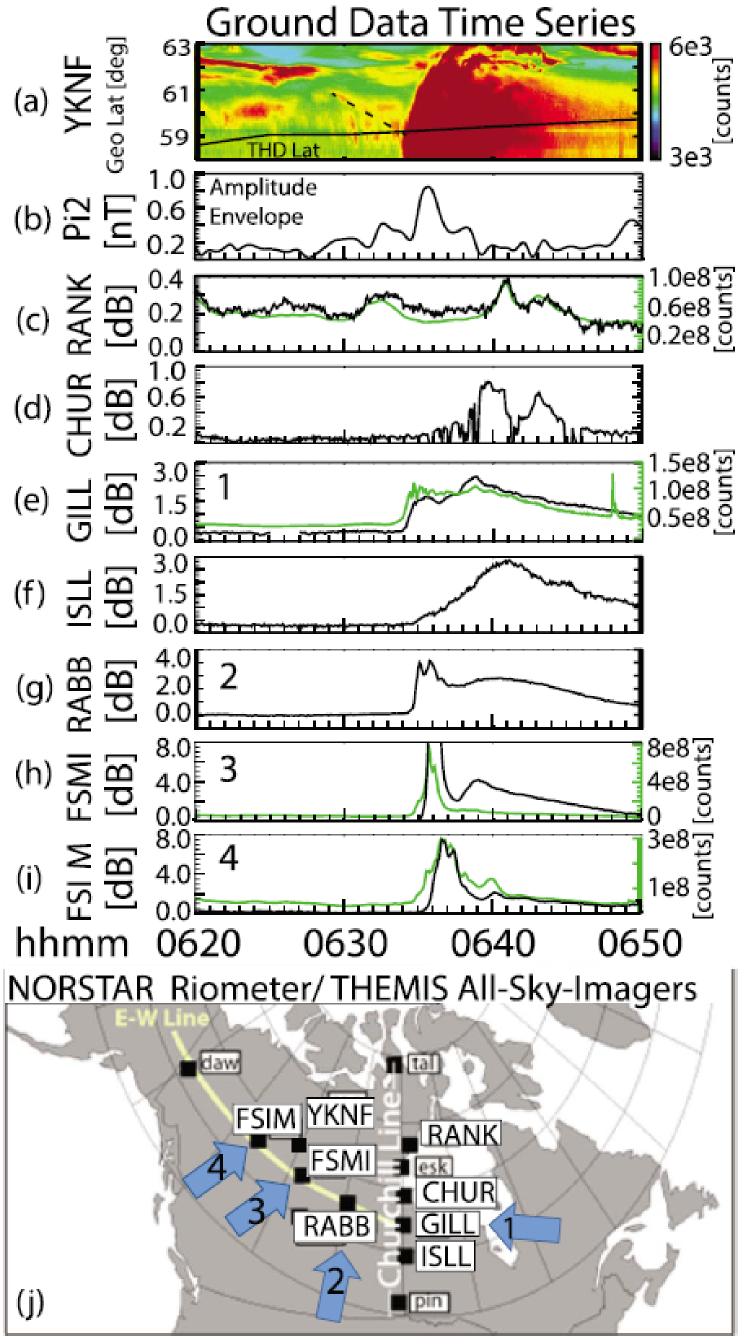


Haaland et al. (AGU JGR, 2007)

This presentation is being recorded. – christine.gabrielse@aero.org

Ground-based Observations

- Riometers: Observe injections via electron precipitation
- Magnetometers: Observe magnetic field variations due to currents



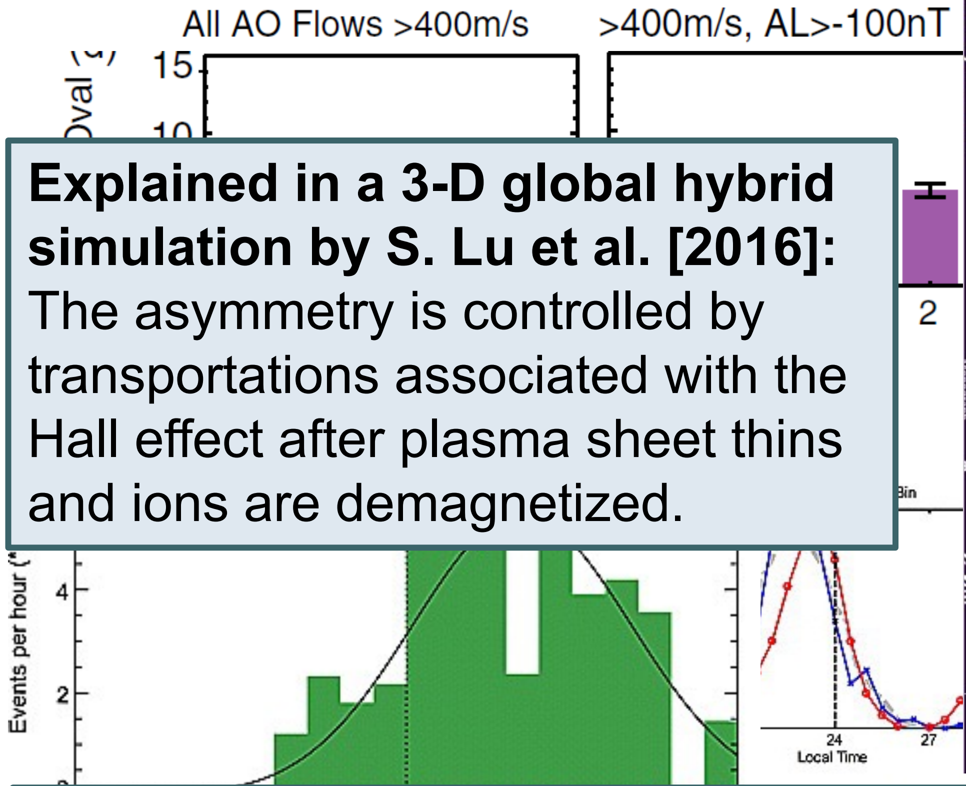
Gabrielse et al. (AGU JGR 2019)

Convection on Mesoscales

Magnetotail Flows and the Ionosphere

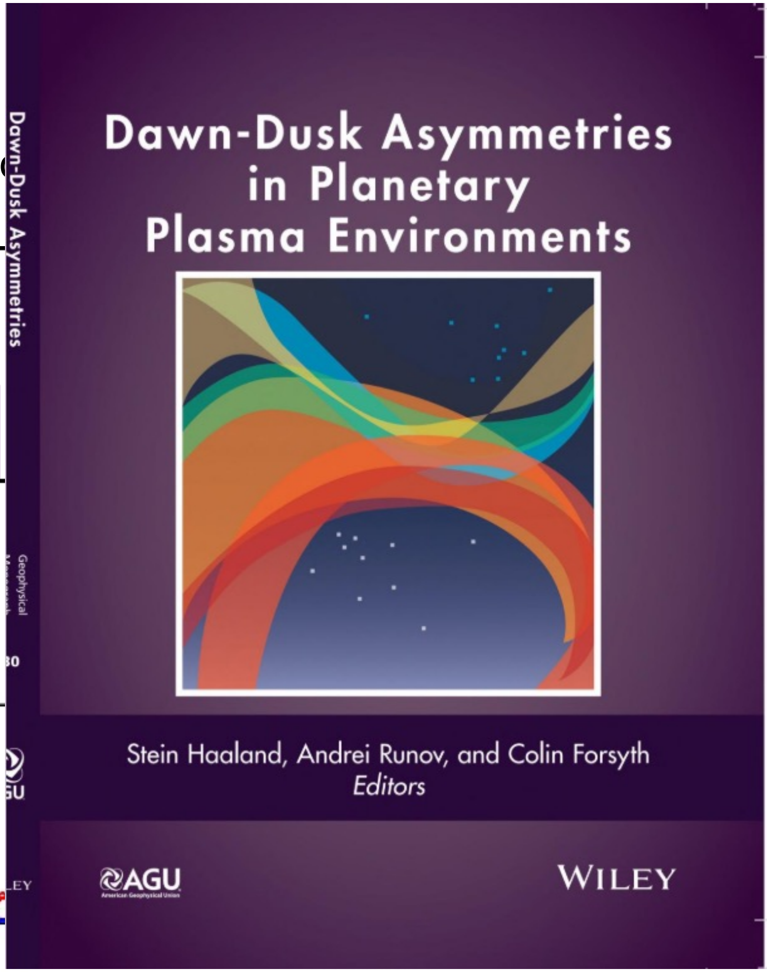
Pre-midnight preference for…

auroral oval (plasma sheet) mesoscale fl



Explained in a 3-D global hybrid simulation by S. Lu et al. [2016]: The asymmetry is controlled by transportations associated with the Hall effect after plasma sheet thins and ions are demagnetized.

Haaland et al. (AGU, 2017)



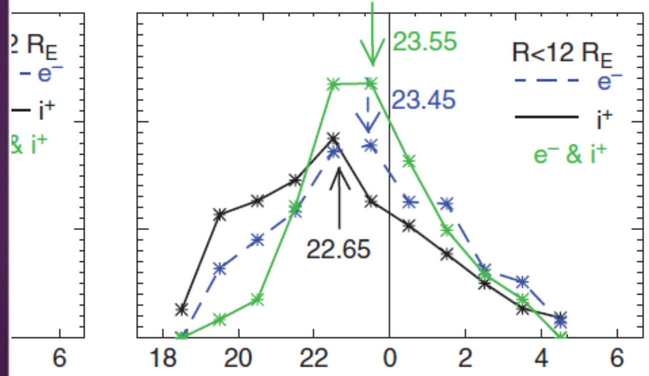
Dawn-Dusk Asymmetries in Planetary Plasma Environments

Stein Haaland, Andrei Runov, and Colin Forsyth
Editors

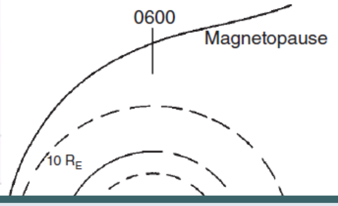
AGU

WILEY

(Mauk and McIlwain, 1974; Birn
; Gabrielse et al., 2014; 2017)

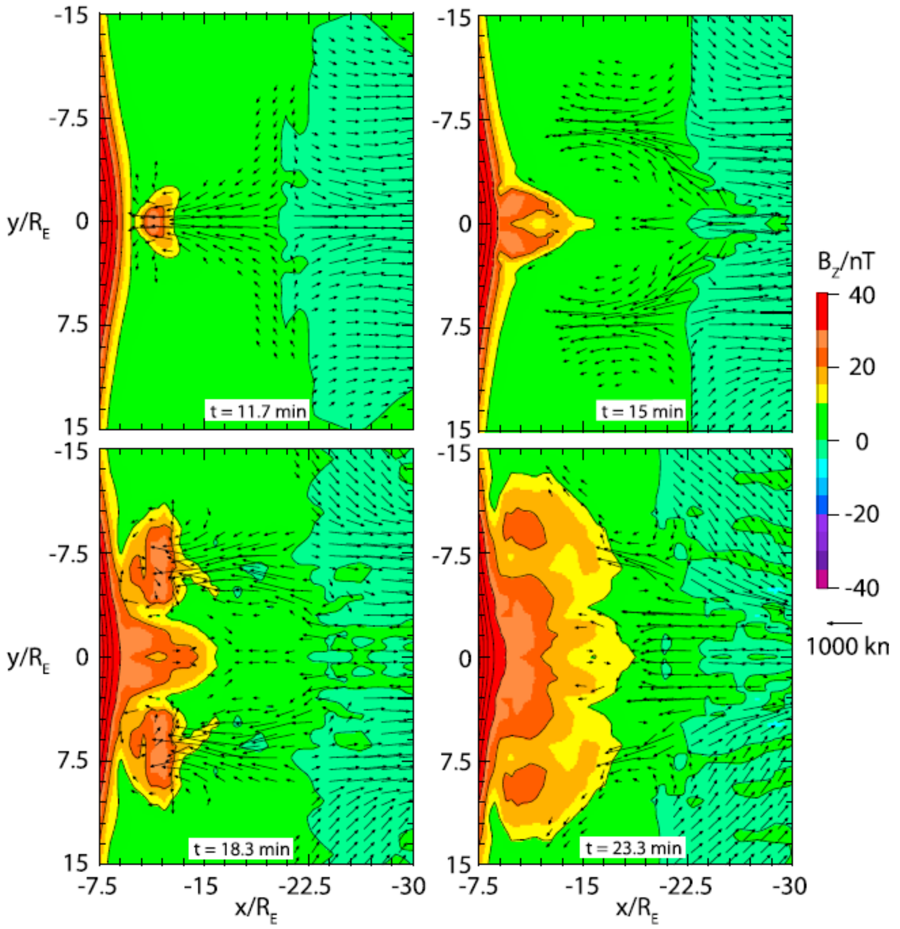


Gabrielse et al. (AGU JGR 2014)



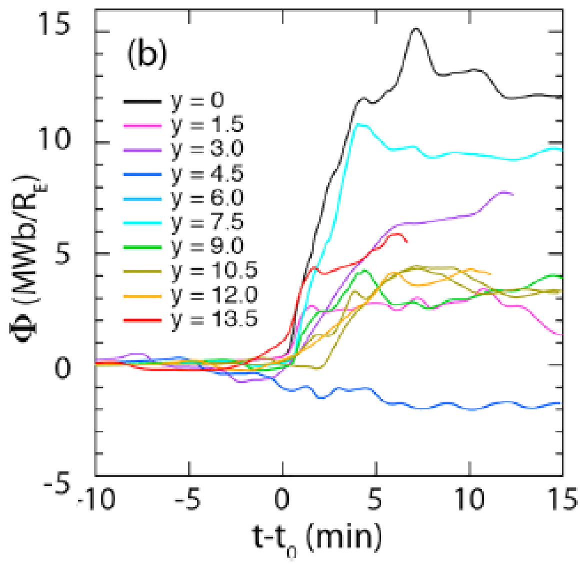
Dawn-dusk asymmetry: **reconnection** [Eastwood et al.,2010; Nagai et al. 2013; Genestreti et al., 2014], **fast flows** [McPherron et al., 2011; Lotko et al., 2014], **auroral oval fast flows** [Gallardo-Lacourt et al., 2014; Gabrielse et al., 2018], **TCRs** [Slavin et al., 2005; Imber et al., 2011], **injections** [Birn et al. 1997; Gabrielse et al. 2014], **energetic proton events at Vela (18 RE)** [Hones et al., 1976] and **IMP-8 (35 RE)** [Sarris et al., 1976]; **dipolarizing flux bundles** [Liu et al., 2013]; **dipolarizations** [Nosé et al., 2016]

Substorm: Magnetic Flux Transport



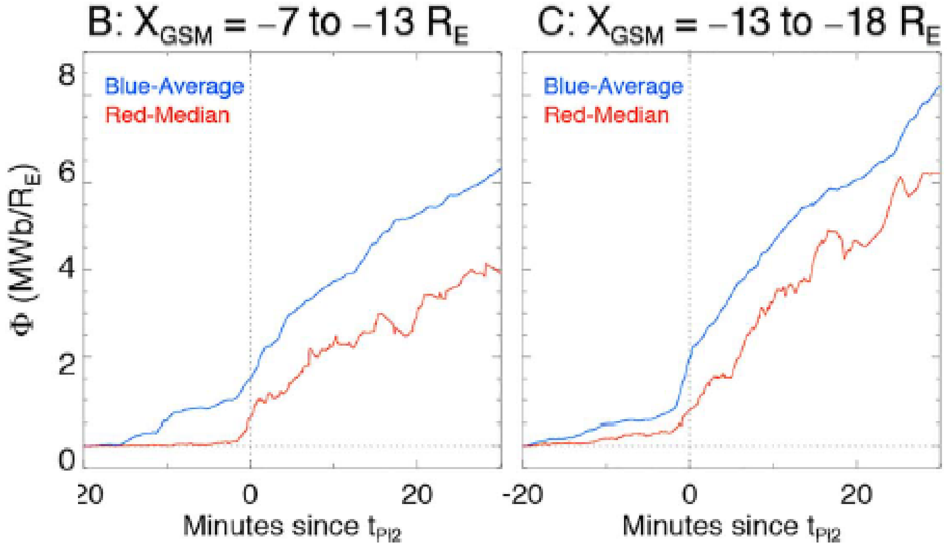
Figures modified from Birn et al. (AGU JGR 2019)

MHD



The total flux transported earthward from the reconnection site was ~ 2.3 × 10^8 Wb, commensurate with estimates of 1–3.6 ×10^8 Wb by Angelopoulos et al. (1994).

This flux was associated with up to seven dipolarization front events localized across the tail.

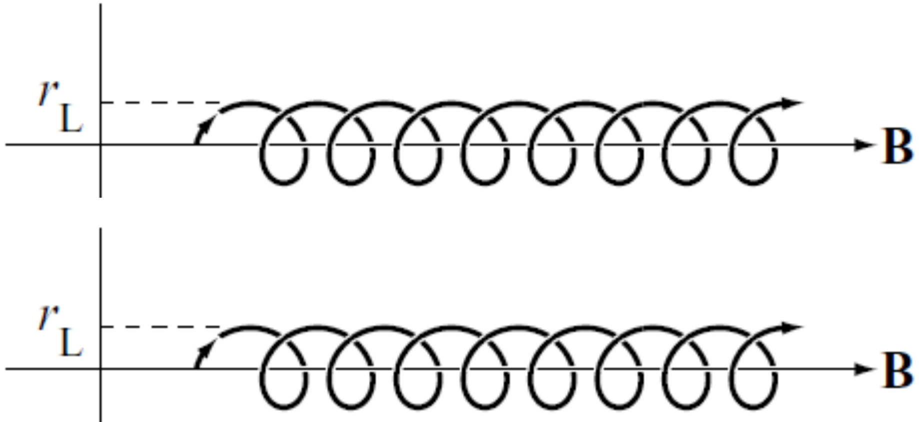


THEMIS statistics

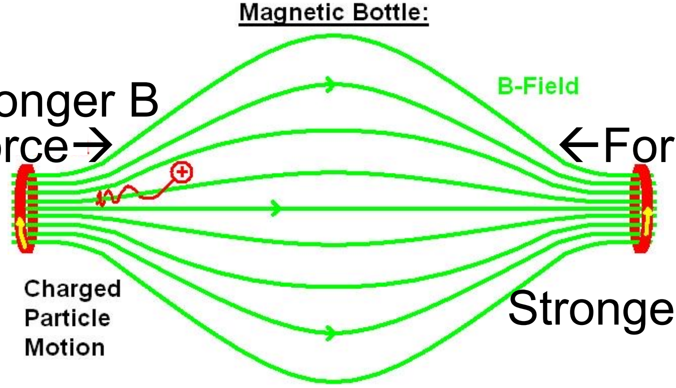
More Reading

- For recent papers, see the GEM Focus Group on Dipolarizations home page:
  - *https://gem.epss.ucla.edu/mediawiki/index.php/FG:_Magnetotail_Dipolarization_and_Its_Effects_on_the_Inner_Magnetosphere*
- For more specifics on dipolarization, see the GEM Plenary tutorial on the topic from 2017:
  - *https://gem.epss.ucla.edu/mediawiki/index.php/GEM_Tutorials#2017_Summer_Workshop*

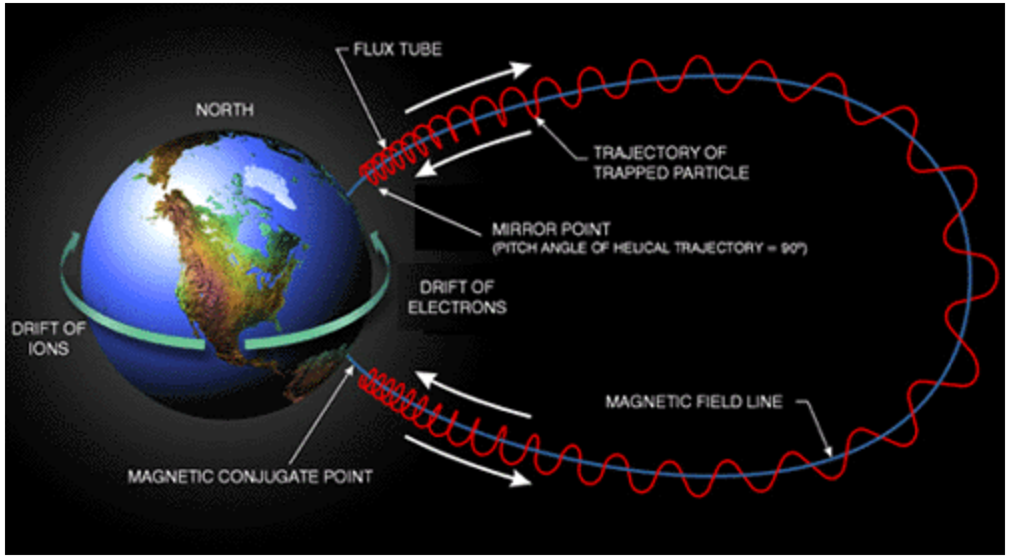
Bounce Motion



Pitch angle nonzero
Parallel velocity nonzero



Stronger B
Force→
←Force
Stronger B
Weaker B

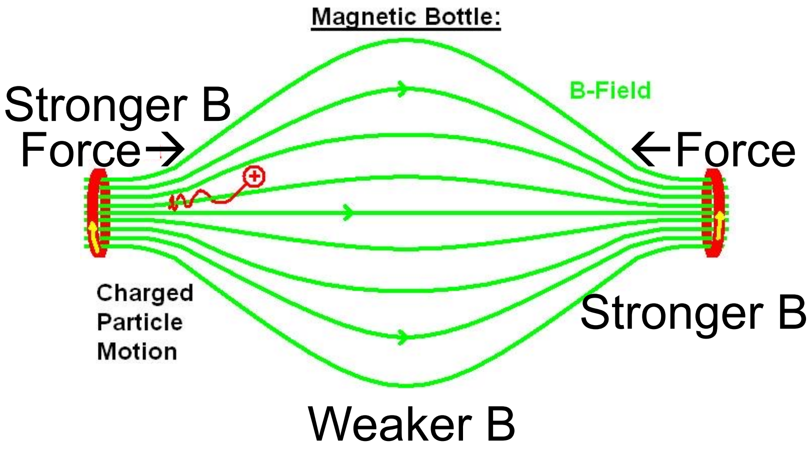


Fermi Acceleration

J = 2nd adiabatic invariant

Constant if field line length changes slower than bounce period

$$J = 2\int_{m1}^{m2} p_{\|} ds$$



Ping pong analogy